

Scalar self-force for highly eccentric orbits in Kerr spacetime

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in collaboration with

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Goals, overall plan of the computation

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Brief review of **effective-source (puncture-function) regularization**

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- separate **2+1D time-domain evolution** for each mode
- **worldtube scheme**
 - finite differencing across the worldtube boundary
 - moving the worldtube
- **computing the effective source** and puncture function
- **finite differencing near the particle** (where fields are only C^2)
- **mesh refinement**

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Sample results

Conclusions, Plans, Lessons Learned

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Work in progress: some goals accomplished, some not yet!

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- fixed mesh refinement; some (finer) grids follow the worldtube/particle
- [now] (almost) causally-disconnected spatial boundaries (with mesh refinement this isn't very expensive)
- [future] hyperboloidal outer boundary (Zenginoğlu, arXiv:1008.3809 = J. Comp. Phys. 230, 2286)

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If φ_p "closely-enough" approximates $\varphi_{\text{singular}}$ near the particle, then the self-force is given (exactly!) by $F_a = q(\nabla_a\varphi_r)|_{\text{particle}}$

I.e., $S_{\text{effective}}$ compensates for the fact that $\varphi_p \neq \varphi_{\text{singular}}$

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- the self-force is given (**exactly!**) by $F_a = q \sum_{m=0}^{\infty} (\nabla_a \varphi_{\text{num},m})|_{\text{particle}}$
- we actually do m -mode decomposition before introducing worldtube
 \Rightarrow worldtube “lives” in (t, r, θ) space, not full spacetime

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Boundary Conditions:

- [now] use almost causally-disconnected spatial boundaries (with mesh refinement this isn't very expensive)
- [future] **hyperboloidal outer boundary** (Zenginoğlu, arXiv:1008.3809 = J. Comp. Phys. 230, 2286)

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Current Status

Equatorial eccentric orbits:

- elliptic-integral puncture fn & effective src
- **worldtube moves** in (r, θ) to follow the particle around the orbit
- fixed mesh refinement: typically 5 levels, $\Delta r_* = M/4$ to $M/64$; finest 3 refinement levels move to follow worldtube
- typical worldtube size particle $\pm 5M$ in r_* , particle $\pm \pi/8$ (22.5°) in θ
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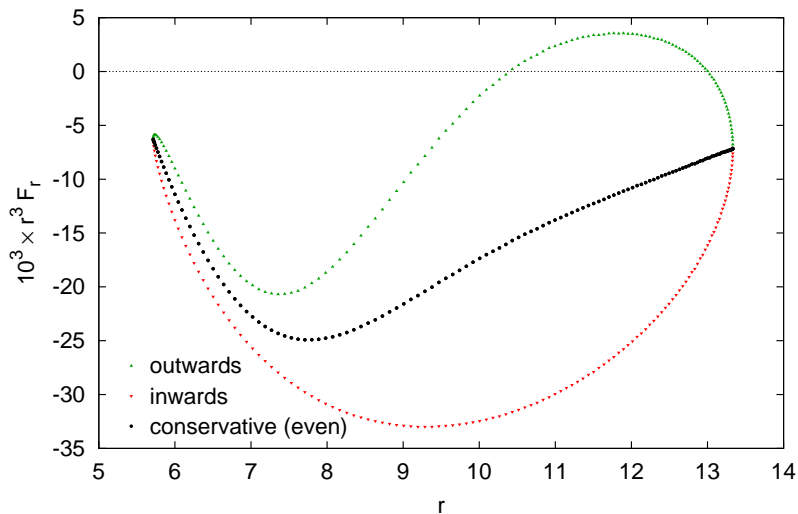
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Generic (inclined eccentric) orbits:

- **our first attempt at an effective source had ~ 20 million terms**
 \Rightarrow impractical to compile machine-generated C code
- we are starting to explore various ideas to reduce the complexity, and are optimistic we can solve this

