

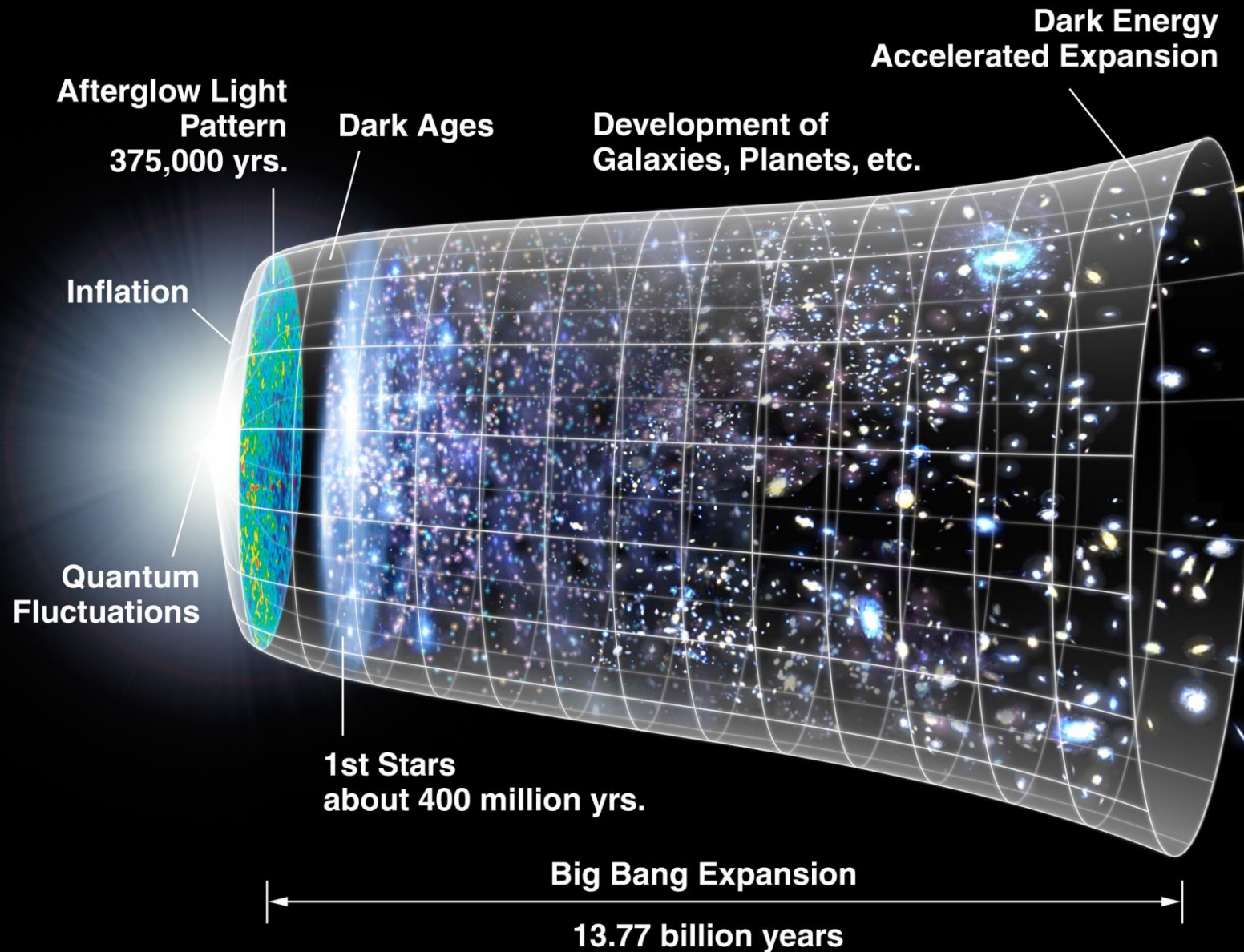
# Non-Standard Primordial Clocks from induced Mass in Alternatives to Inflation Scenarios

