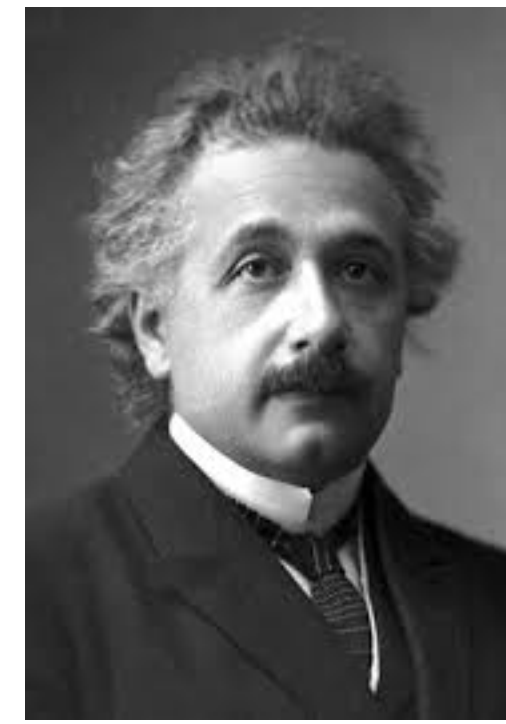
SUSY dark matter, axion, LHC and ILC

Howard Baer
University of Oklahoma

DSU2015, Kyoto

twin pillars of guidance:
naturalness & simplicity



``The appearance of fine-tuning in a scientific theory is like a cry of distress from nature, complaining that something needs to be better explained''

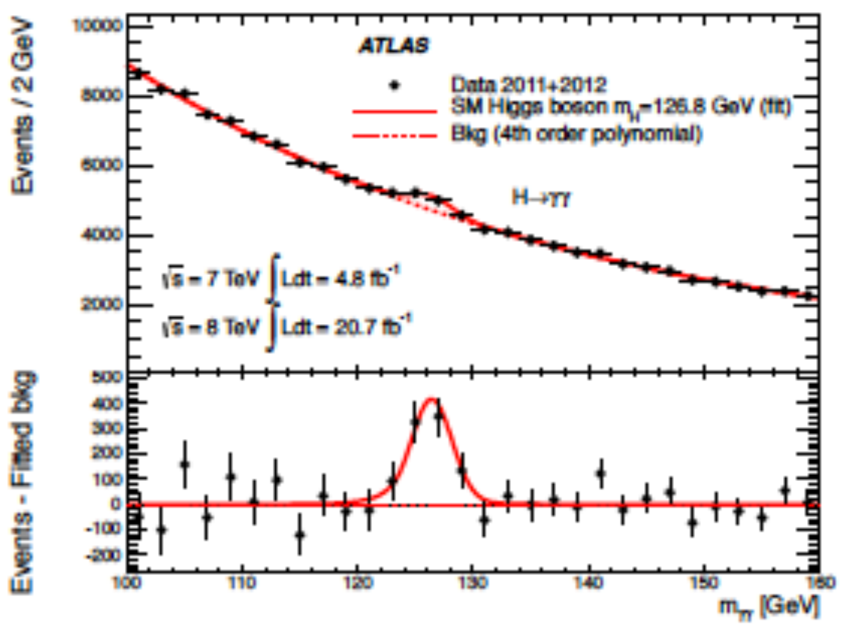
S. Weinberg

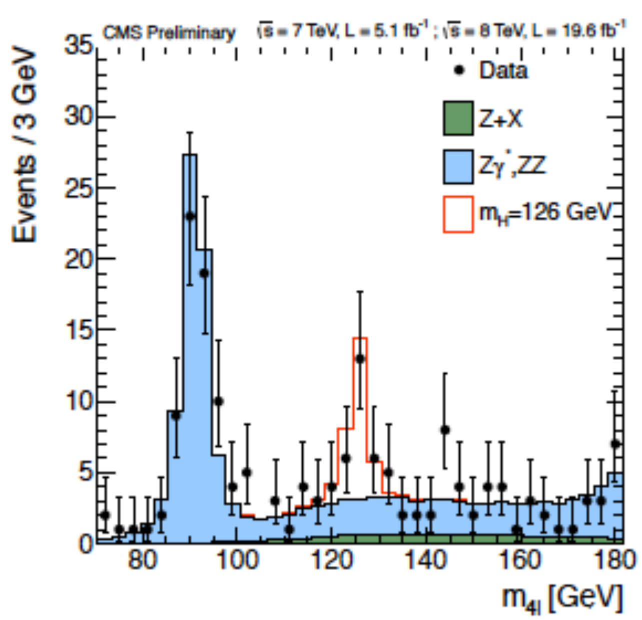
``Everything should be made as simple as possible, but not simpler''

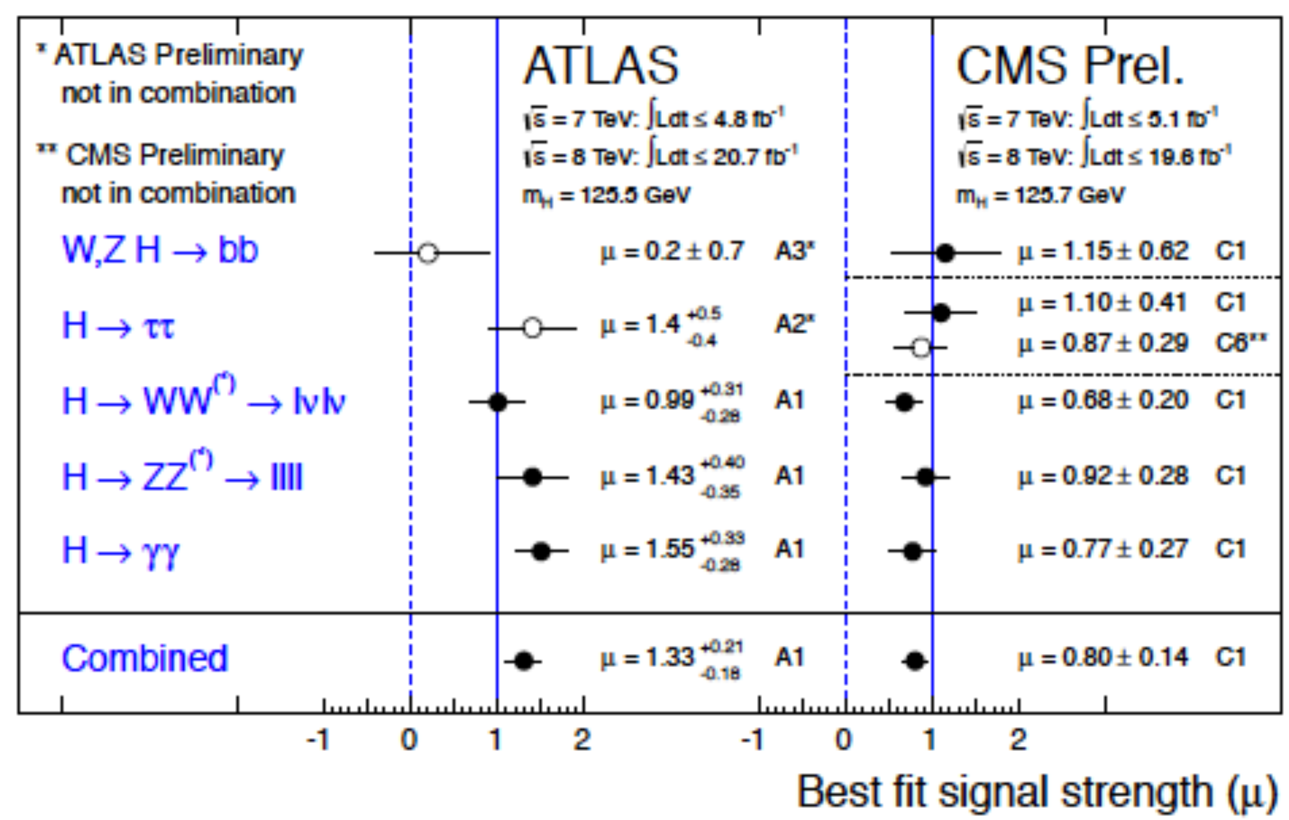
A. Einstein

Brief summary of SM:

We have found a very SM-like Higgs boson with m(h)=125 GeV at LHC but nothing yet beyond the SM

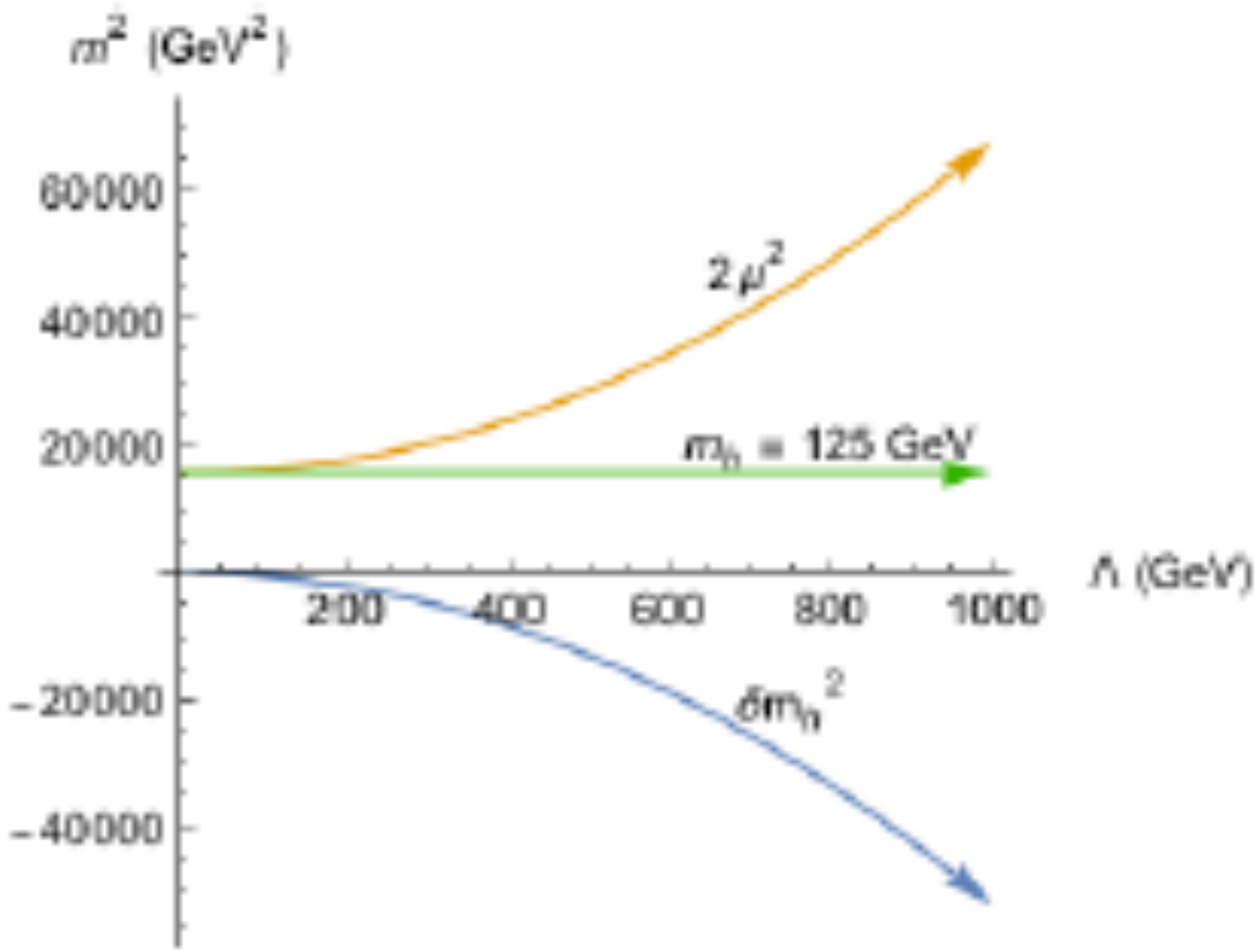






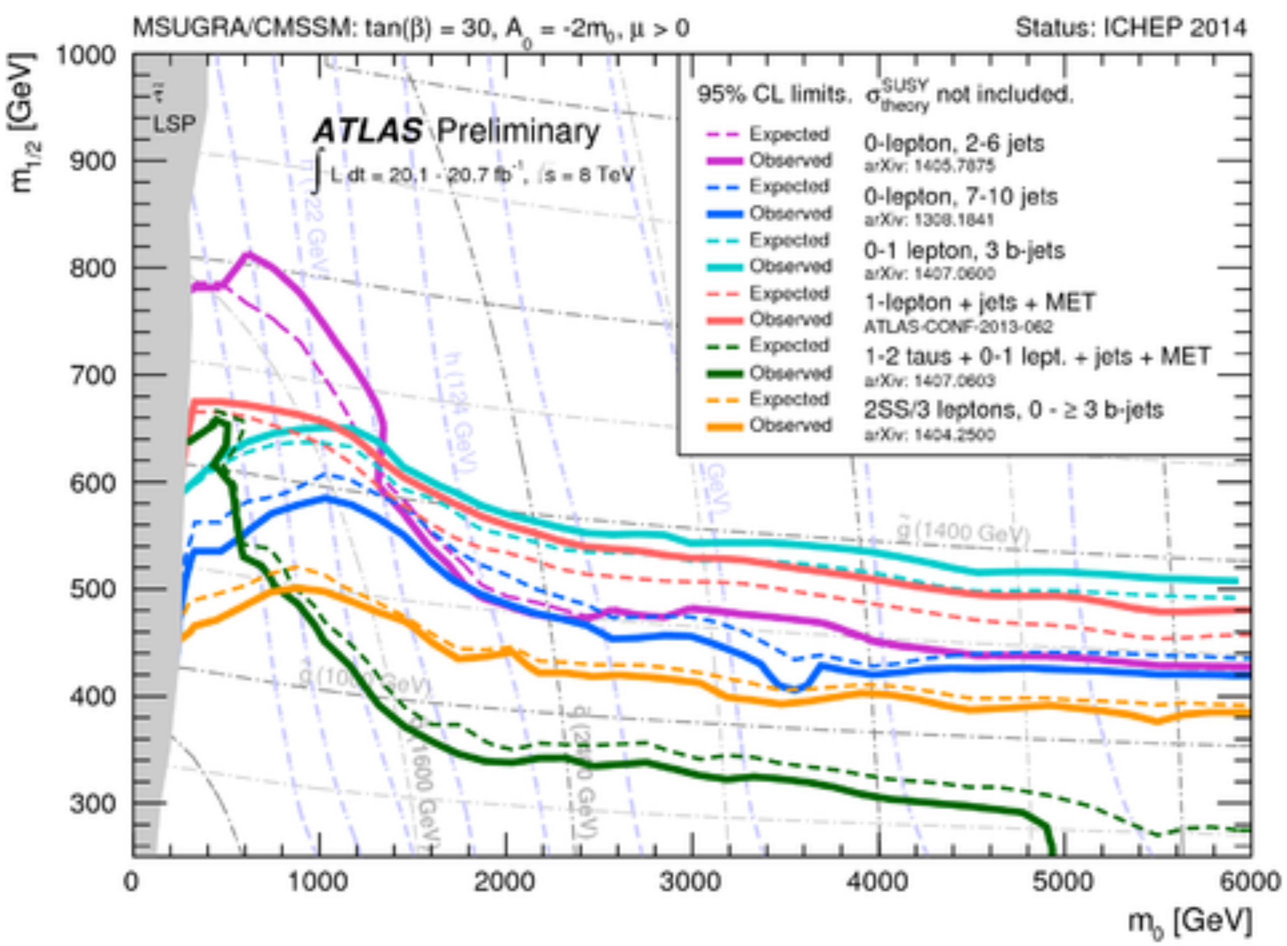
But the Higgs mass problem remains

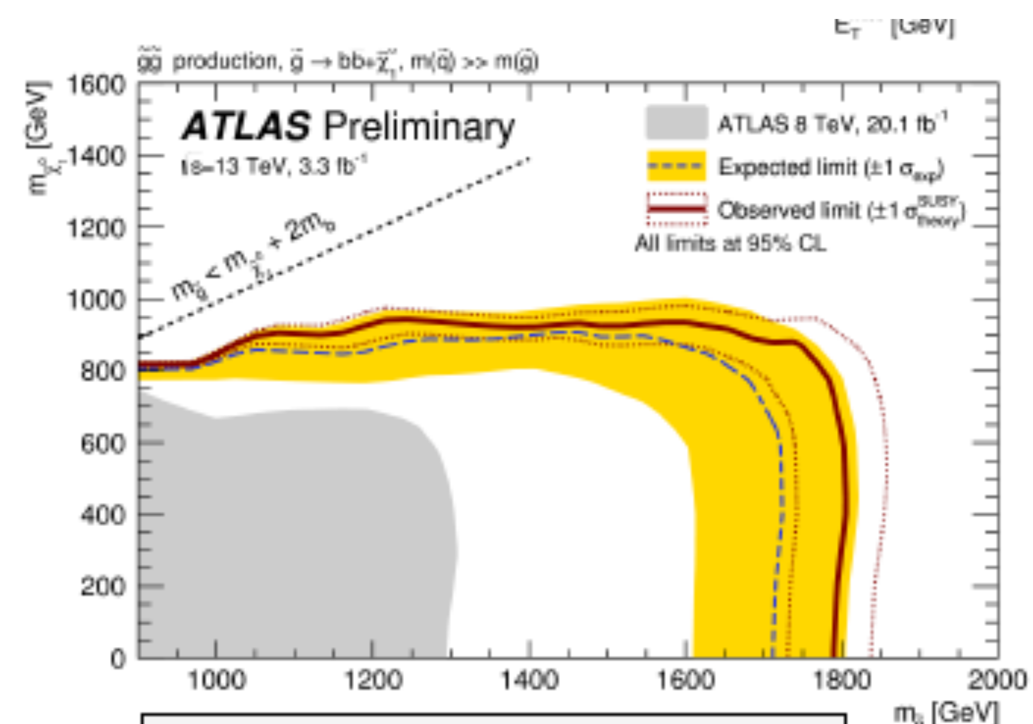
Wilson 1972
Weinberg
Gildener



Independent radiative corrections to m(h) should be smaller than m(h): SM only valid below ~1 TeV

SUSY tames scalar mass problem in elegant fashion: but where are sparticles?