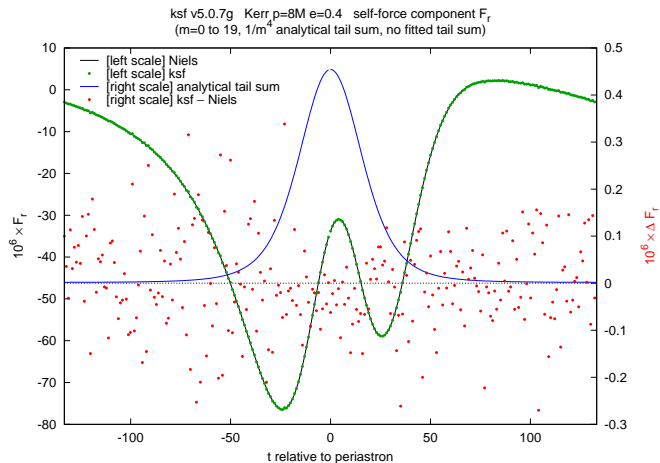
Self-force for $e = 0.4$ orbit

BH spin 0.6 orbit: $p=8M$, $e=0.4$



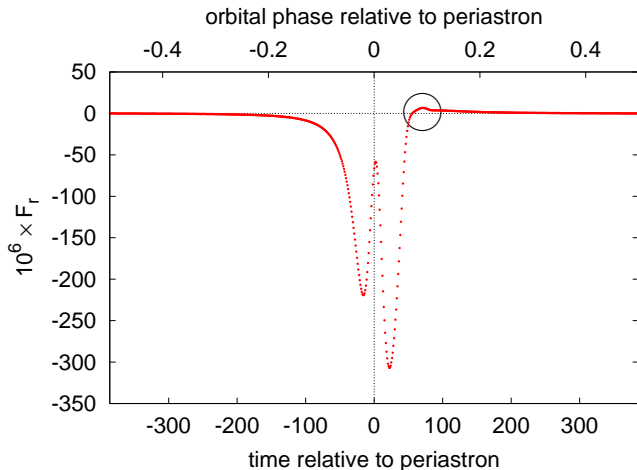
Comparison with Frequency-Domain Results

Big thanks to **Niels Warburton (Dublin)** for frequency-domain results!



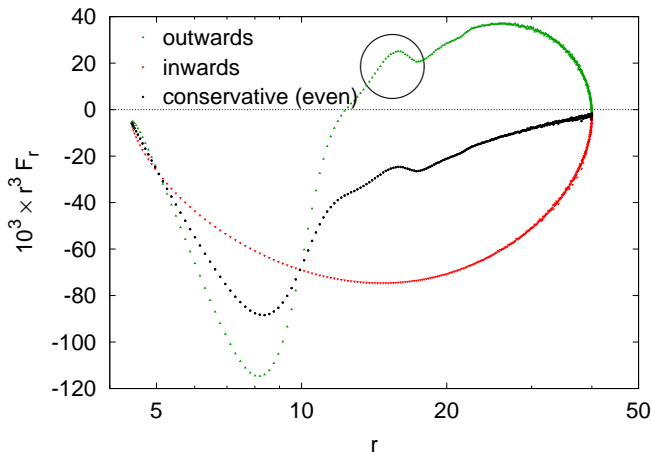
Wiggles/Ripples (1)

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Wiggles/Ripples (2)

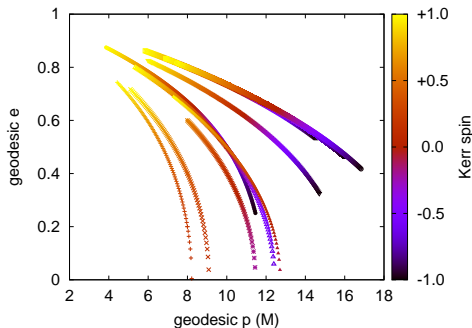
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Isofrequency Kerr Geodesics

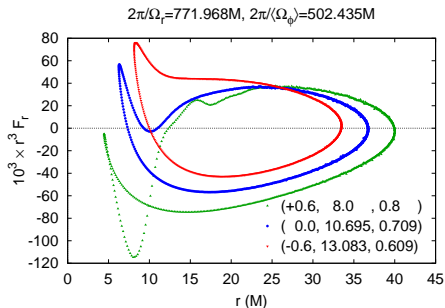
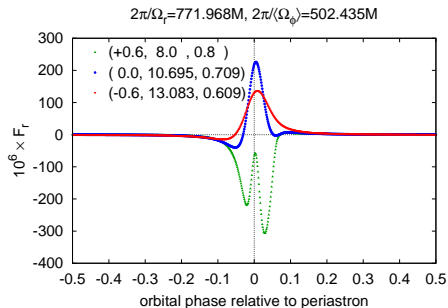
Basic idea:

- parameterize bound equatorial geodesics in Kerr spacetime by (a, p, e) where $a = \text{Kerr spin}$ and (p, e) characterize the geodesic
- since $\Omega_r = \Omega_r(a, p, e)$ and $\langle \Omega_\phi \rangle = \langle \Omega_\phi \rangle(a, p, e)$, in general we can find a 1-parameter family of (a, p, e) with the **same Ω_r and $\langle \Omega_\phi \rangle$**
- that is, these are **geodesics in different-spin Kerr spacetimes which have the same dominant GW frequencies**



Self-force for Isofrequency Kerr Geodesics

Very different self-forces \Rightarrow very different orbital evolutions \Rightarrow break degeneracies



Things that work well

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 - note that this does **not** require geodesic approximation

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 - compute $S_{\text{effective}}$ integrals via elliptic integrals
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Highly eccentric orbits:

- numerical errors & cost per unit of evolution time seem to be only weakly dependent on eccentricity
- main limits are cost due to long orbital period \Rightarrow long evolution time

Directions for further research

Things that don't (yet) work well:

- $S_{\text{effective}}$ is very expensive to calculate ($\gtrsim \frac{2}{3}$ of code's CPU time)
- dynamically adjust mesh refinement and/or worldtube around the orbit
- evolved fields only $C^2 \Rightarrow$ hard to get higher-order finite-diff convergence
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- $\mathcal{O}(\mu^2)$