Zun Wang, The Hong Kong University of Science and Technology  
王尊, 香港科技大學

## Reference:

1. Xingang Chen, Mohammad Hossein Namjoo and Yi Wang, 1509.03930
2. Yi Wang, Zun Wang and Yuhang Zhu, 2007.09677

# Universe evolution



**-Expanding?  
-Contracting?**

**How Fast?**

**How to describe it?**

## Alternative scenarios of the inflation

The relative expansion of the universe is parametrized by a dimensionless scale factor  $a$ . In an expanding or contracting FLRW universe, starting from time  $t_0$  to any arbitrary time  $t$ , it should be:

$$d(t) = a(t) d_0(t_0),$$

where  $d(t)$  refers to the proper distance at  $t$  and  $d_0$  is the distance at the reference time  $t_0$ .

And we usually use conformal time  $\tau$  but not physical time  $t$  for calculation simplification which is defined as:

$$dt = a(\tau) d(\tau),$$

where the cosmological conformal time is the proper time reconciled by the fundamental observers.

# Cosmological scale factor and conformal time

## Four possibilities for alternatives scenarios

$$a = a_0 \left( \frac{t}{t_0} \right)^p = a_0 \left( \frac{\tau}{\tau_0} \right)^{p/(1-p)}$$

**fast expansion**

$$|p| > 1$$

**fast contraction**

$$0 < p \sim O(1) < 1$$

**slow contraction**

$$0 < p \ll 1$$

**slow expansion**

$$-1 \ll p < 0$$

Figure 1. Four possibilities for alternatives scenarios with power law scale factor  $a \sim t^p$ . Each coefficient  $p$  represents different alternatives of cosmic evolution.

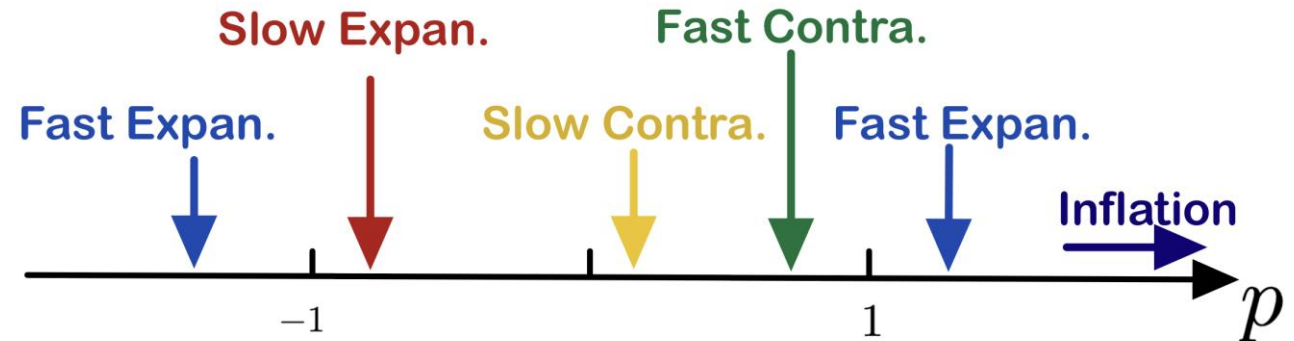


Figure 2. Axial Schematic diagram to explain four possibilities for alternatives scenarios with power law scale factor  $a \sim t^p$ .

# Alternative scenarios of the inflation

We know fluctuations as functions of scales ( $k$ ) very well that

$$k \sim -1/\tau \text{ (conformal time).}$$

Thus we know that: fluctuation  $\leftrightarrow$  conformal time

But the problem is that: fluctuation  $\leftrightarrow$  physical time

## Classical Standard Clock (Massive Oscillations) (X. Chen, 1104.1323, 1106.1635)

$$\int d\tau f(\tau) e^{-ik\tau} e^{imt},$$

Saddle point approximation:

inverse function of  $a(t)$

$$\Delta P_\zeta \sim \sin \left[ \frac{p^2}{p-1} \frac{\omega}{H_0} \left( \frac{k}{k_r} \right)^{1/p} + \text{phase} \right].$$