Gluino exclusion reach 1.8 TeV

$m_{\tilde{g}} > 1.8 \text{ TeV}$

$m_{\tilde{g}} > 1.3\ TeV$ for $m_{\tilde{q}} \gg m_{\tilde{g}}$

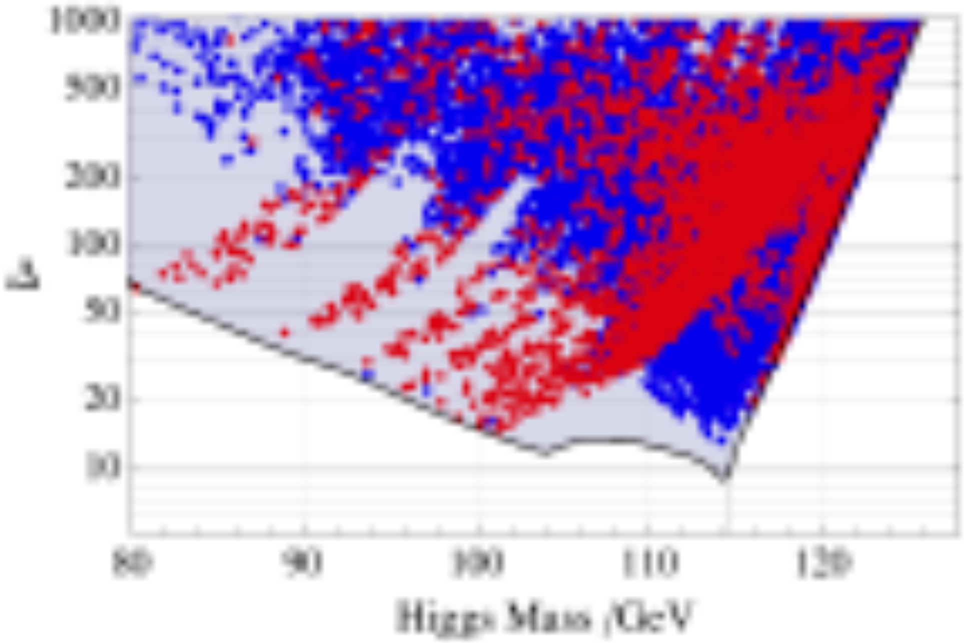
$m_{\tilde{g}} > 1.8\ TeV$ for $m_{\tilde{q}} \sim m_{\tilde{g}}$

$m_{\tilde{t}_1} \sim \text{multi} - \text{TeV}$ for $m_h \simeq 125\ GeV$

These bounds appear in sharp conflict with EW ``naturalness''

	mass
gluino	400 GeV
uR	400 GeV
eR	350 GeV
chargino	100 GeV
neutralino	50 GeV

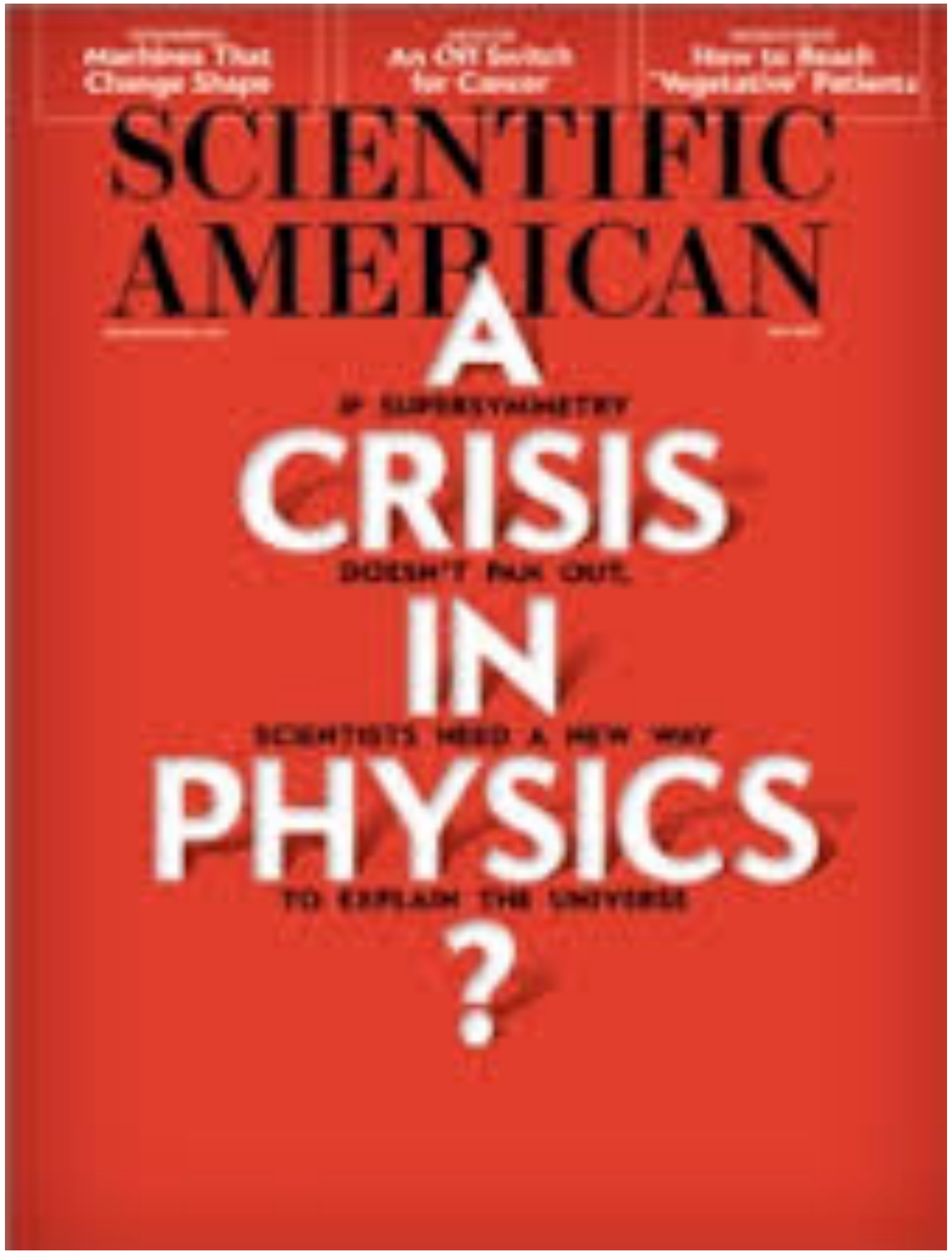
Barbieri-Giudice 10% bounds, 1987



Cassel, Ghilencea, Ross,2009

$\Delta \to 1000$
as $m_h \to 125$ GeV
0.1% tuning!?

Is there a crisis in physics?



short answer:
No! but there may be a crisis
in how theorists
calculate naturalness

This unshakable fidelity to supersymmetry is widely shared. Particle theorists do admit, however, that the idea of natural supersymmetry is already in trouble and is headed for the dustbin of history unless superpartners are discovered soon...

Lykken, Spiropulu,2014

``…settling the ultimate fate of naturalness is perhaps the most profound theoretical question of our time''

Arkani-Hamed et al., arXiv:1511.06495

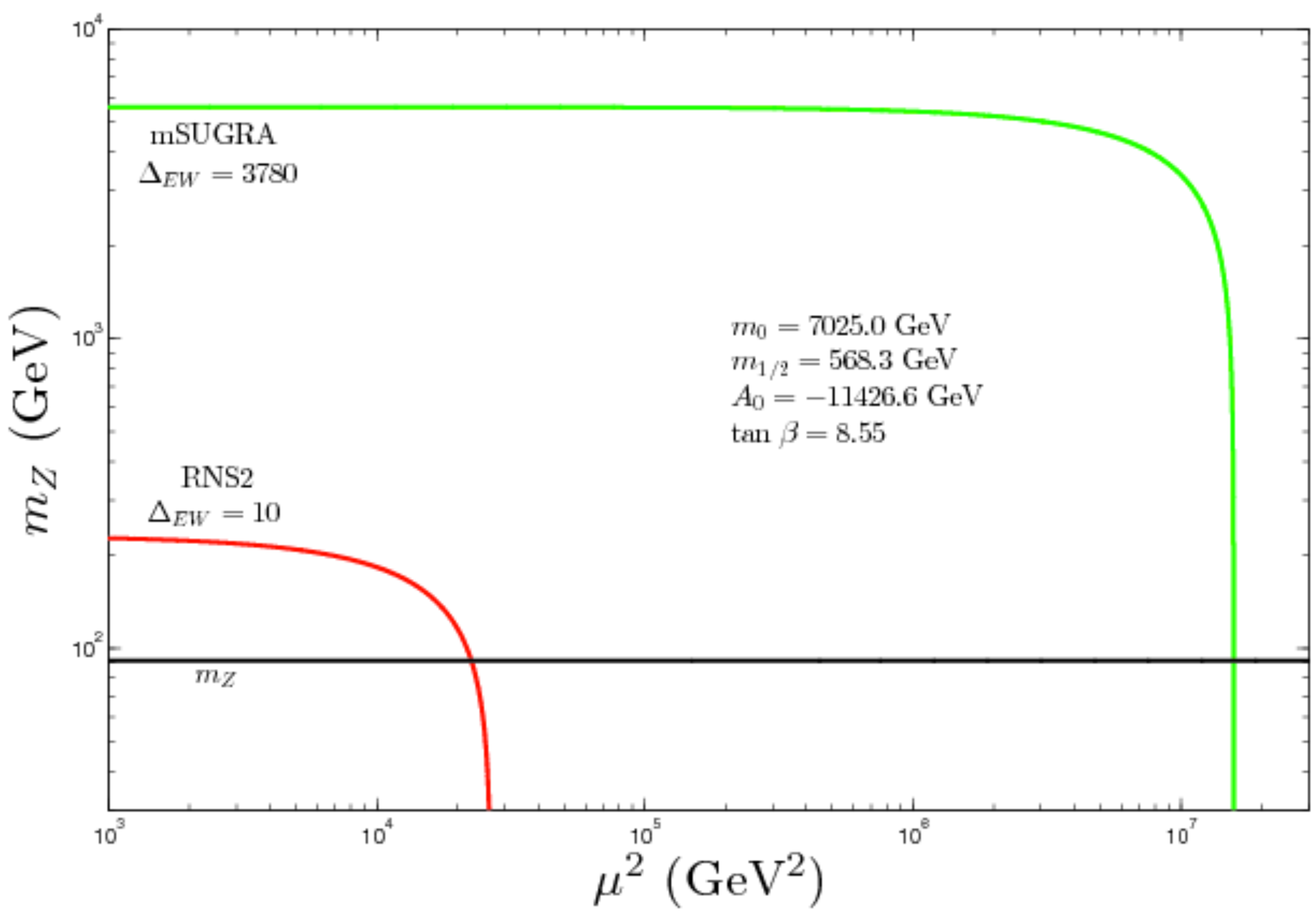
``Given the magnitude of the stakes involved, it is vital to get a clear verdict on naturalness from experiment''

This should be matched by theoretical scrutiny of what we mean by naturalness

Three measures of fine-tuning:

First: simple electroweak fine-tuning in SUSY:
dial value of mu so that Z mass comes out right:
everybody does it but it is hidden inside spectra codes (Isajet, SuSpect, SoftSUSY, Spheno, SSARD)

$$\frac{m_Z^2}{2} = \frac{m_{H_d}^2 + \Sigma_d^d - (m_{H_u}^2 + \Sigma_u^u)\tan^2\beta}{\tan^2\beta - 1} - \mu^2 \simeq -m_{H_u}^2 - \Sigma_u^u - \mu^2$$



e.g. in CMSSM/ mSUGRA: one then concludes nature gives this:

#1: Most conservative, simple measure: Δ_{EW}

Working only at the weak scale, minimize scalar potential: calculate m(Z) or m(h)

No large uncorrelated cancellations in m(Z) or m(h)

$$\frac{m_Z^2}{2} = \frac{m_{H_d}^2 + \Sigma_d^d - (m_{H_u}^2 + \Sigma_u^u)\tan^2\beta}{\tan^2\beta - 1} - \mu^2 \sim -m_{H_u}^2 - \Sigma_u^u - \mu^2$$

$\Delta_{EW} \equiv max_i |C_i| / (m_Z^2/2)$ with $C_{H_u} = -m_{H_u}^2 \tan^2\beta/(\tan^2\beta - 1)$ etc.

simple, direct, unambiguous interpretation:

- $|\mu| \sim m_Z \sim 100 - 200$ GeV
- $m_{H_u}^2$ should be driven to small negative values such that $-m_{H_u}^2 \sim 100 - 200$ GeV at the weak scale and
- that the radiative corrections are not too large: $\Sigma_u^u \lesssim 100 - 200$ GeV

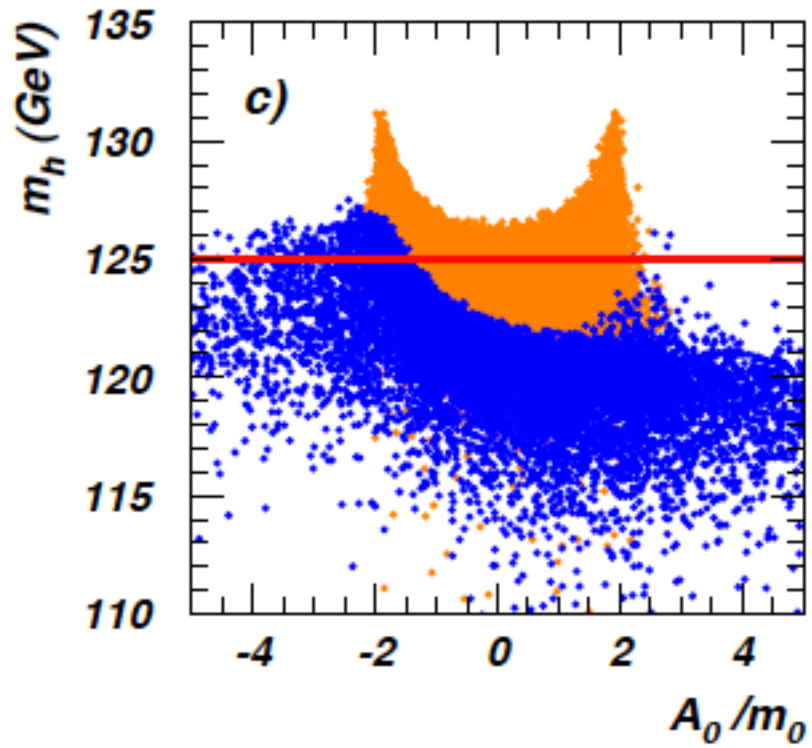
CETUP*-12/002, FTPI-MINN-12/22, UMN-TH-3109/12, UH-511-1195-12

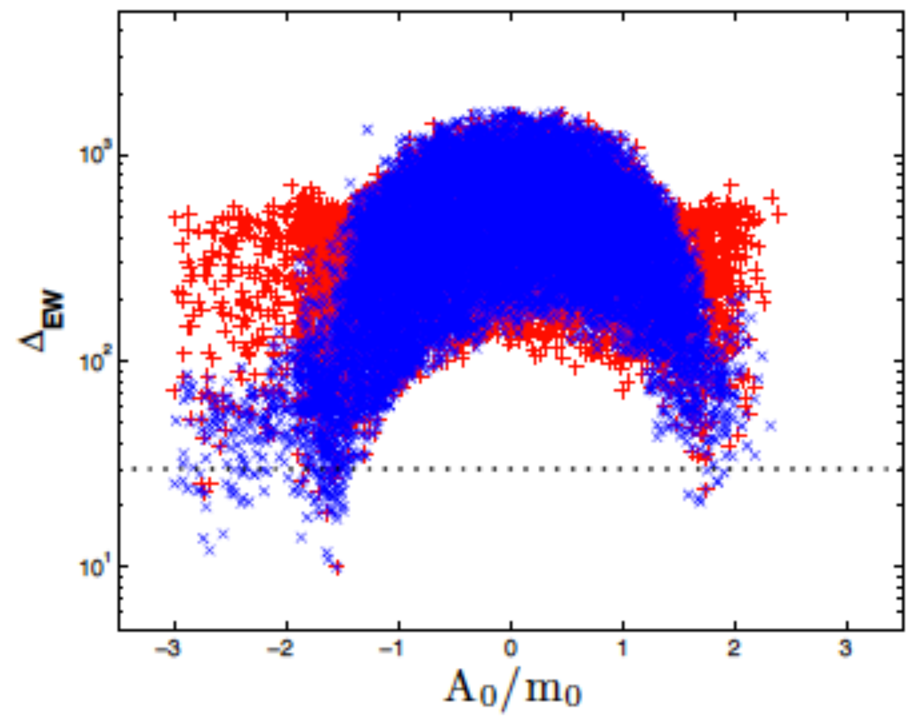
Radiative natural SUSY with a 125 GeV Higgs boson

Howard Baer,[1] Vernon Barger, Peisi Huang,[2] Azar Mustafayev,[3] and Xerxes Tata[4]
[1]Dept. of Physics and Astronomy, University of Oklahoma, Norman, OK, 73019, USA
[2]Dept. of Physics, University of Wisconsin, Madison, WI 53706, USA
[3]W. I. Fine Institute for Theoretical Physics, University of Minnesota, Minneapolis, MN 55455, USA

PRL109 (2012) 161802

Large value of A_t reduces $\Sigma_u^u(\tilde{t}_{1,2})$ contributions to Δ_{EW} while uplifting m_h to ~ 125 GeV





$$\Sigma_u^u(\tilde{t}_{1,2}) = \frac{3}{16\pi^2} F(m_{\tilde{t}_{1,2}}^2) \left[f_t^2 - g_Z^2 \mp \frac{f_t^2 A_t^2 - 8 g_Z^2 (\frac{1}{4} - \frac{2}{3} x_W) \Delta_t}{m_{\tilde{t}_2}^2 - m_{\tilde{t}_1}^2} \right]$$

$$\Delta_t = (m_{\tilde{t}_L}^2 - m_{\tilde{t}_R}^2)/2 + M_Z^2 \cos 2\beta (\frac{1}{4} - \frac{2}{3} x_W)$$

$$F(m^2) = m^2 \left(\log \frac{m^2}{Q^2} - 1 \right) \qquad Q^2 = m_{\tilde{t}_1} m_{\tilde{t}_2}$$

#2: Higgs mass or large-log fine-tuning Δ_{HS}

It is tempting to pick out one-by-one
quantum fluctuations but must combine log divergences
before taking any limit

$$m_h^2 \simeq \mu^2 + m_{H_u}^2 + \delta m_{H_u}^2|_{rad}$$

$$\frac{dm_{H_u}^2}{dt} = \frac{1}{8\pi^2}\left(-\frac{3}{5}g_1^2 M_1^2 - 3g_2^2 M_2^2 + \frac{3}{10}g_1^2 S + 3f_t^2 X_t\right) \qquad X_t = m_{Q_3}^2 + m_{U_3}^2 + m_{H_u}^2 + A_t^2$$

neglect gauge pieces, S, mHu and running;
then we can integrate from m(SUSY) to Lambda

$$\delta m_{H_u}^2 \sim -\frac{3f_t^2}{8\pi^2}\left(m_{Q_3}^2 + m_{U_3}^2 + A_t^2\right)\ln(\Lambda/m_{SUSY})$$

$$\Delta_{HS} \sim \delta m_h^2/(m_h^2/2) < 10$$

$$m_{\tilde{t}_{1,2},\tilde{b}_1} < 500\ \text{GeV}$$

$$m_{\tilde{g}} < 1.5\ \text{TeV}$$

old natural SUSY

then A_t can't be too big

Most claims against SUSY stem from overestimates of EW fine-tuning.