# Classical Standard Clock

## Reference:

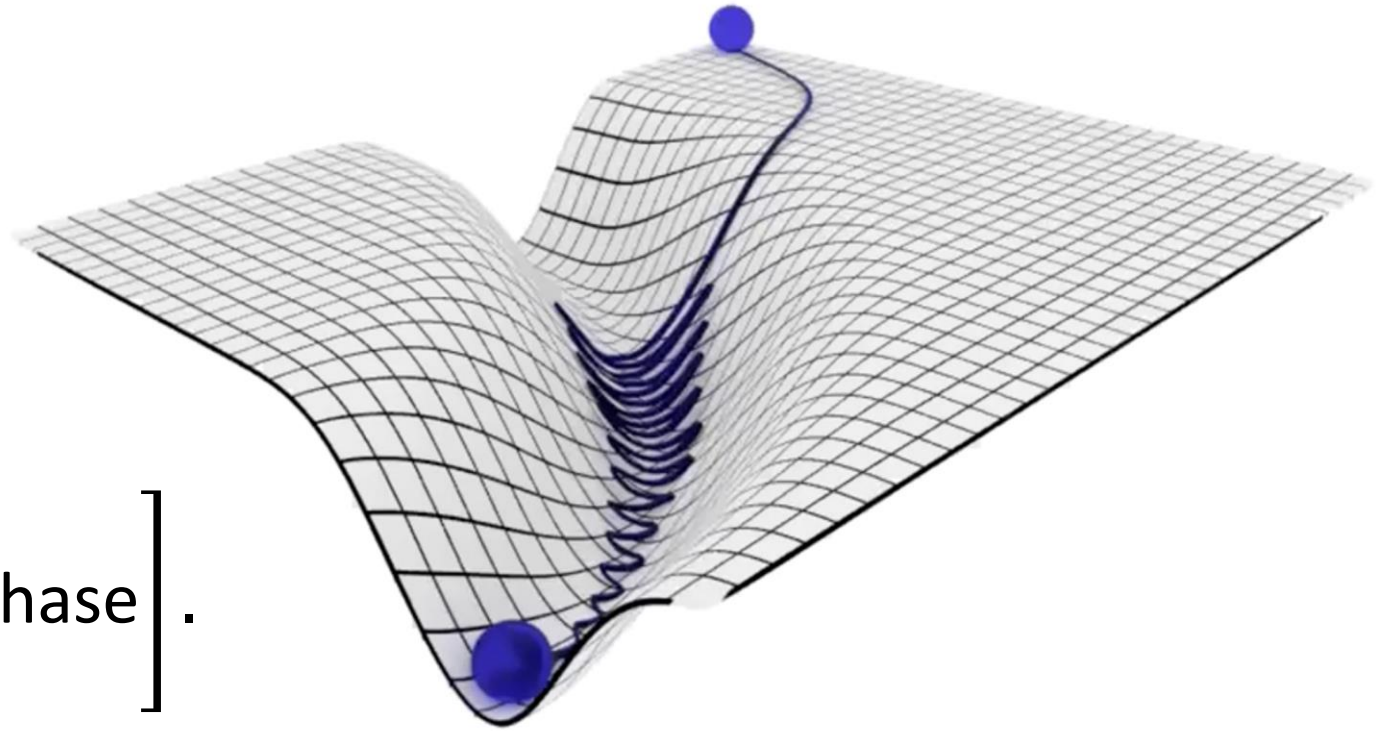
“Primordial Features as Evidence for Inflation” arxiv: 1104.1323

“Fingerprints of Primordial Universe Paradigms as Features in Density Perturbations” arxiv: 1106.1635

$$\int d\tau f(\tau) e^{-ik\tau} e^{imt},$$

Saddle point approximation:

$$\Delta P_\zeta \sim \sin \left[ \frac{p^2}{p-1} \frac{\omega}{H_0} \left( \frac{k}{k_r} \right)^{1/p} + \text{phase} \right].$$



# Non-standard Classical Clocks (Chen,1104.1323; Huang&Pi, 1610.00115 | Domnech, Rubio&Wons, 1811.08224)

Couplings:

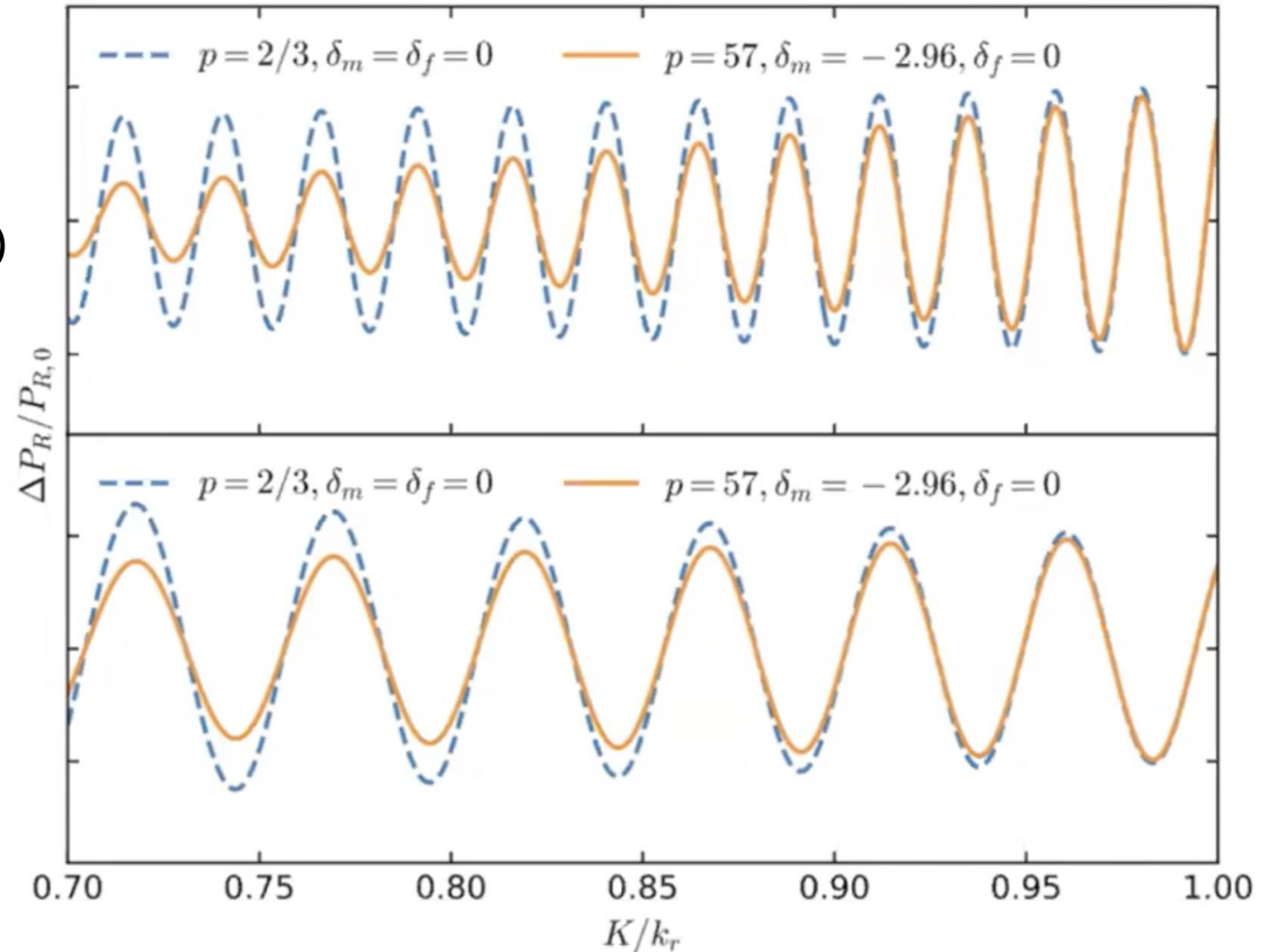
$$\frac{-L}{\sqrt{-g}} = \frac{M_p^2}{2} r - \frac{1}{2} \omega^2(\chi) g^{\mu\nu} \partial_\mu \phi \partial_\nu \phi - V(\phi)$$