These arise from violations of the

Prime directive on fine-tuning:

``Thou shalt not claim fine-tuning of dependent quantities one against another!”

HB, Barger, Mickelson, Padeffke-Kirkland, arXiv:1404.2277

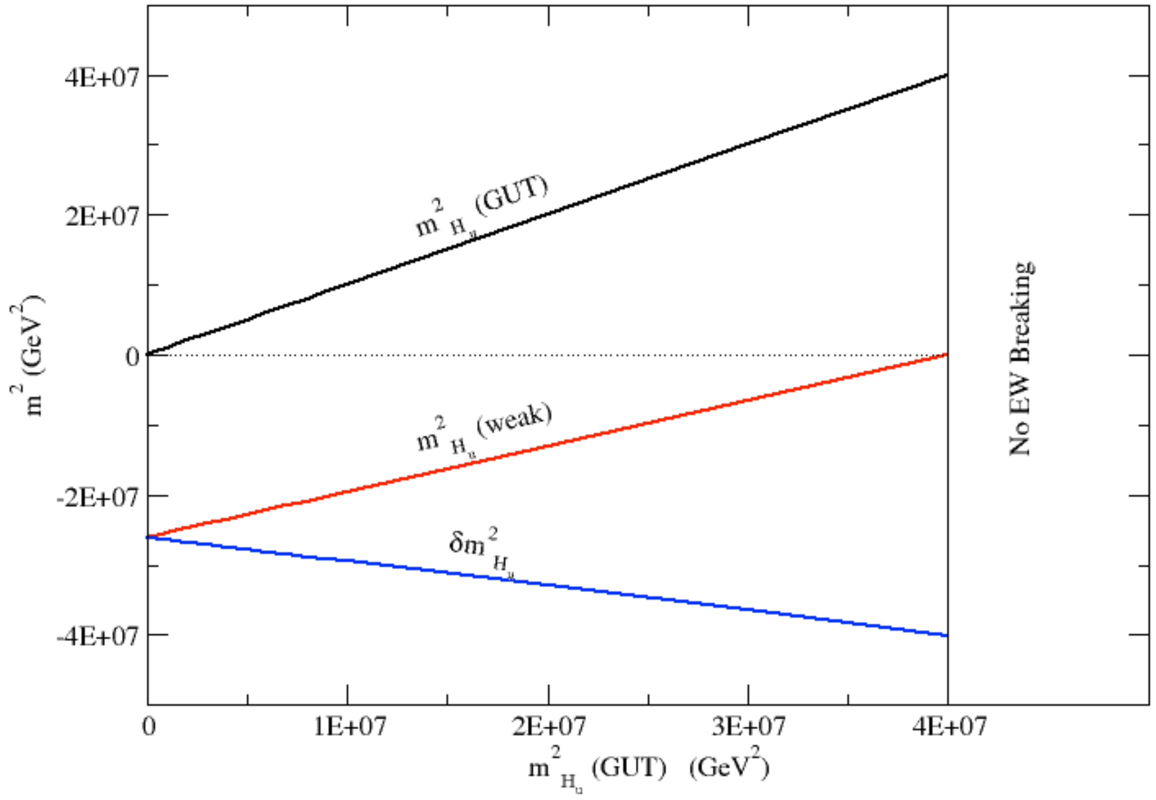
Is $\mathcal{O} = \mathcal{O} + b - b$ fine-tuned for $b > \mathcal{O}$?

What's wrong with this argument?

In zeal for simplicity, have made several simplifications: most egregious is that one sets m(Hu)^2=0 at beginning to simplify

$m^2_{H_u}(\Lambda)$ and $\delta m^2_{H_u}$ are *not* independent!

violates prime directive!



The larger $m^2_{H_u}(\Lambda)$ becomes, then the larger becomes the cancelling correction!

HB, Barger, Savoy

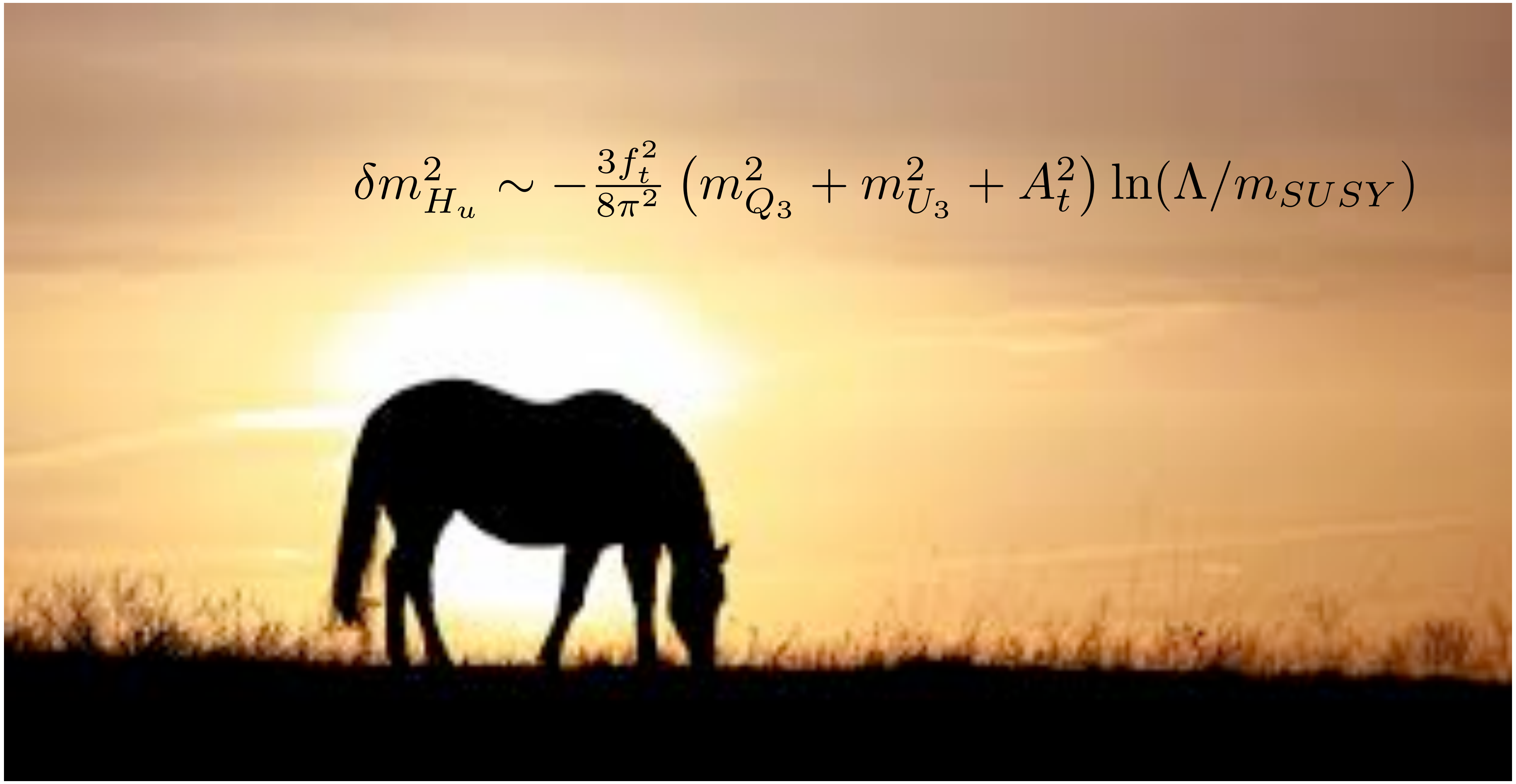
To fix: combine dependent terms:

$m_h^2 \simeq \mu^2 + \left(m_{H_u}^2(\Lambda) + \delta m_{H_u}^2\right)$ where now both μ^2 and $\left(m_{H_u}^2(\Lambda) + \delta m_{H_u}^2\right)$ are $\sim m_Z^2$

After re-grouping: $\boxed{\Delta_{HS} \simeq \Delta_{EW}}$

Instead of: the radiative correction $\delta m_{H_u}^2 \sim m_Z^2$
we now have: the radiatively-corrected $m_{H_u}^2 \sim m_Z^2$

Recommendation: put this horse out to pasture



R.I.P.

sub-TeV 3rd generation squarks not required for naturalness

#3: EENZ/BG traditional measure Δ_{BG}

Such a re-grouping is properly used in the EENZ/BG measure:

$$\Delta_{BG} \equiv max_i \left[c_i\right], \quad \text{where} \quad c_i = \left|\frac{\partial \ln m_Z^2}{\partial \ln p_i}\right| = \left|\frac{p_i}{m_Z^2}\frac{\partial m_Z^2}{\partial p_i}\right|$$

the p_i constitute the fundamental parameters of the model.

for pMSSM, obviously $\Delta_{BG} \simeq \Delta_{EW}$

What about models defined at high scale?

$$\frac{m_Z^2}{2} = \frac{m_{H_d}^2 - m_{H_u}^2 \tan^2\beta}{\tan^2\beta - 1} - \mu^2 \simeq -m_{H_u}^2 - \mu^2$$

↑

express weak scale value in terms of high scale parameters

Express m(Z) in terms of GUT scale parameters:

$$m_Z^2 \simeq -2m_{H_u}^2 - 2\mu^2$$ (weak scale relation)

$$-2\mu^2(m_{SUSY}) = -2.18\mu^2$$

$$\begin{aligned}-2m_{H_u}^2(m_{SUSY}) = & \; 3.84M_3^2 + 0.32M_3M_2 + 0.047M_1M_3 - 0.42M_2^2 \\ & +0.011M_2M_1 - 0.012M_1^2 - 0.65M_3A_t - 0.15M_2A_t \\ & -0.025M_1A_t + 0.22A_t^2 + 0.004m_3A_b \\ & -1.27m_{H_u}^2 - 0.053m_{H_d}^2 \\ & +0.73m_{Q_3}^2 + 0.57m_{U_3}^2 + 0.049m_{D_3}^2 - 0.052m_{L_3}^2 + 0.053m_{E_3}^2 \\ & +0.051m_{Q_2}^2 - 0.11m_{U_2}^2 + 0.051m_{D_2}^2 - 0.052m_{L_2}^2 + 0.053m_{E_2}^2 \\ & +0.051m_{Q_1}^2 - 0.11m_{U_1}^2 + 0.051m_{D_1}^2 - 0.052m_{L_1}^2 + 0.053m_{E_1}^2,\end{aligned}$$

all GUT scale parameters

Ibanez, Lopez, Munoz; Lleyda, Munoz

Kane, King

Abe, Kobayashi, Omura; S. P. Martin

For generic parameter choices, Δ_{BG} *is large*

But if: $m_{Q_{1,2}} = m_{U_{1,2}} = m_{D_{1,2}} = m_{L_{1,2}} = m_{E_{1,2}} \equiv m_{16}(1,2)$ then $\sim 0.007m_{16}^2(1,2)$

Even better: $m_{H_u}^2 = m_{H_d}^2 = m_{16}^2(3) \equiv m_0^2$ => $-0.017m_0^2$

For correlated parameters, EWFT collapses in 3rd gen. sector!

Feng, Matchev, Moroi

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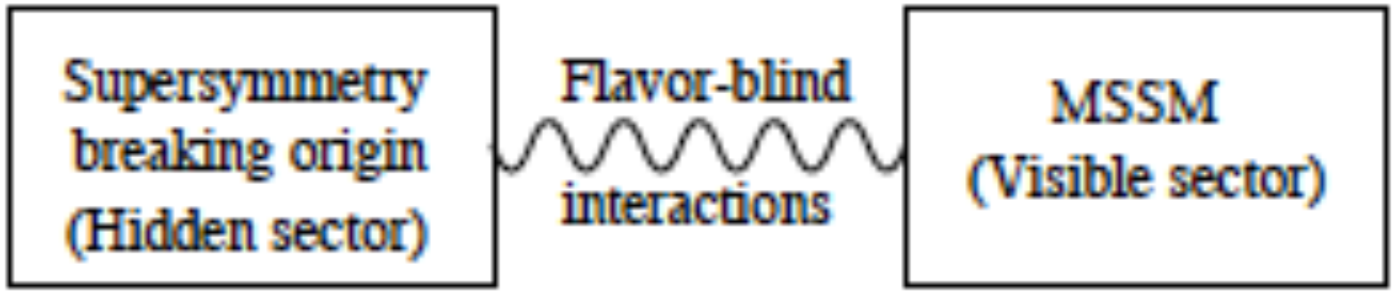
Feng, Matchev, Moroi

- Usually Δ_{BG} is applied to *multi-parameter effective theories* where multiple soft terms are adopted as parameter set.
- For these theories, the multiple soft terms parametrize our ignorance of details of the hidden sector SUSY breaking.
- But in supergravity, for any given hidden sector, soft terms are all *dependent* and can be computed as multiples of $m_{3/2}$.

Thus, the usual evaluation of Δ_{BG} also violates the prime directive!

To properly apply BG measure, need to identify independent soft breaking terms



examine gravity mediation

For any particular SUSY breaking hidden sector, each soft term is some multiple of gravitino mass m(3/2)

$$\begin{aligned} m_{H_u}^2 &= a_{H_u} \cdot m_{3/2}^2, \\ m_{Q_3}^2 &= a_{Q_3} \cdot m_{3/2}^2, \\ A_t &= a_{A_t} \cdot m_{3/2}, \\ M_i &= a_i \cdot m_{3/2}, \\ & \cdots . \end{aligned}$$

Soni, Weldon (1983);
Kaplunovsky, Louis (1992);
Brignole, Ibanez, Munoz (1993)

Since we don't know hidden sector, we impose parameters which parameterize our ignorance:
but this doesn't mean each parameter is independent

e.g. dilaton-dominated SUSY breaking: $m_0^2 = m_{3/2}^2$ with $m_{1/2} = -A_0 = \sqrt{3} m_{3/2}$

Writing each soft term as a multiple of m(3/2) then we allow for correlations/cancellations:

GUT scale param's

$$m_Z^2 = -2.18\mu^2 + a \cdot m_{3/2}^2$$

numerical co-efficient which depends on hidden sector

for naturalness, then

$\mu^2 \sim m_Z^2$ and $a \cdot m_{3/2}^2 \sim m_Z^2$

either $m_{3/2} \sim m_Z$ *or* a is small

$$m_Z^2 \simeq -2\mu^2(weak) - 2m_{H_u}^2(weak) \simeq -2.18\mu^2(GUT) + a \cdot m_{3/2}^2$$

then $-m_{H_u}^2(weak) \sim a \cdot m_{3/2}^2 \sim m_Z^2$

$$\lim_{n_{SSB}\to 1} \Delta_{BG} \to \Delta_{EW}$$

Thus, correctly applying these measures by first collecting dependent quantities, we find that- at tree level- all agree:

$$\Delta_{HS} \simeq \Delta_{BG} \simeq \Delta_{EW}$$

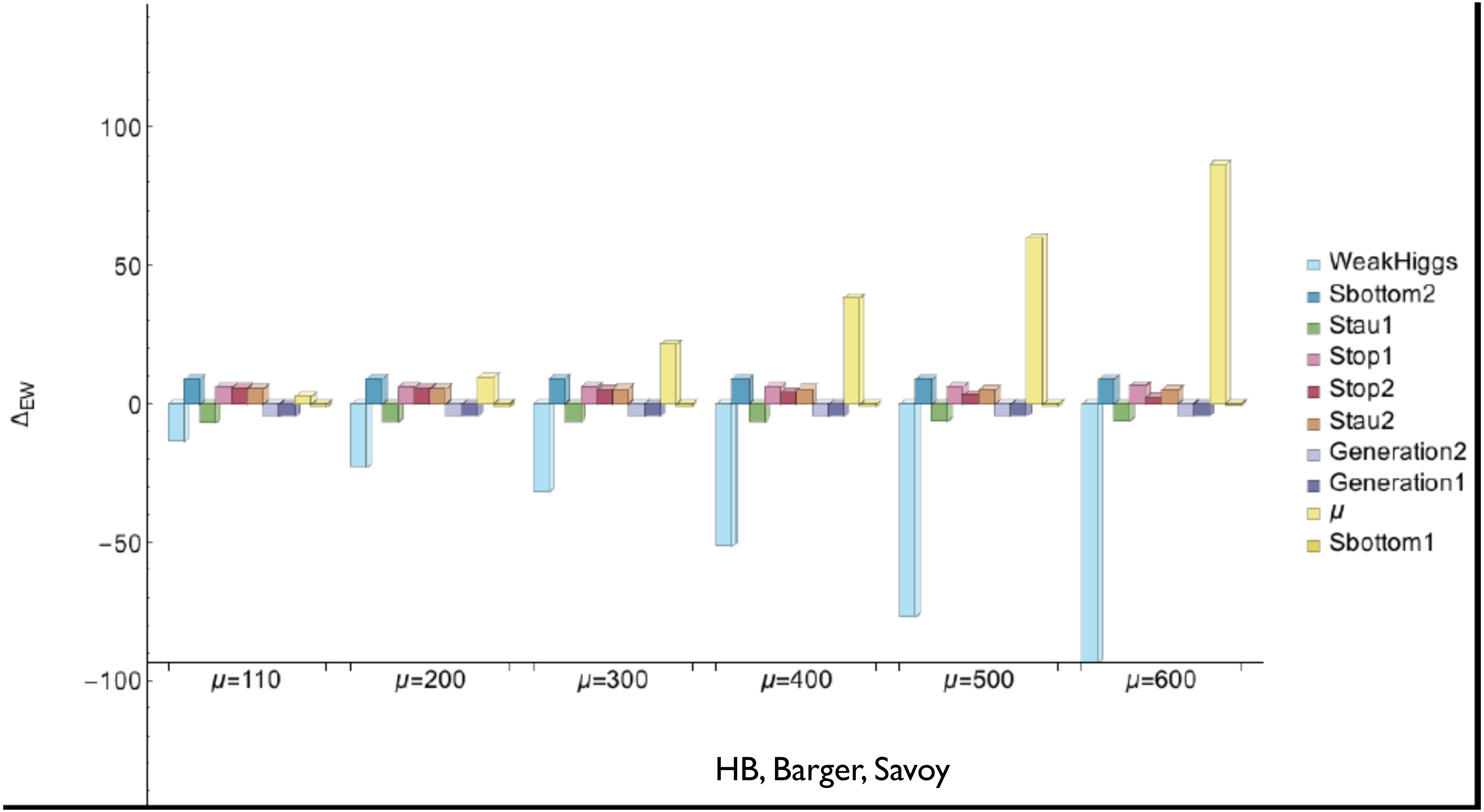
Due to ease of use and including radiative corrections, and due to its explicit model independence, we will use

$$\Delta_{EW}$$

for remainder of talk

hard wired in Isasugra

How much is too much fine-tuning?



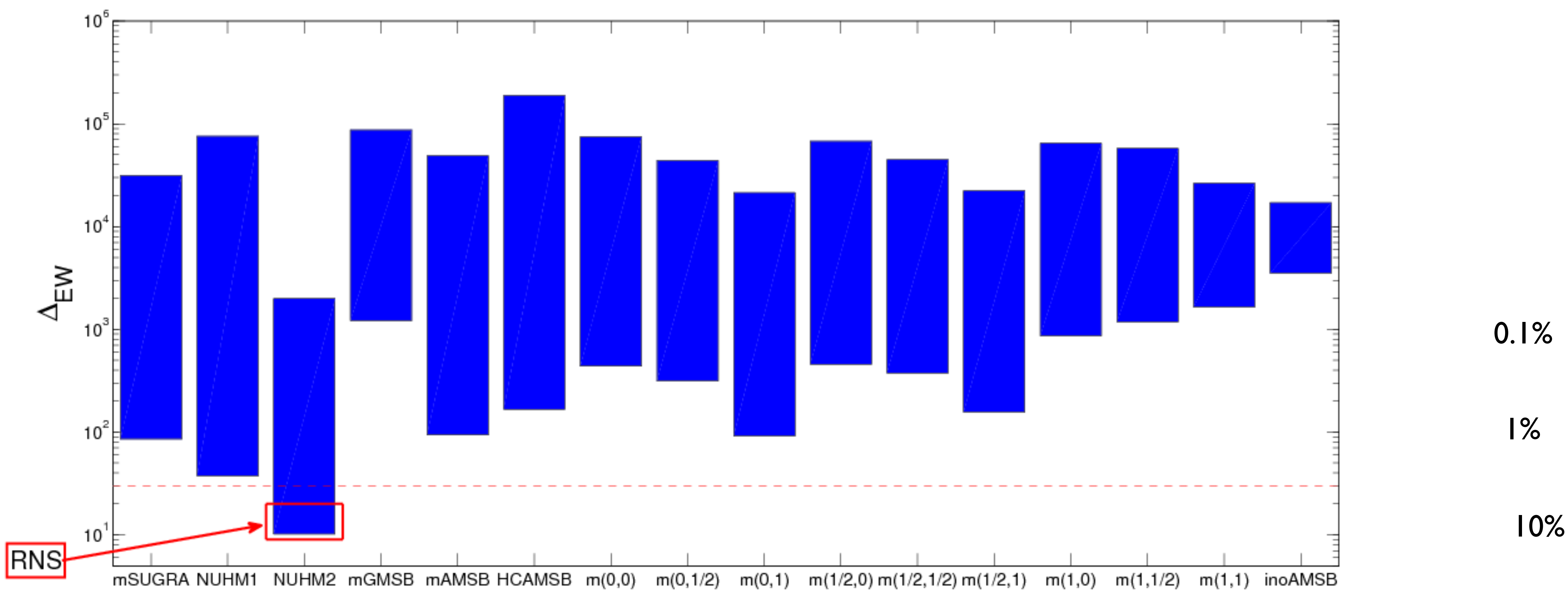
HB, Barger, Savoy

Visually, large fine-tuning has already developed by $\mu \sim 350$ or $\Delta_{EW} \sim 30$

Δ_{EW} is highly selective:
most constrained models are ruled out
except NUHM2 and its generalizations:

J. Ellis, K. Olive and Y. Santoso, *Phys. Lett.* B **539** (2002) 107; J. Ellis, T. Falk, K. Olive and Y. Santoso, *Nucl. Phys.* B **652** (2003) 259; H. Baer, A. Mustafayev, S. Profumo, A. Belyaev and X. Tata, *J. High Energy Phys.* 0507 (2005) 065.

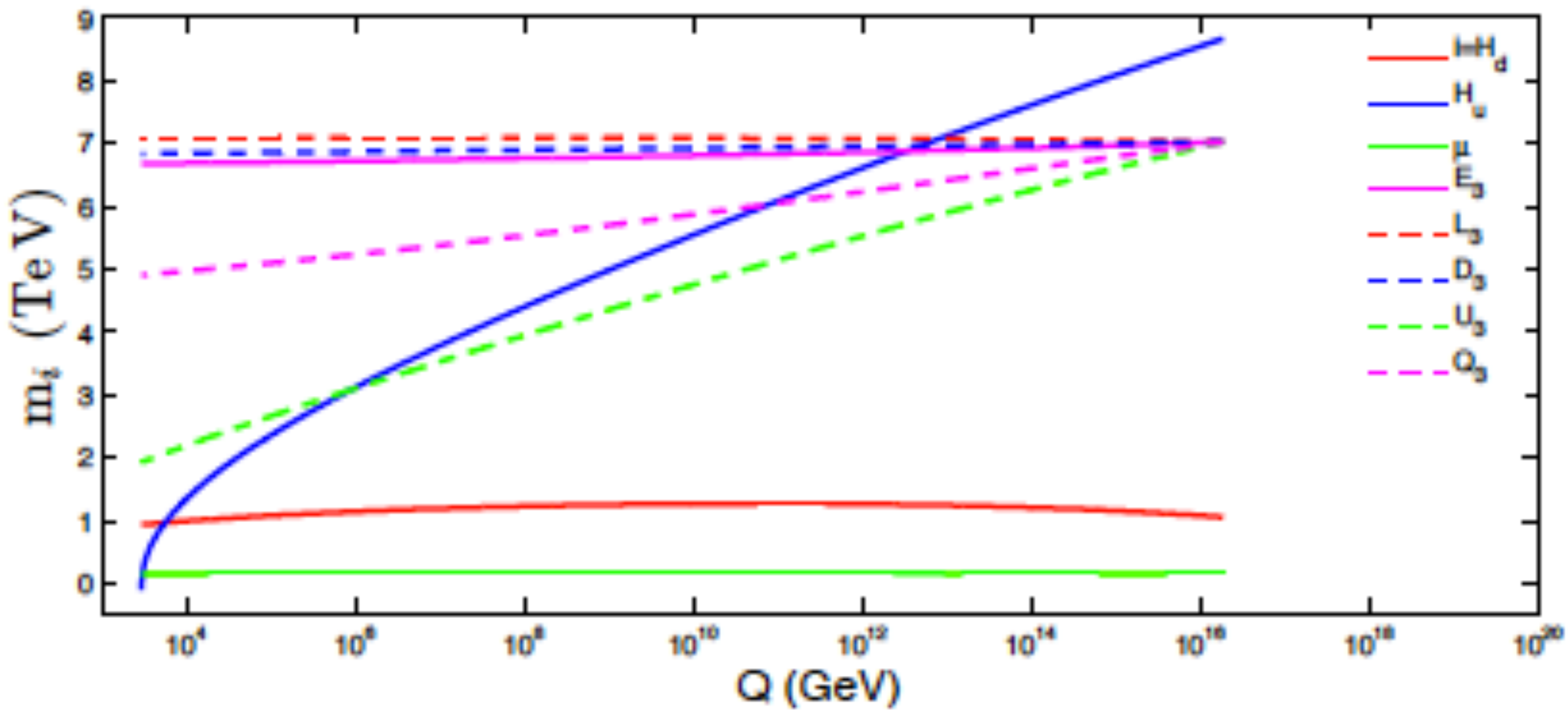
scan over p-space with m(h)=125.5+-2.5 GeV:



HB, Barger, Mickelson,Padeffke-Kirkland, PRD89 (2014) 115019

Applied properly, all three measures agree:

naturalness is unambiguous and highly predictive!



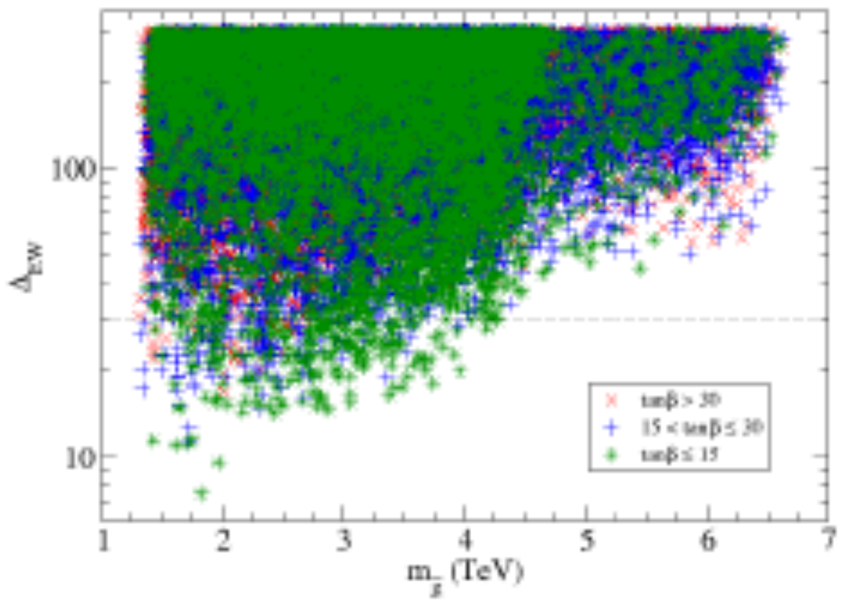
Radiatively-driven natural SUSY, or RNS:

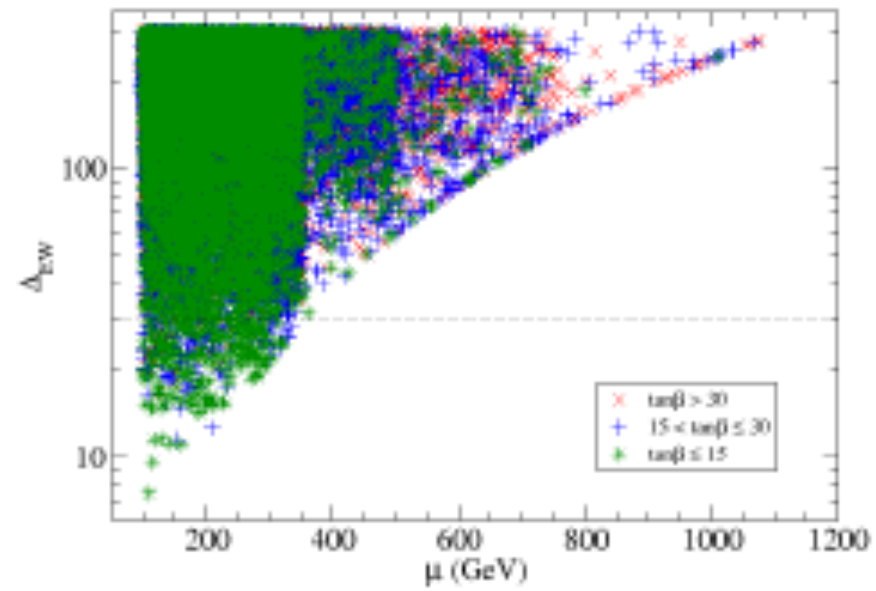
(typically need mHu~25-50% higher than m0)

H. Baer, V. Barger, P. Huang, A. Mustafayev and X. Tata, *Phys. Rev. Lett.* **109** (2012) 161802.

H. Baer, V. Barger, P. Huang, D. Mickelson, A. Mustafayev and X. Tata, *Phys. Rev.* **D 87** (2013) 115028 [arXiv:1212.2655 [hep-ph]].

Upper bounds on sparticle masses:



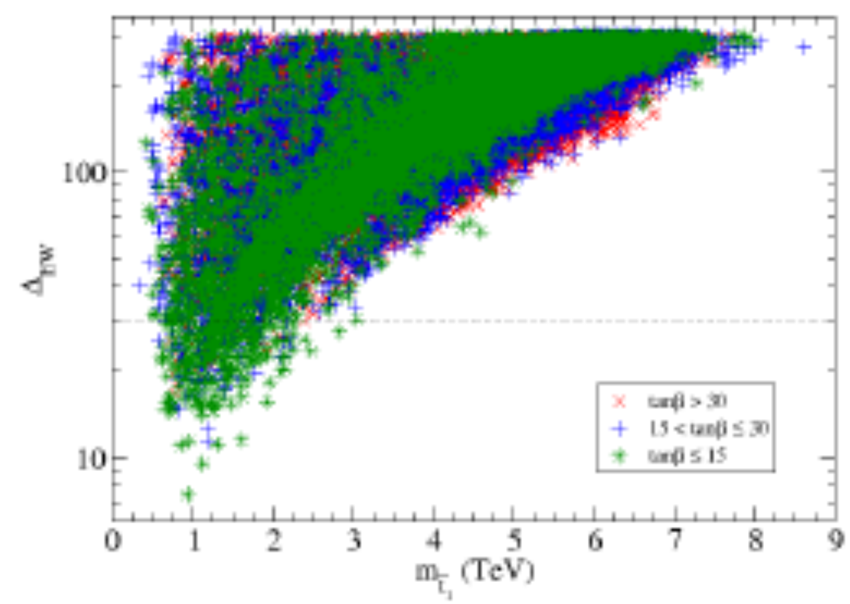


$\Delta_{EW} < 30$ upper bounds:

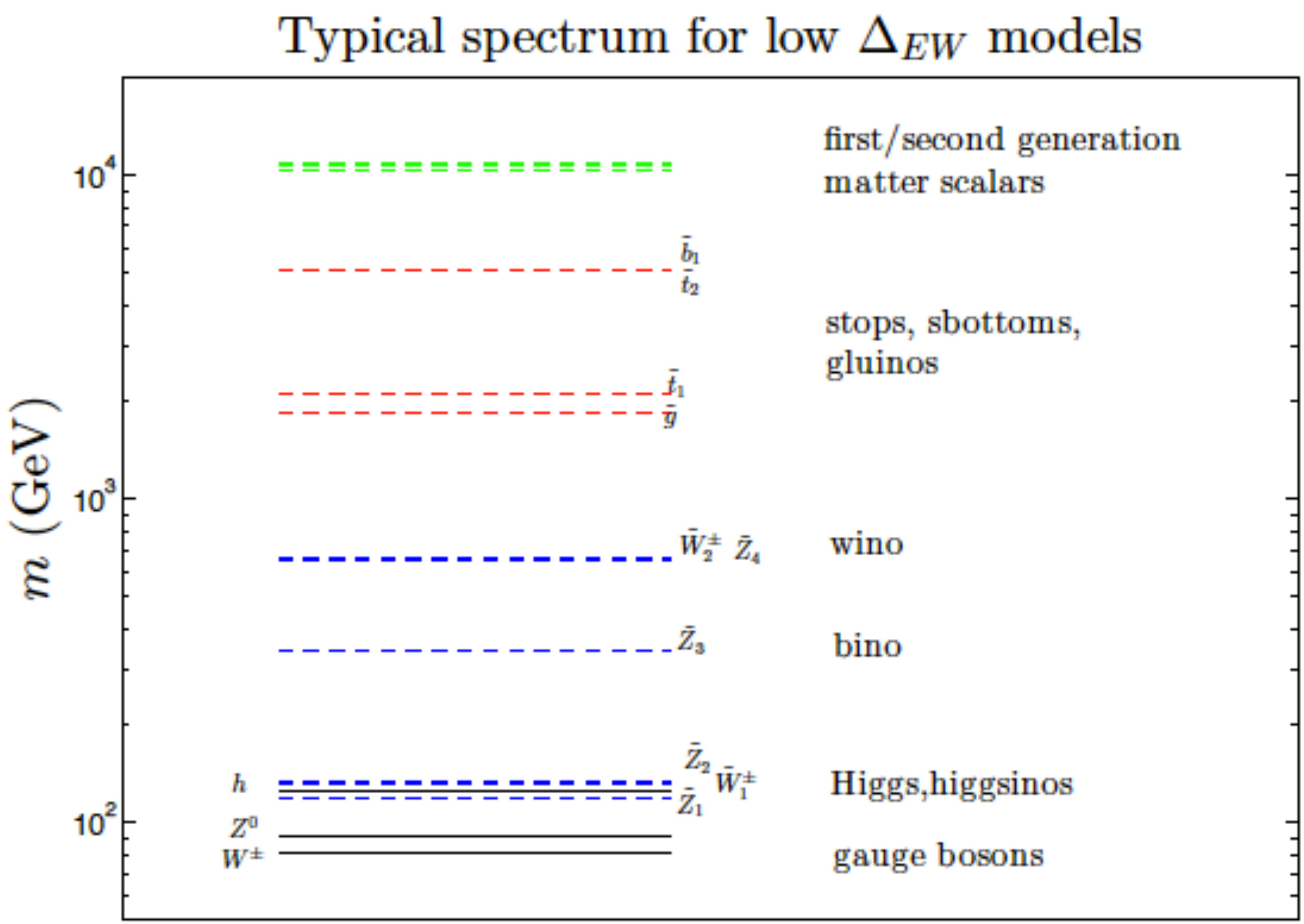
m(gluino)<4 TeV

mu<350 GeV

m(t1)<3 TeV



higher than old NS models and
allows for m(h)~125 GeV within MSSM



There is a Little Hierarchy, but it is no problem

$$\mu \ll m_{3/2}$$

SUSY mu problem: mu term is SUSY, not SUSY breaking: expect mu~M(Pl) but phenomenology requires mu~m(Z)

- NMSSM: mu~m(3/2); beware singlets!
- Giudice-Masiero: mu forbidden by some symmetry: generate via Higgs coupling to hidden sector
- Kim-Nilles: invoke SUSY version of DFSZ axion solution to strong CP:

KN: PQ symmetry forbids mu term, but then it is generated via PQ breaking

$$\mu \sim \lambda f_a^2 / M_P$$

$$m_{3/2} \sim m_{hid}^2 / M_P$$

Little Hierarchy due to mismatch between PQ breaking and SUSY breaking scales?

$$f_a \ll m_{hid}$$

Higgs mass tells us where to look for axion!

$$m_a \sim 6.2 \mu\text{eV} \left(\frac{10^{12}\ \text{GeV}}{f_a} \right)$$

Little Hierarchy from radiative PQ breaking? exhibited within context of MSY model

Murayama, Suzuki, Yanagida (1992);
Gherghetta, Kane (1995)

Choi, Chun, Kim (1996)