$$- \frac{1}{2} f^2(\phi) g^{\mu\nu} \partial_\mu \chi \partial_\nu \chi - \frac{1}{2} m^2(\phi) \chi^2$$

where

$$\delta_f \equiv \frac{d \ln f}{dN}$$

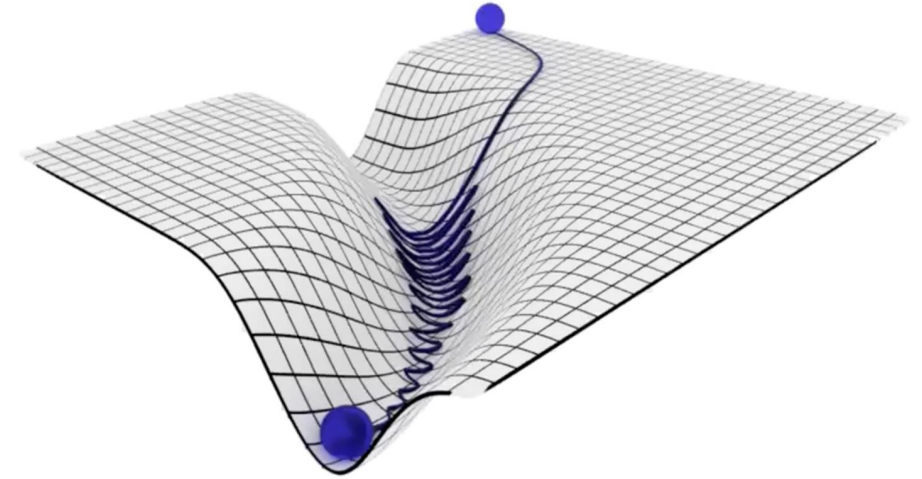
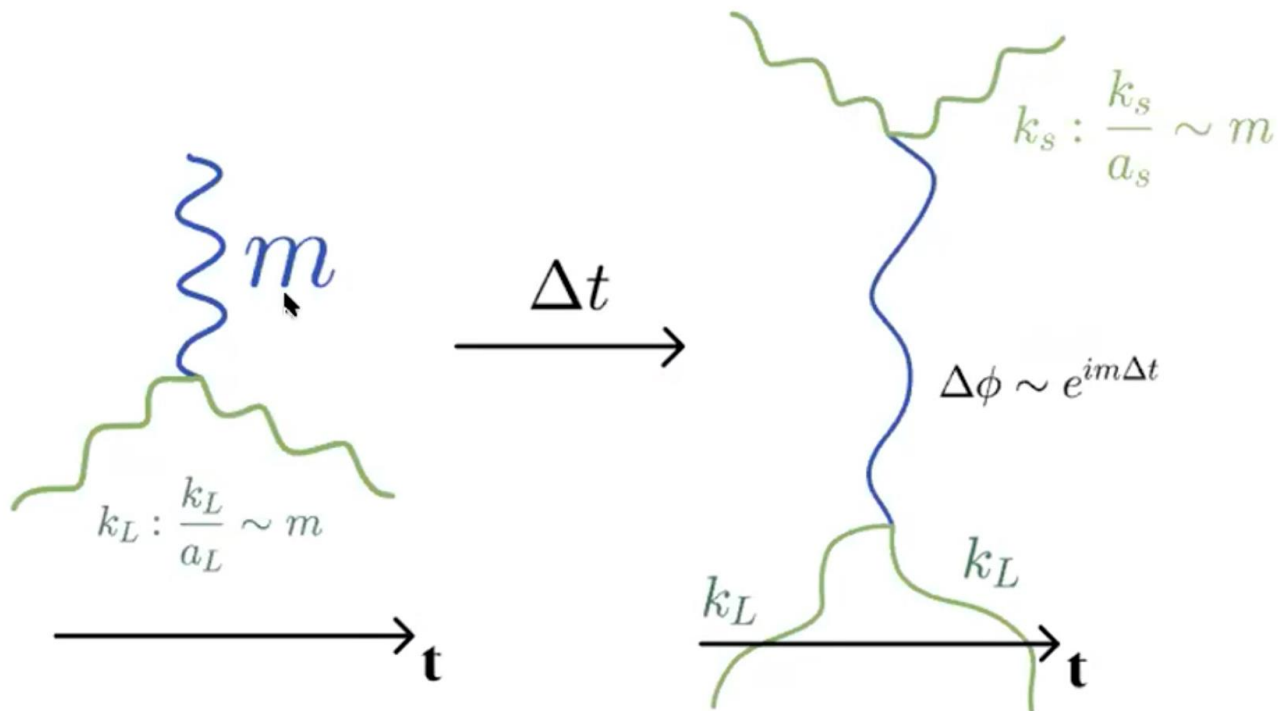
$$\delta_m \equiv \frac{d \ln m}{dN}$$