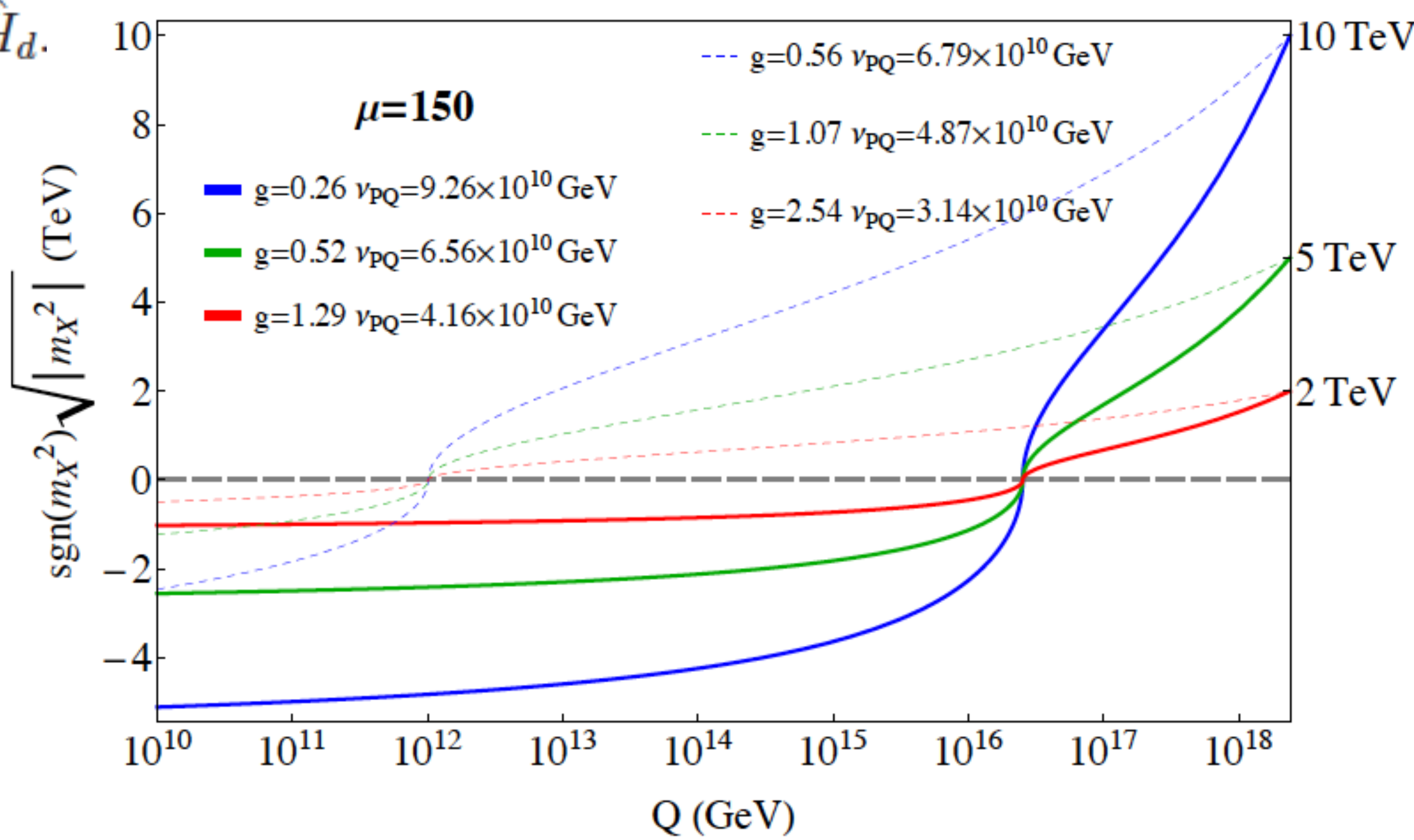
Bae, HB, Serce, PRD91 (2015) 015003

augment MSSM with PQ charges/fields:

$$\hat{f}' = \frac{1}{2}h_{ij}\hat{X}\hat{N}_i^c\hat{N}_j^c + \frac{f}{M_P}\hat{X}^3\hat{Y} + \frac{g}{M_P}\hat{X}\hat{Y}\hat{H}_u\hat{H}_d.$$

$$M_{N_i^c} = v_X \; h_i|_{Q=v_X}$$

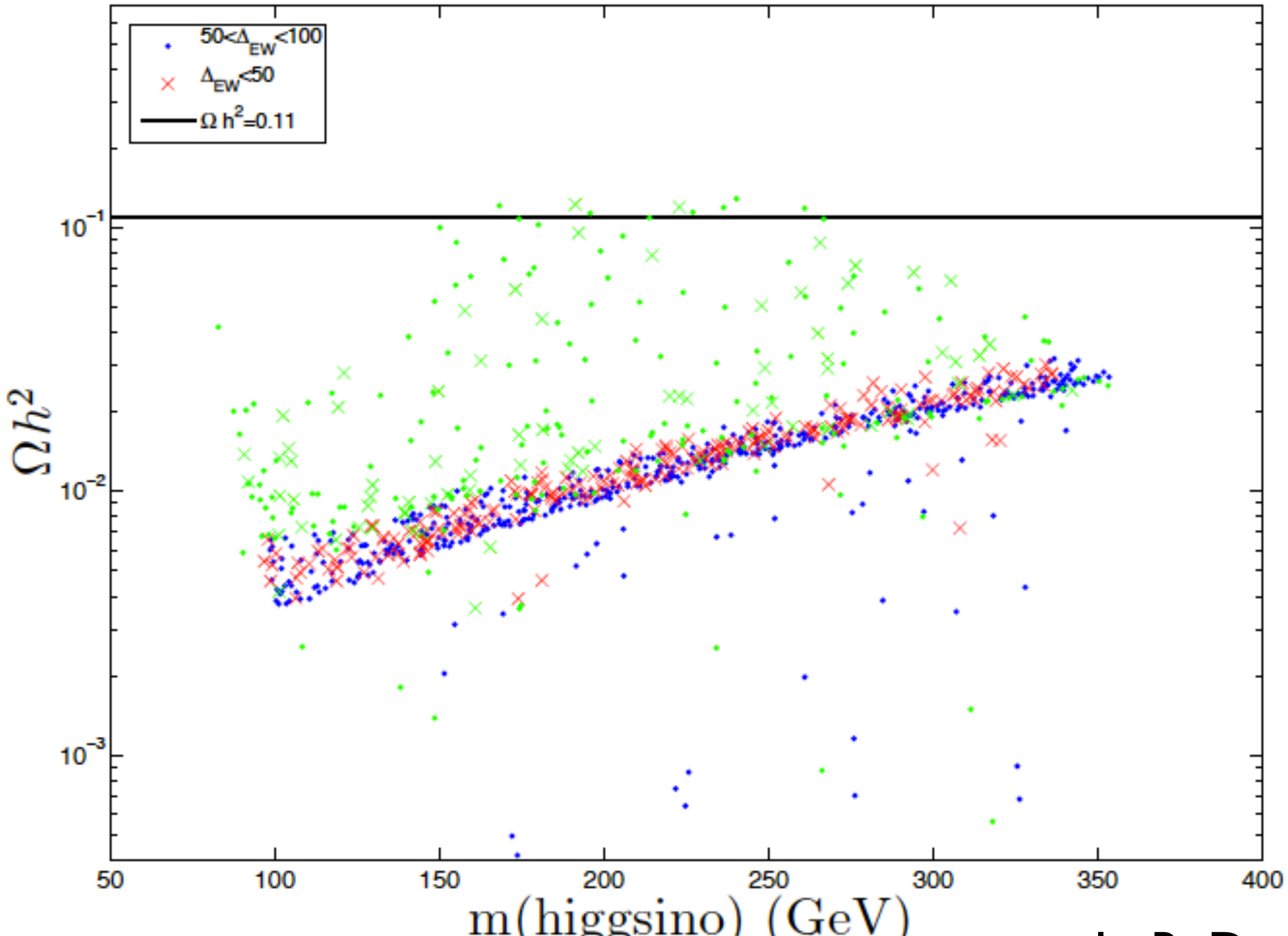
$$\mu = g\frac{v_X v_Y}{M_P} \; .$$



Large $m_{3/2}$ generates small $\mu \sim 100 - 200$ GeV!

Dark matter in Radiatively-driven Natural SUSY

Mainly higgsino-like WIMPs thermally underproduce DM



IsaReD

green: excluded;
red/blue:allowed

HB, Barger, Mickelson

Factor of 10-15 too low

But so far we have addressed only Part 1
of fine-tuning problem:

In QCD sector, the term $\frac{\bar{\theta}}{32\pi^2} F_{A\mu\nu} \tilde{F}_A^{\mu\nu}$ must occur

But neutron EDM says it is not there: strong CP problem

(frequently ignored by SUSY types)

Best solution after 35 years:
PQWW/KSVZ/DFSZ invisible axion

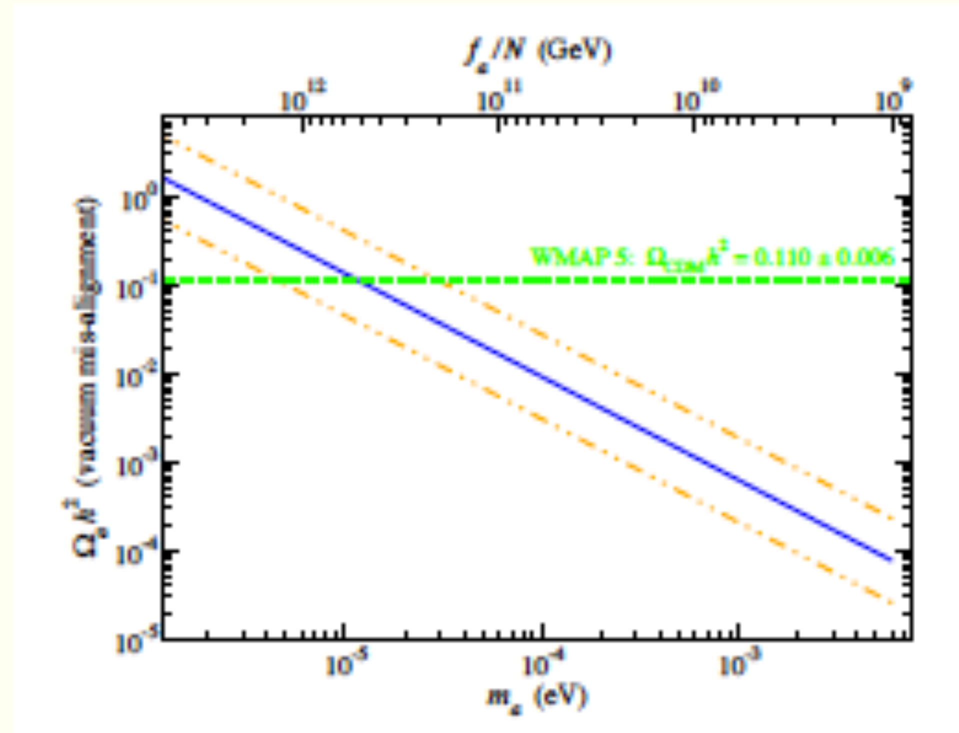
In SUSY, axion accompanied by axino and saxion

Changes DM calculus:
expect mixed WIMP/axion DM (2 particles)

Choi, Kim,Lee, Seto; HB,Lessa, Rajagopalan,Sreethawong

Axion cosmology

- ⋆ Axion field eq'n of motion: $\theta = a(x)/f_a$
  - $\ddot{\theta} + 3H(T)\dot{\theta} + \frac{1}{f_a^2}\frac{\partial V(\theta)}{\partial\theta} = 0$
  - $V(\theta) = m_a^2(T) f_a^2 (1 - \cos\theta)$
  - Solution for T large, $m_a(T) \sim 0$:
    $\theta = const.$
  - $m_a(T)$ turn-on ~ 1 GeV
- ⋆ a(x) oscillates,
  creates axions with $\vec{p} \sim 0$:
  production via vacuum mis-alignment
- ⋆ $\Omega_a h^2 \sim \frac{1}{2}\left[\frac{6\times 10^{-6} eV}{m_a}\right]^{7/6} \theta_i^2 h^2$
- ⋆ astro bound: stellar cooling $\Rightarrow f_a \gtrsim 10^9 GeV$

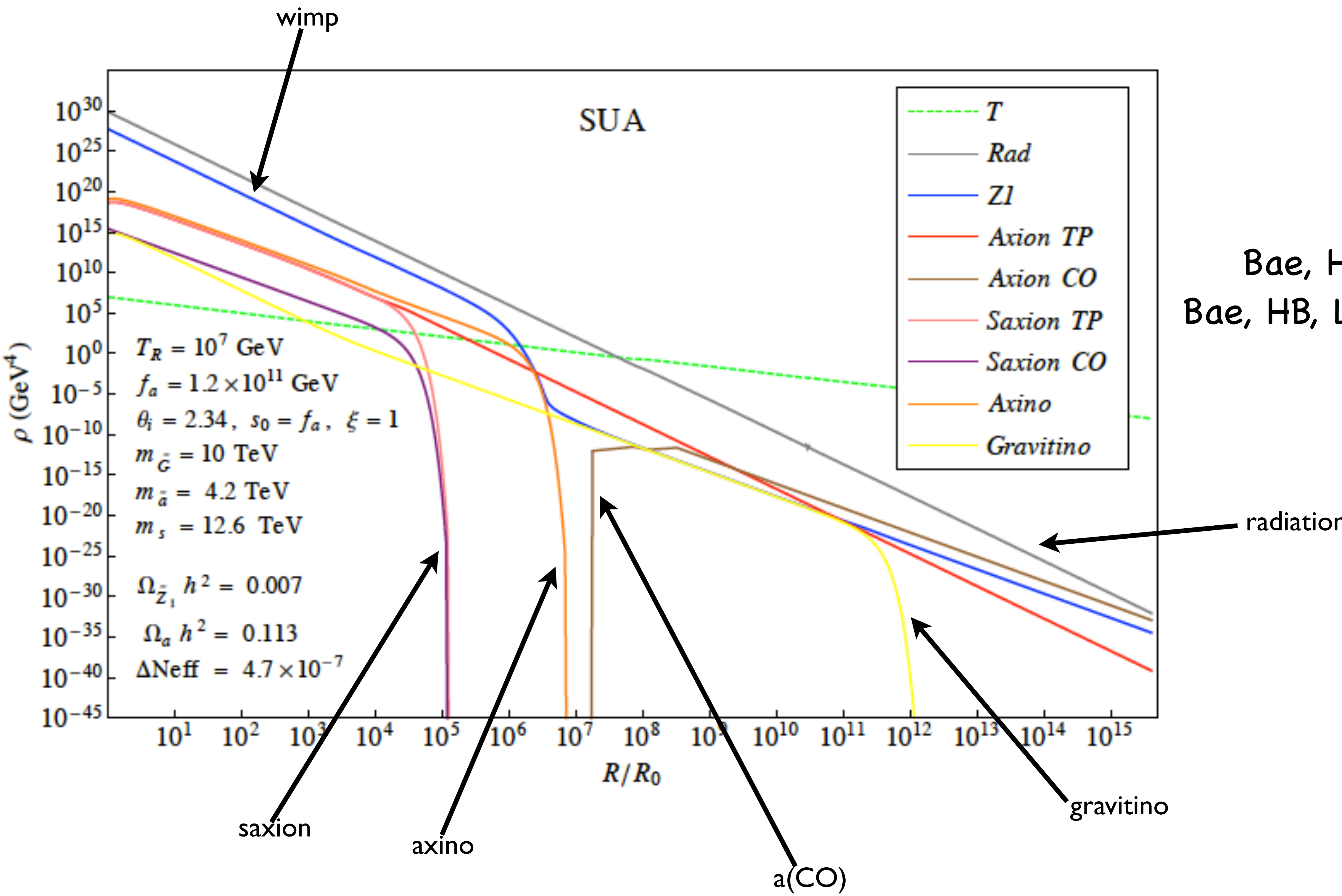


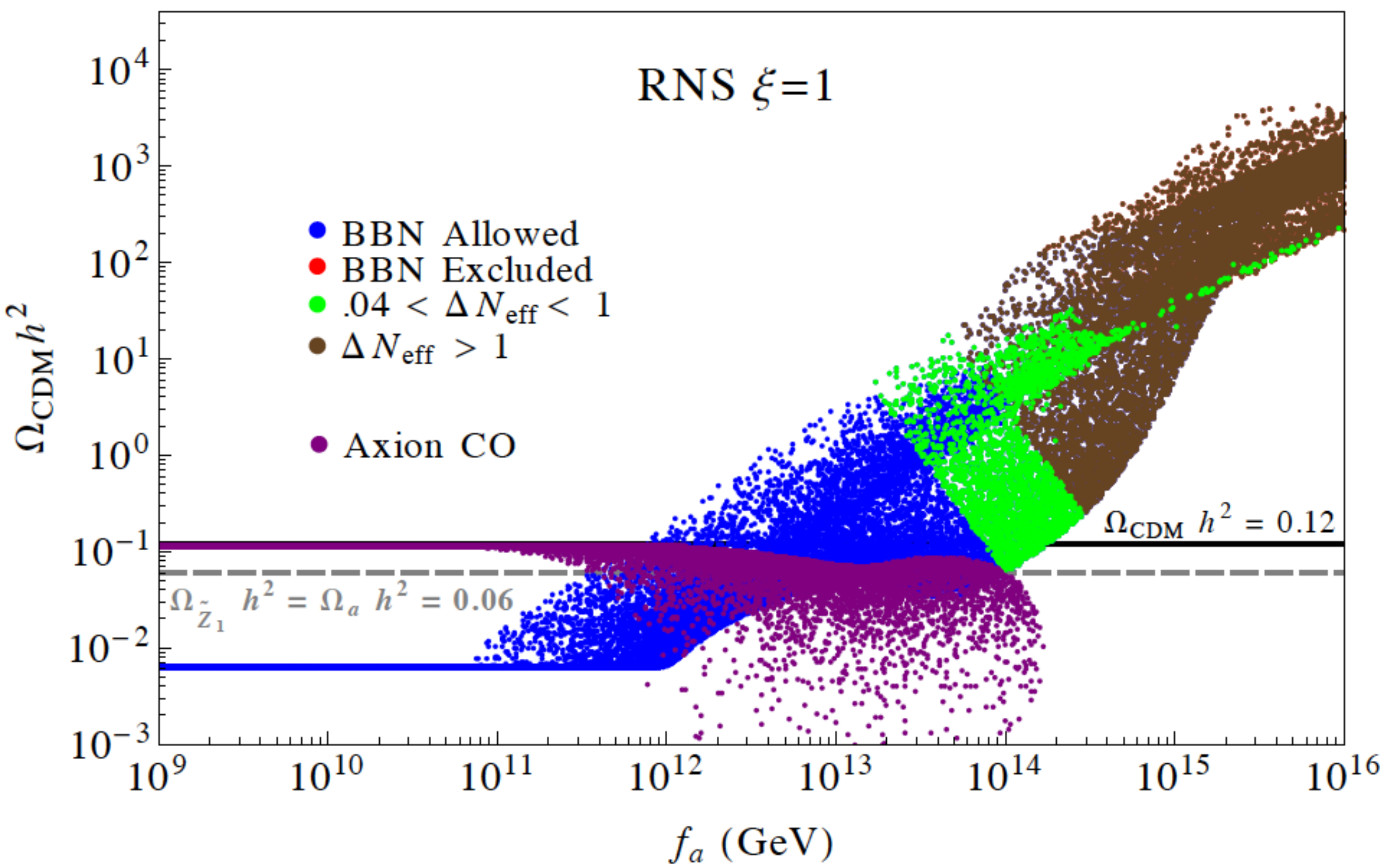
mixed axion-neutralino production in early universe

- neutralinos: thermally produced (TP) or NTP via $\tilde{a}$, s or $\tilde{G}$ decays
  - re-annihilation at $T_D^{s,\tilde{a}}$
- axions: TP, NTP via $s \to aa$, bose coherent motion (BCM)
- saxions: TP or via BCM
  - $s \to gg$: entropy dilution
  - $s \to SUSY$: augment neutralinos
  - $s \to aa$: dark radiation ($\Delta N_{eff} < 1.6$)
- axinos: TP
  - $\tilde{a} \to SUSY$ augments neutralinos
- gravitinos: TP, decay to SUSY

DM production in SUSY DFSZ: solve eight coupled Boltzmann equations

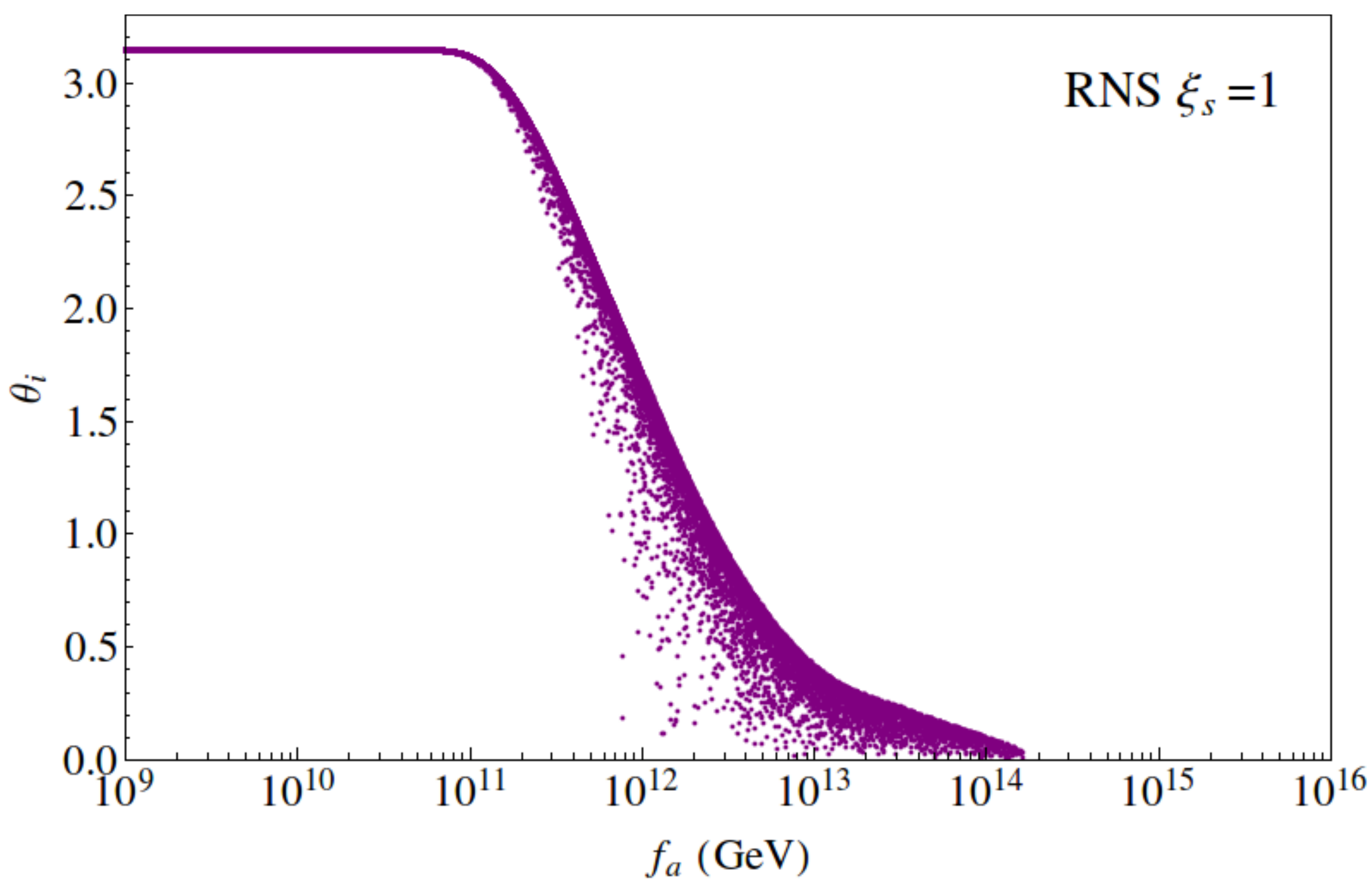
Bae, HB, Chun;
Bae, HB, Lessa, Serce





higgsino abundance

axion abundance

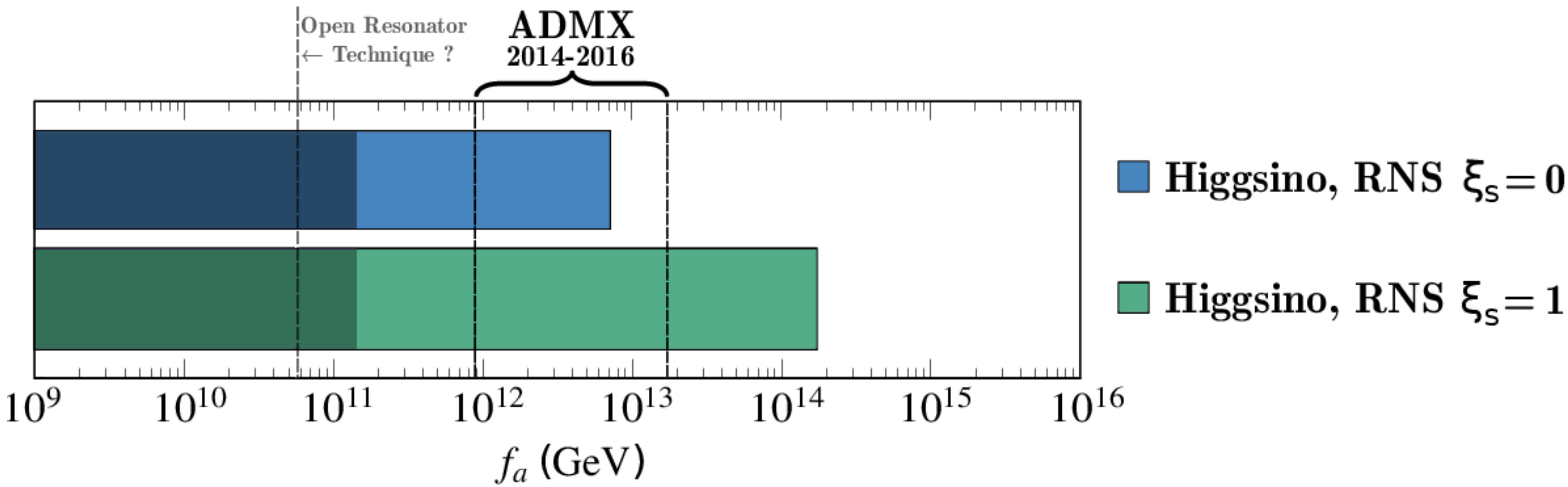


mainly axion CDM

for fa<~10^12 GeV;

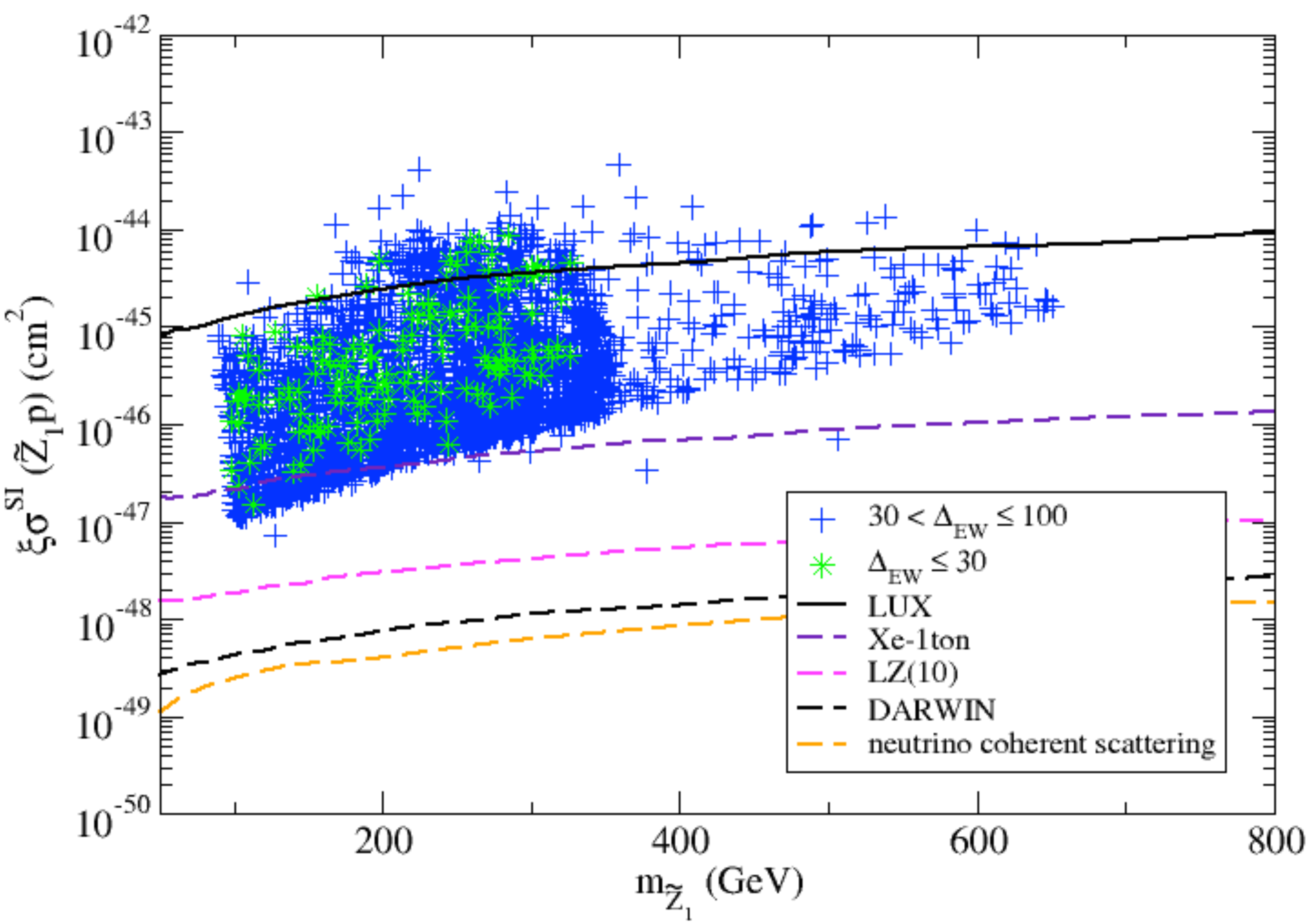
for higher fa, then get increasing wimp abundance

Bae, HB,Lessa,Serce



range of f_a expected from SUSY
with radiatively-driven naturalness
compared to ADMX axion reach

Direct higgsino detection rescaled for minimal local abundance



Bae, HB, Barger,Savoy,Serce

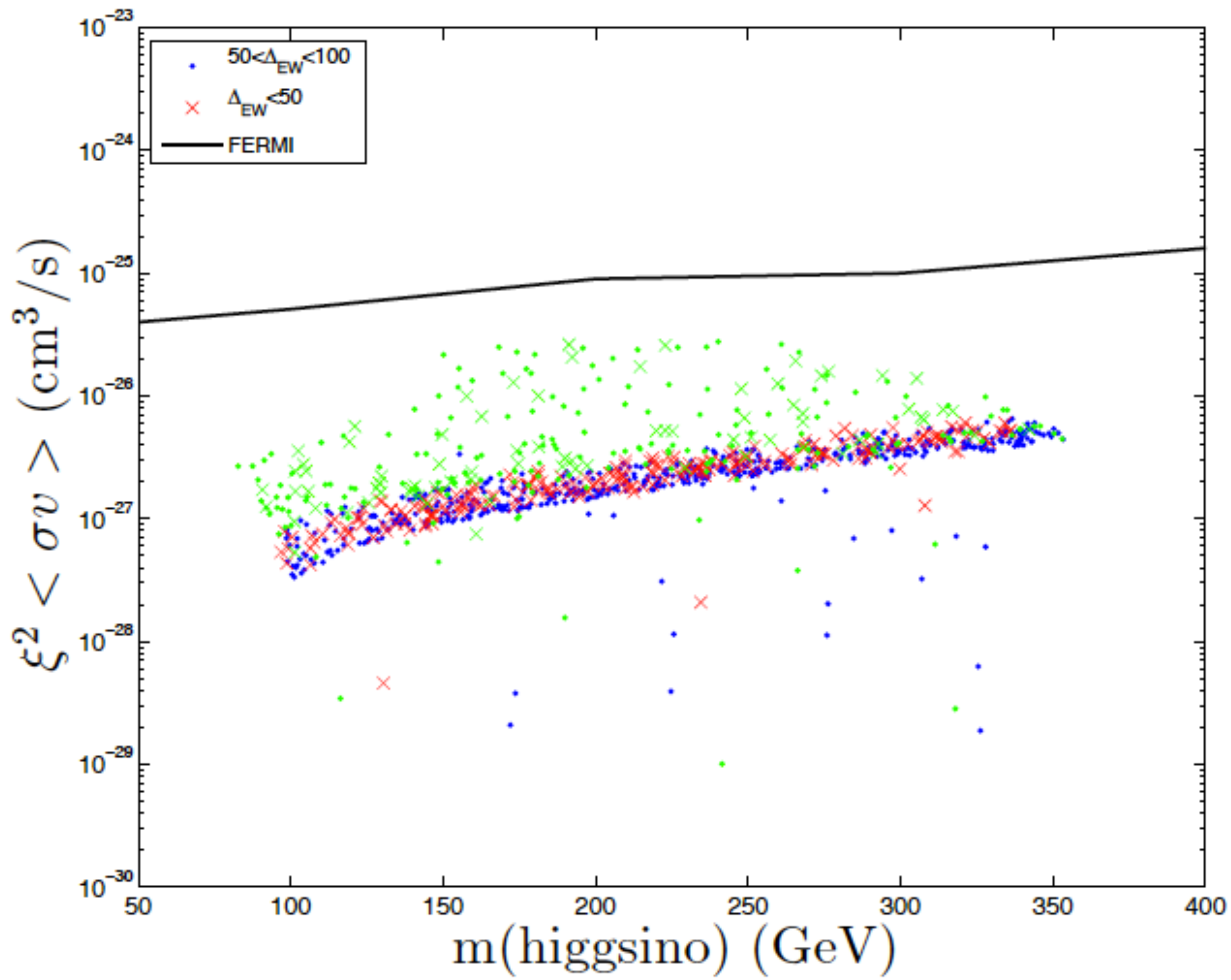
$$\mathcal{L} \ni -X_{11}^h \overline{\widetilde{Z}}_1 \widetilde{Z}_1 h$$

$$X_{11}^h = -\frac{1}{2}\left(v_2^{(1)} \sin\alpha - v_1^{(1)} \cos\alpha\right)\left(g v_3^{(1)} - g' v_4^{(1)}\right)$$

Deployment of Xe-1ton, LZ, SuperCDMS coming soon!

Can test completely with ton scale detector or equivalent (subject to minor caveats)

Higgsino detection via halo annihilations:



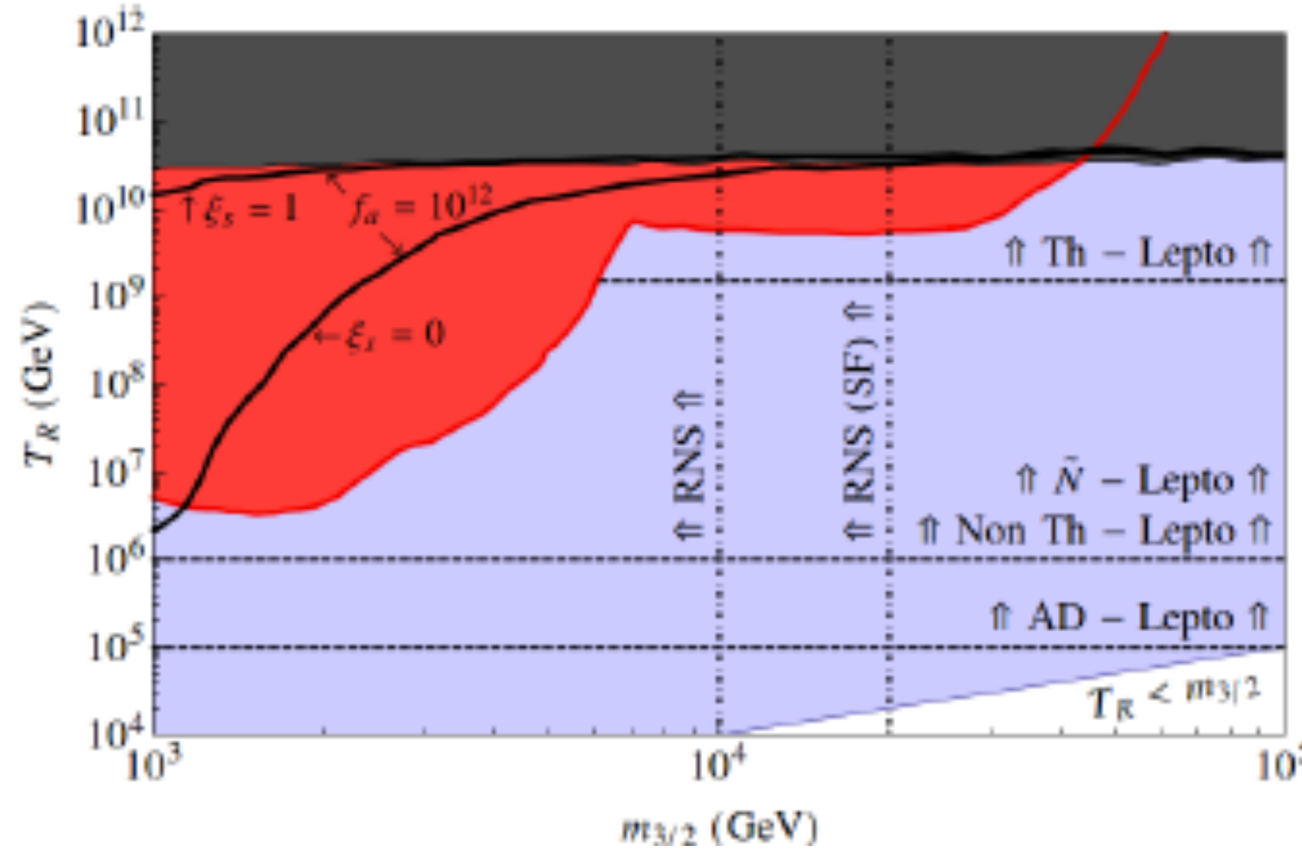
green: excluded by Xe-100

annihilation rate is high but rescaling is squared

Gamma-ray sky signal is factor 10-20 below current limits

Baryogenesis scenarios for radiative natural SUSY

- thermal leptogenesis
- non-thermal (inflaton decay)
- oscillating sneutrino
- Affleck-Dine (AD)



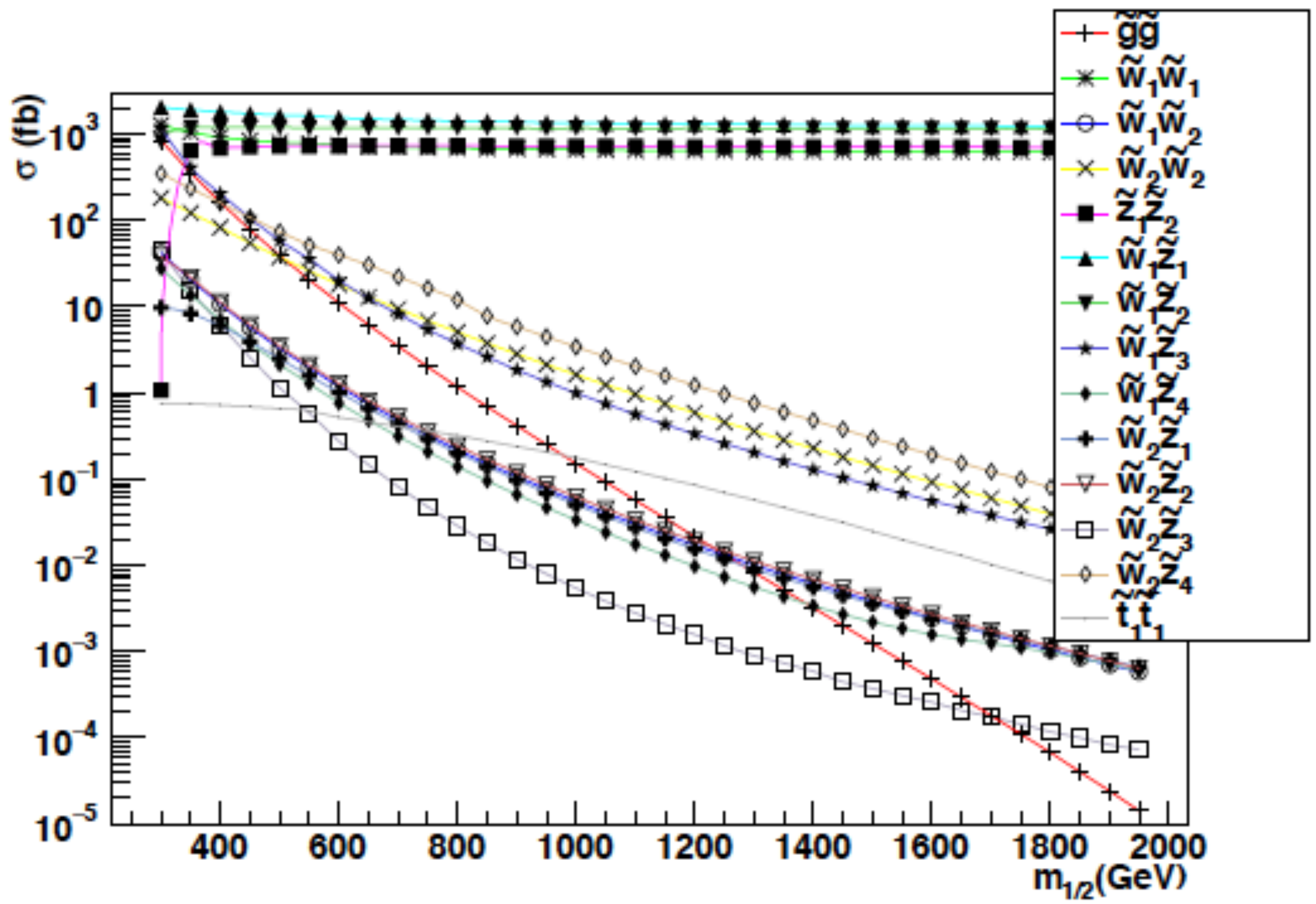
$$f_a = 10^{11},\ 10^{12}\ \text{GeV}$$

gravitino problem plus axino/saxion problem: still plenty room

Bae,HB, Serce, Zhang, arXiv:1510.00724

Prospects for discovering RNS at LHC and ILC

Sparticle prod'n along RNS model-line at LHC14:



higgsinos

gauginos

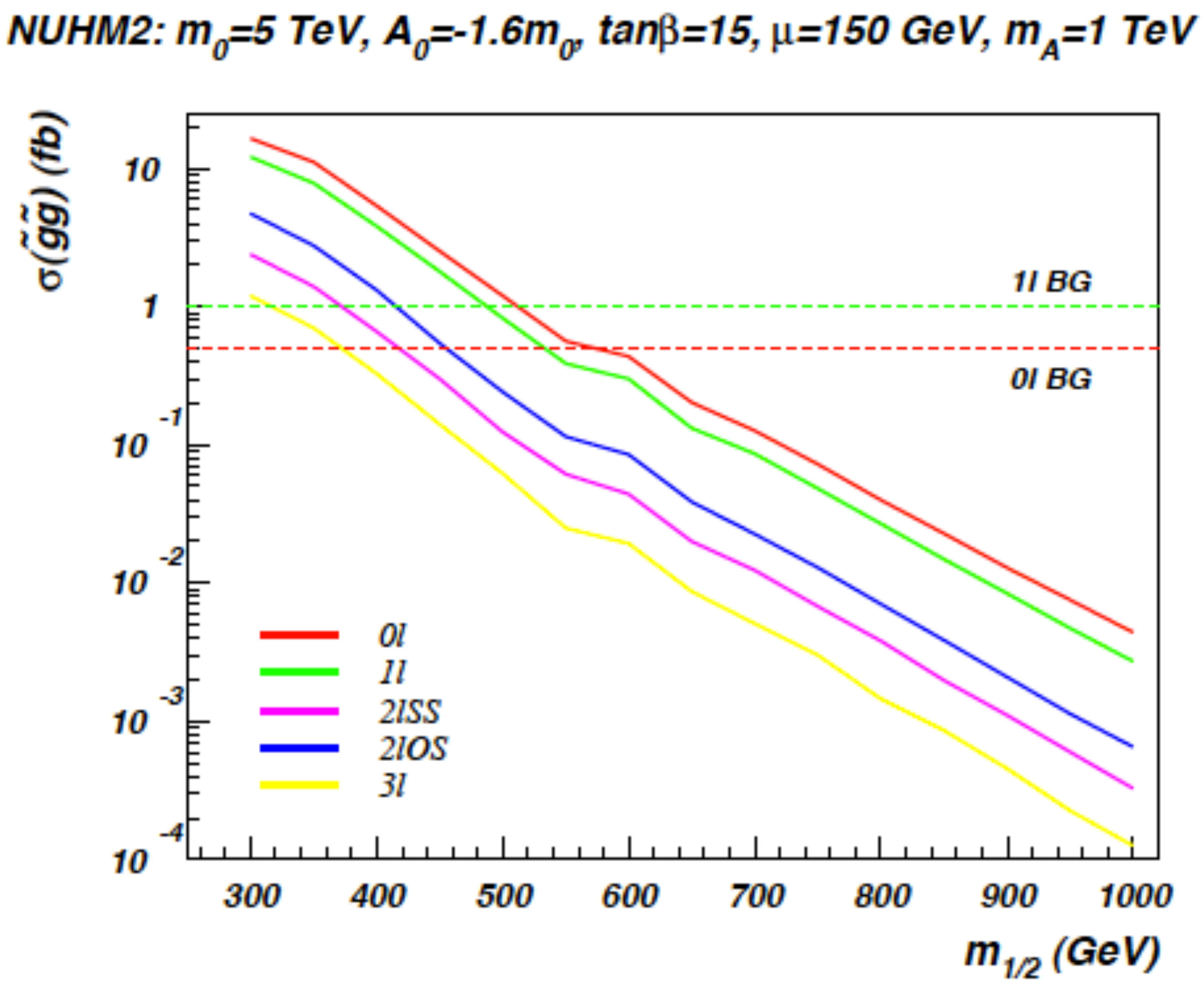
gluinos

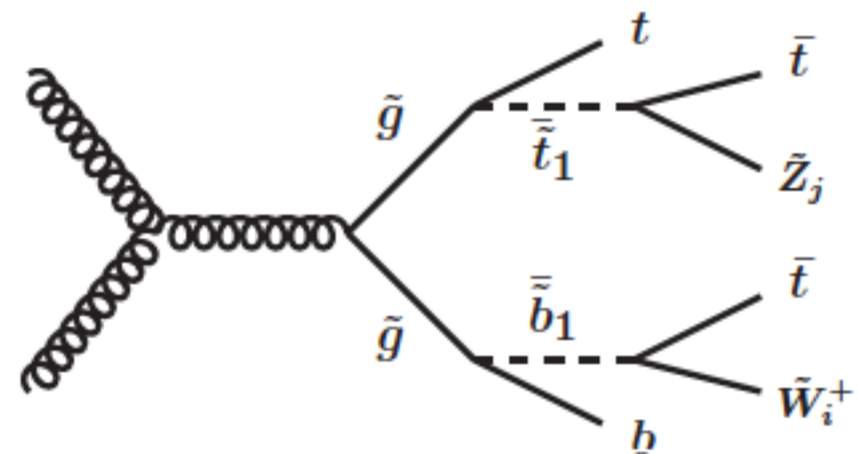
higgsino pair production dominant-but only soft visible energy release from higgsino decays

largest visible cross section: wino pairs

gluino pairs sharply dropping

gluino pair cascade decay signatures





Particle	dom. mode	BF
$\tilde{g}$	$\tilde{t}_1 t$	$\sim 100\%$
$\tilde{t}_1$	$b\widetilde{W}_1$	$\sim 50\%$
$\widetilde{Z}_2$	$\widetilde{Z}_1 f\bar{f}$	$\sim 100\%$
$\widetilde{Z}_3$	$\widetilde{W}_1^{\pm}W^{\mp}$	$\sim 50\%$
$\widetilde{Z}_4$	$\widetilde{W}_1^{\pm}W^{\mp}$	$\sim 50\%$
$\widetilde{W}_1$	$\widetilde{Z}_1 f\bar{f}'$	$\sim 100\%$
$\widetilde{W}_2$	$\widetilde{Z}_i W$	$\sim 50\%$

Table 1: Dominant branching fractions of various sparticles along the RNS model line for $m_{1/2} =$ 1 TeV.

Int. lum. (fb^{-1})	$\tilde{g}\tilde{g}$
10	1.4
100	1.6
300	1.7
1000	1.9

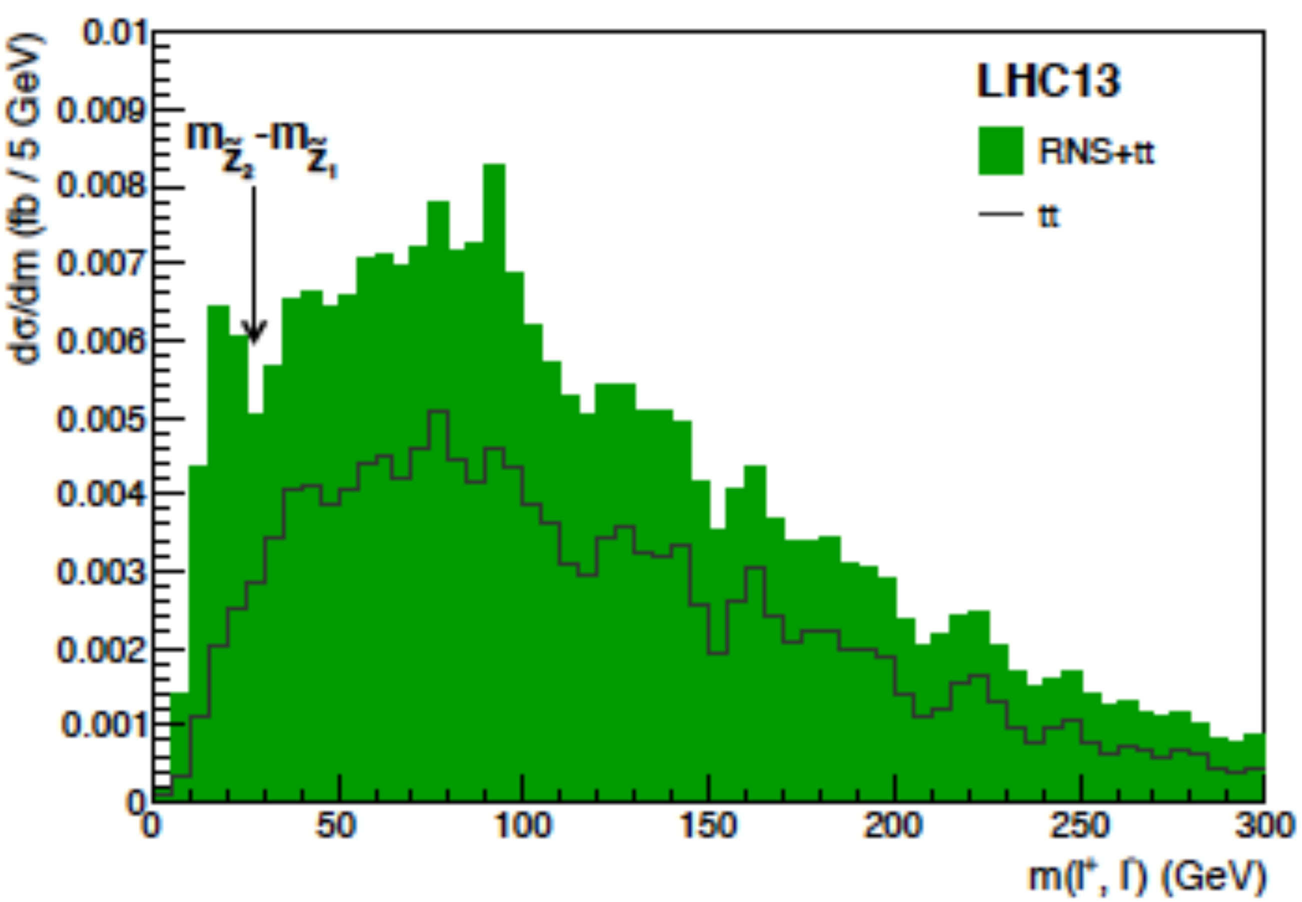
LHC14 reach in m(gluino) (TeV)

since m(gluino) extends to ~4 TeV, LHC14 can see about half the low EWFT parameter space in these modes

LHC14 has some reach for RNS; if a signal is seen, should be characteristic

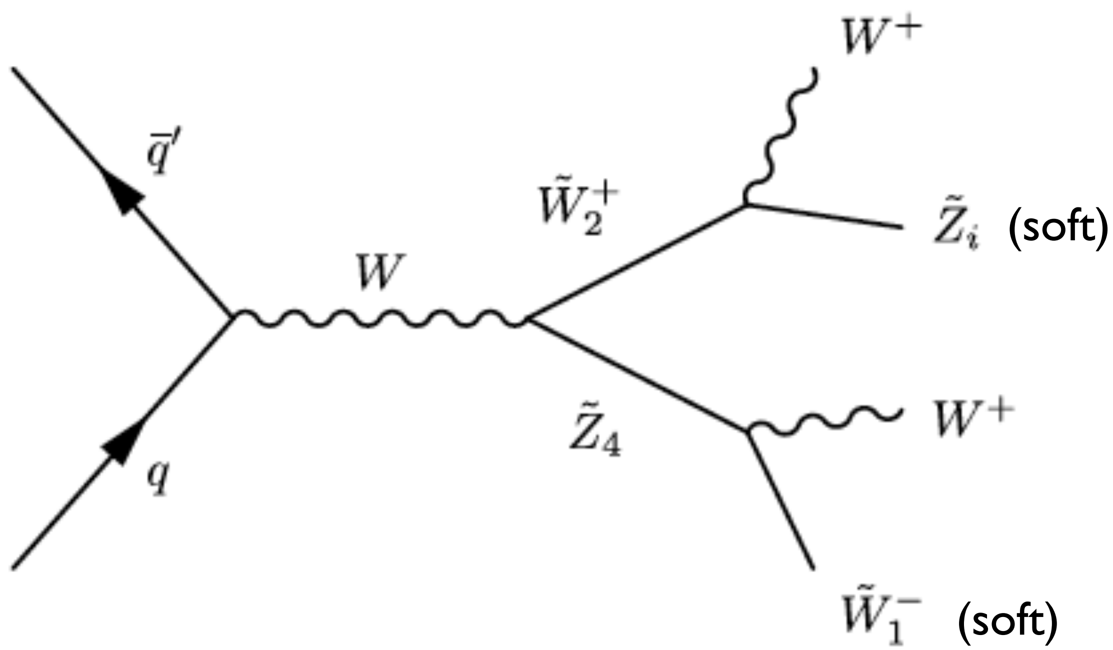
Int. lum. (fb^{-1})	$\tilde{g}\tilde{g}$	SSdB	$WZ \to 3\ell$	4ℓ
10	1.4	–	–	–
100	1.6	1.6	–	~ 1.2
300	1.7	2.1	1.4	$\gtrsim 1.4$
1000	1.9	2.4	1.6	$\gtrsim 1.6$

5σ reach of LHC14 in terms of $m_{\tilde{g}}$ for various Int. Lum.

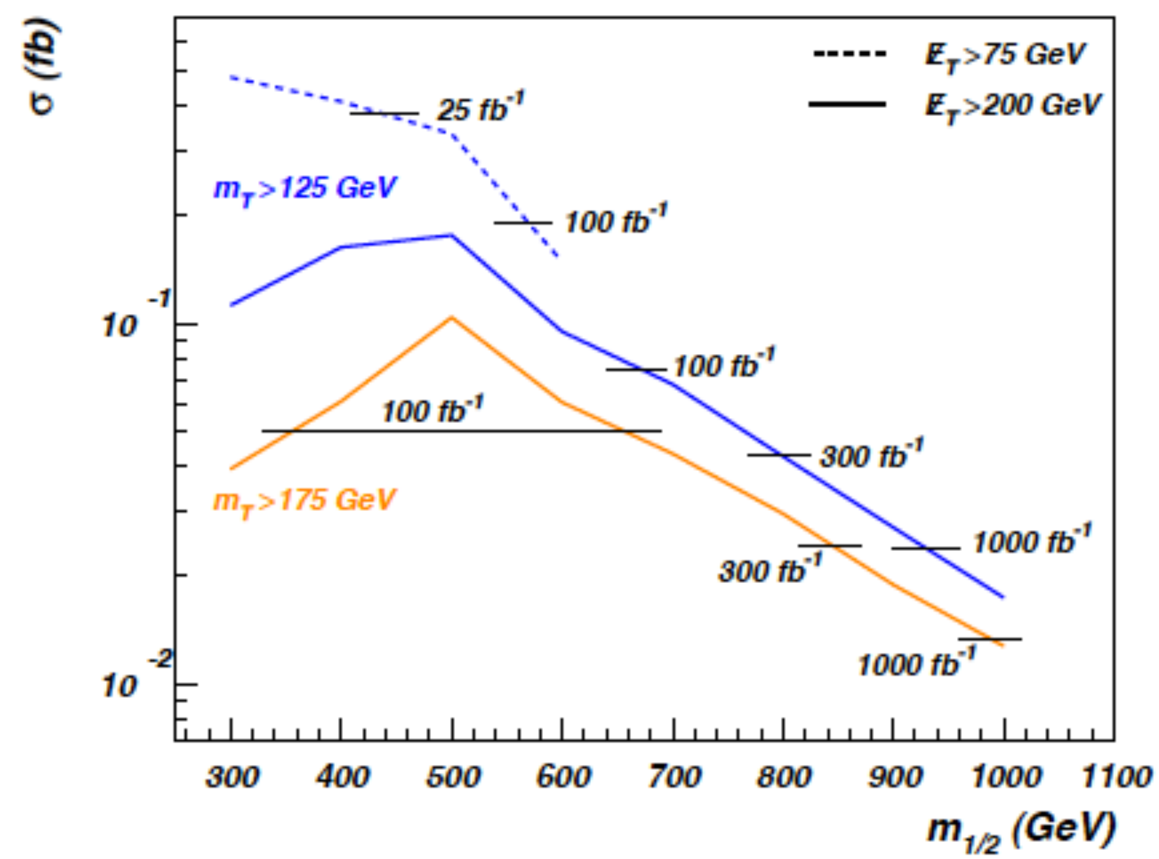


OS/SF dilepton mass edge apparent from cascade decays with z2->z1+l+lbar

Characteristic same-sign diboson (SSdB) signature from SUSY models with light higgsinos!



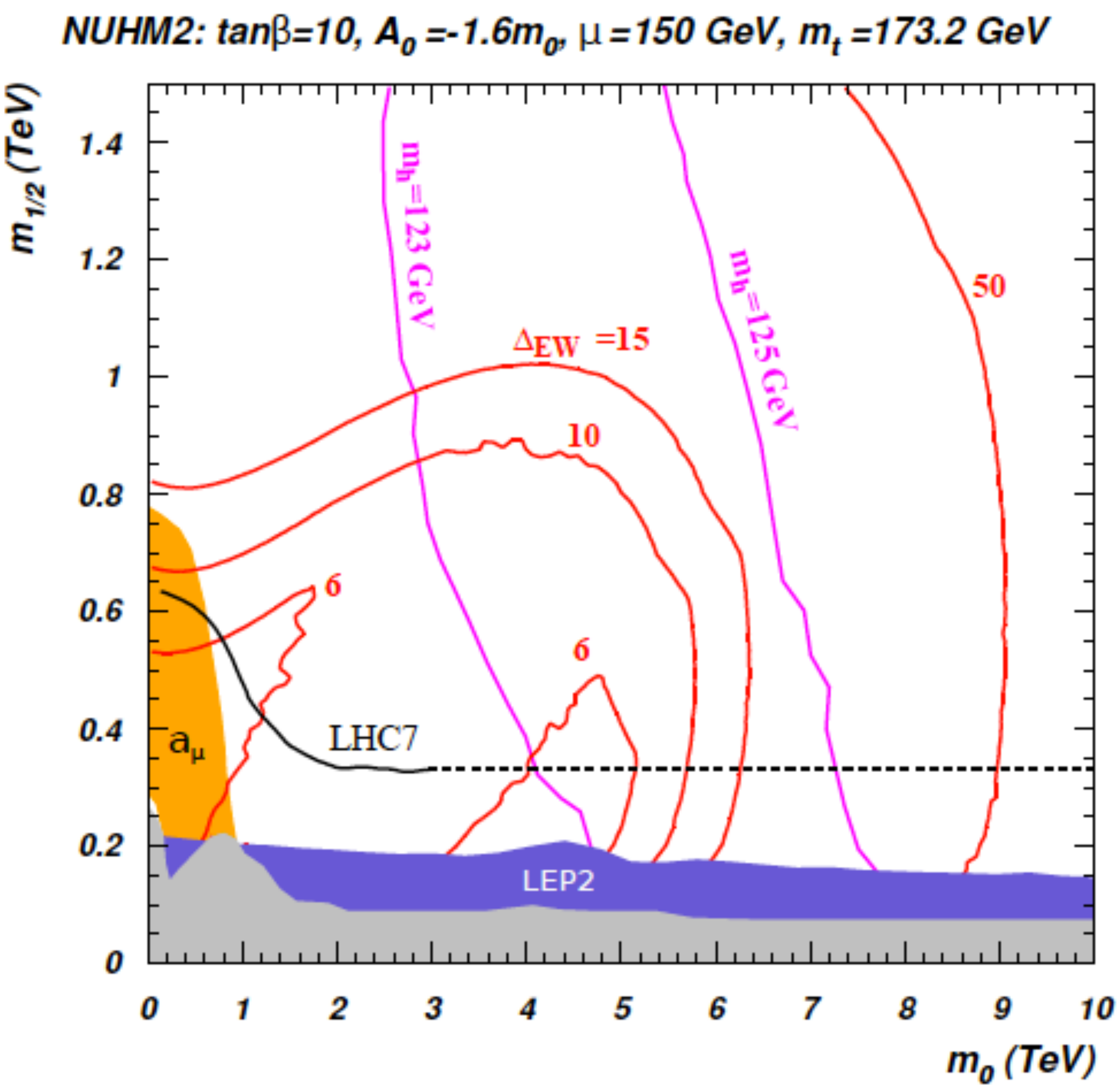
wino pair production



This channel offers best reach of LHC14 for RNS; it is also indicative of wino-pair prod'n followed by decay to higgsinos

H. Baer, V. Barger, P. Huang, D. Mickelson, A. Mustafayev, W. Sreethawong and X. Tata, *Phys. Rev. Lett.* **110** (2013) 151801.

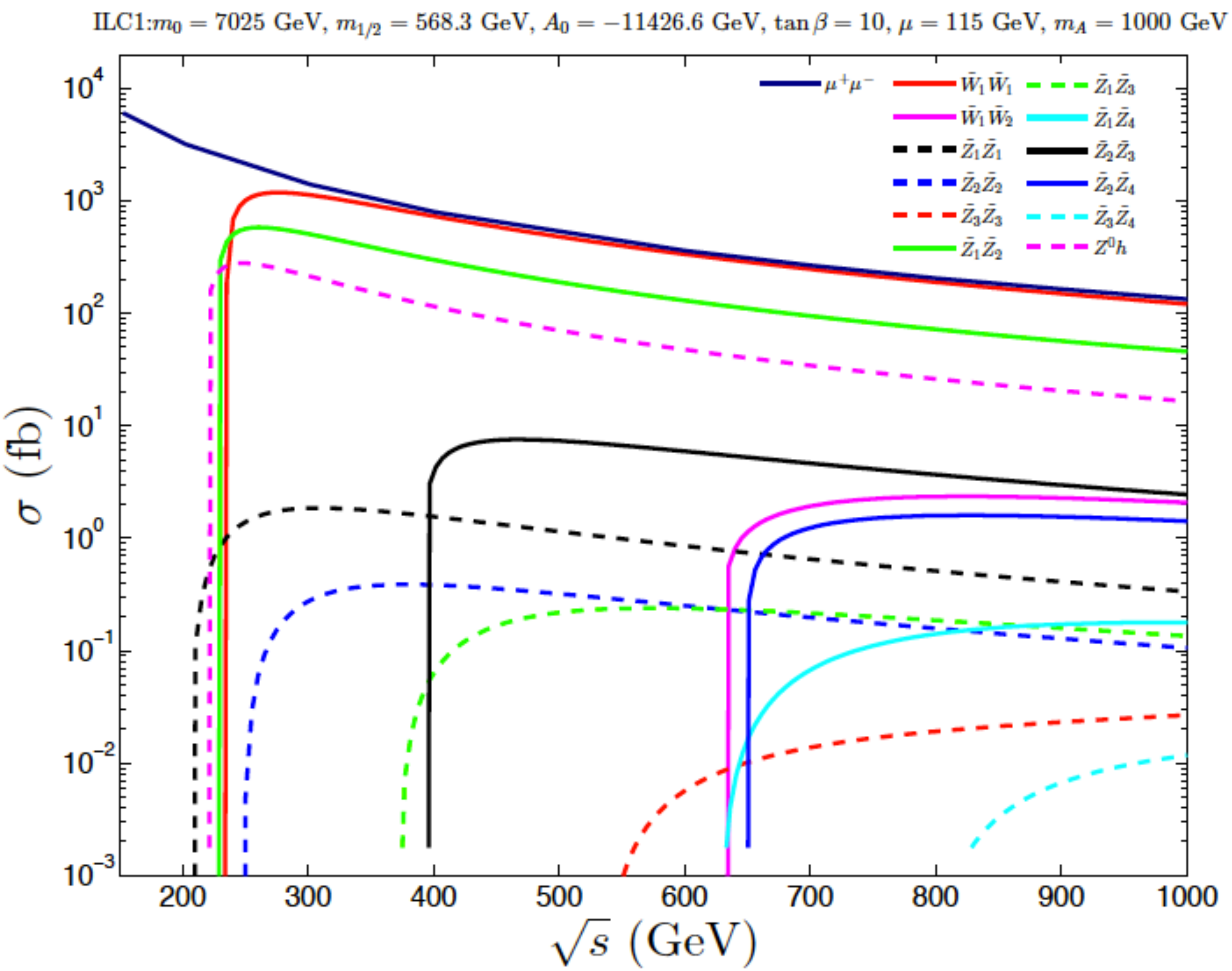
Good old m0 vs. mhf plane still viable, but require low mu (NUHM2)



$\mu = 150$ GeV throughout which is allowed for NUHM2

Smoking gun signature: light higgsinos at ILC:

ILC is Higgs/higgsino factory!



$$\sigma(higgsino) \gg \sigma(Zh)$$

10-15 GeV higgsino mass gaps no problem in clean ILC environment

HB, Barger, Mickelson, Mustafayev, Tata
arXiv:1404:7510

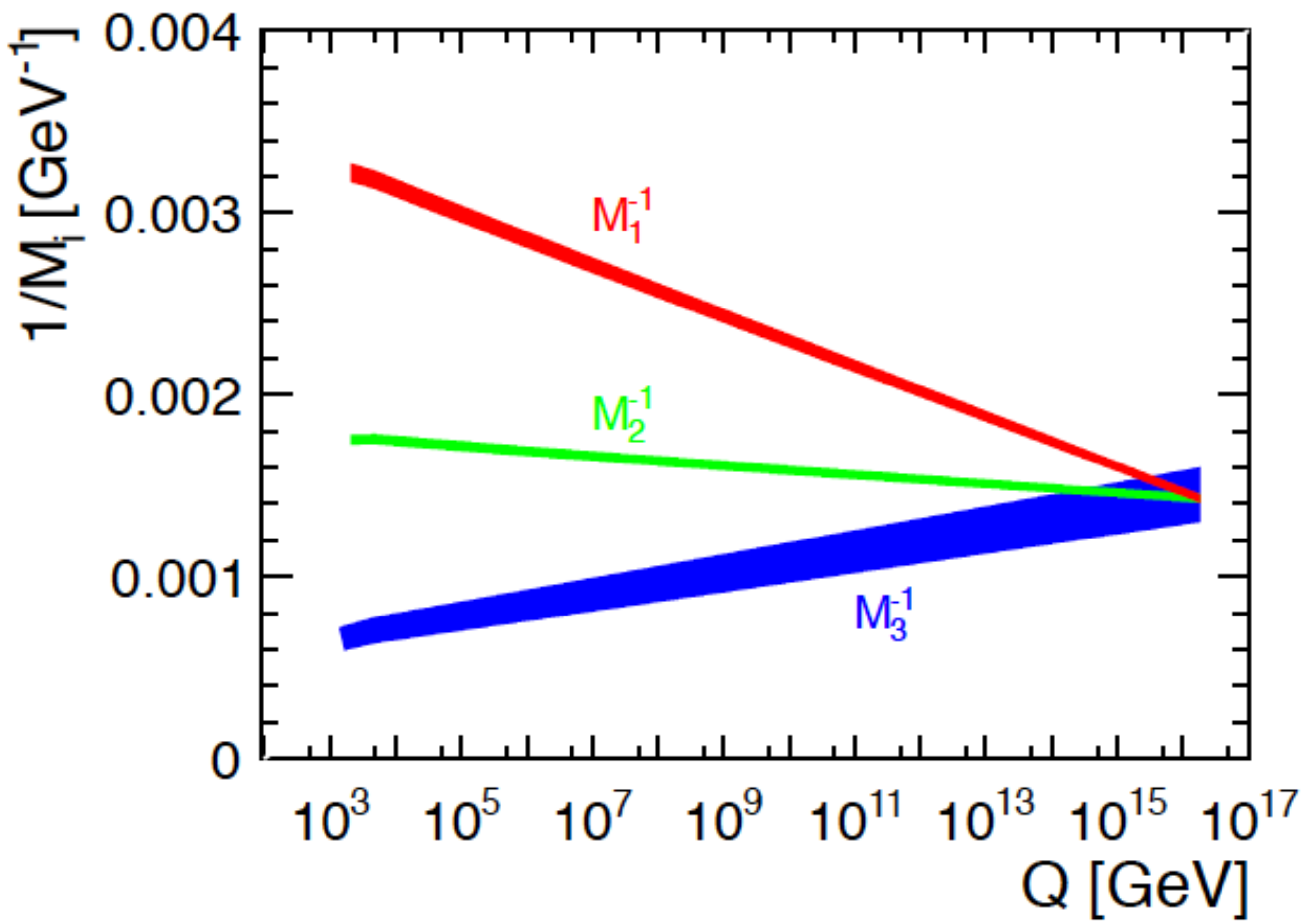
ILC either sees light higgsinos or natural SUSY dead

In SUSY with radiatively-driven naturalness, can still test unification

Higgsino mass $\Rightarrow \mu$

Higgsino mass gaps $\Rightarrow M_1$ and M_2

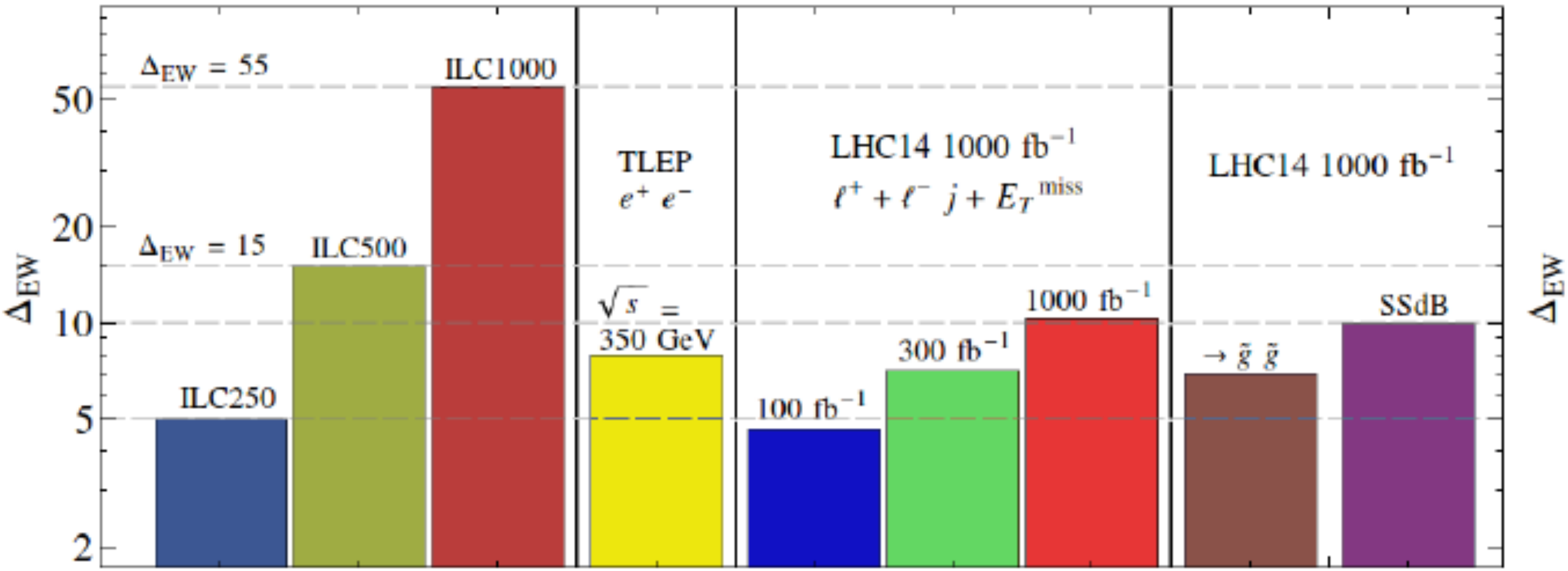
Measure M_3 at LHC else use unification to predict $m_{\tilde{g}}$



testing gaugino mass unification at ILC

HB, J. List

Future collider reach for naturalness



Bae, HB, Nagata, Serce

When to give up on naturalness in MSSM?
If ILC(600-700 GeV) sees no light higgsinos

message to MEXT committee

The simplest, most natural version of
supersymmetry predicts light higgsinos
with mass 100-300 GeV
(the closer to m(W,Z,h) the better)

These states are difficult, maybe impossible, to
detect at LHC while ILC would be a
higgsino factory:
can test naturalness, unification, dark matter…

ILC is highly worthy of the investment!

Conclusions: status of SUSY post LHC8

- SUSY EWFT non-crisis: EWFT allowed at 10% level in radiatively-driven natural SUSY: SUGRA GUT paradigm is just fine in NUHM2 but CMSSM/others fine-tuned
- naturalness maintained for mu~100-200 GeV; t1~1-2 TeV, t2~2-4 TeV, highly mixed; m(glno)~1-4 TeV
- LHC14 w/ 300 fb^-1 can see about half of RNS parameter space
- e+e- collider with sqrt(s)~500-600 GeV needed to find predicted light higgsino states
- Discovery of and precision measurements of light higgsinos at ILC!
- RNS spectra characterized by mainly higgsino-like WIMP: standard relic underabundance
- SUSY DFSZ/MSY invisible axion model: solves strong CP and mu problems while allowing for mu~m(Z)
- Expect mainly axion CDM with 5-10% higgsino-like WIMPs over much of p-space
- Ultimately detect both axion and higgsino-like WIMP