# Classical Clock to the Quantum Clock

Classical Clock Signal: Large signal, but fortune needed

Quantum Clock Signal: One single-particle needed, existence, large signal if lucky



$$\text{corr.f.} \sim \exp[im(t_s - t_L)]$$

$$\sim \exp \left[ i \frac{m}{H} (\log(k_S/m)) - \log(k_L/m) \right]$$

$$\sim \left( \frac{k_S}{k_L} \right)^{im/H} \quad (\text{inflation})$$



# Non-standard quantum clocks (Yi Wang, Zun Wang and Yuhang Zhu, 2007.09677)

1. Time-dependent effective mass from coupling to the Ricci curvature

$$\mu^2 = m^2 + \xi R = m^2 + 6\xi \frac{p(2p-1)}{t^2} \rightarrow \mu \sim t^{-1} \rightarrow H$$

2. Time-dependent effective mass from nonlinear corrections ( $V = \frac{1}{4}\lambda\phi^4$ )

$$\dot{\sigma} = -\frac{\lambda}{3p} t\sigma^3 + \frac{\sqrt{\Lambda}}{3p} t^{(-3/2)}\eta(x, t) \Rightarrow \mu \sim t^{-1} \rightarrow H$$

## Standard/ Non-standard Quantum Clock Signal

$$m \sim H \sim t^{-1}: \quad S^{\text{clock}} \sim f(p, \xi, k_3) \left(\frac{2k_1}{k_3}\right)^{\frac{1+p}{2(1-p)}} \sin \left\{ S \log \left[ \left(\frac{2k_1}{k_3}\right)^{-1} \right] + \phi(p, \xi, k_3) \right\}$$

# Non-standard quantum clocks (Yi Wang, Zun Wang and Yuhang Zhu, 2007.09677)

1. Time-dependent effective mass from coupling to the Ricci curvature

$$\mu^2 = m^2 + \xi R = m^2 + 6\xi \frac{p(2p-1)}{t^2} \rightarrow \mu \sim t^{-1} \rightarrow H$$

2. Time-dependent effective mass from nonlinear corrections ( $V = \frac{1}{4}\lambda\phi^4$ )

$$\dot{\sigma} = -\frac{\lambda}{3p} t\sigma^3 + \frac{\sqrt{\Lambda}}{3p} t^{(-3/2)}\eta(x, t) \Rightarrow \mu \sim t^{-1} \rightarrow H$$

## Standard/ Non-standard Quantum Clock Signal

$$m \sim H \sim t^{-1}: \quad S^{\text{clock}} \sim f(p, \xi, k_3) \left(\frac{2k_1}{k_3}\right)^{\frac{1+p}{2(1-p)}} \sin \left\{ S \log \left[ \left(\frac{2k_1}{k_3}\right)^{-1} \right] + \phi(p, \xi, k_3) \right\}$$

# Non-standard quantum clocks (Yi Wang, Zun Wang and Yuhang Zhu, 2007.09677)

## Standard/ Non-standard Quantum Clock Signal

$$m \sim H \sim t^{-1}: \quad S^{\text{clock}} \sim f(p, \xi, k_3) \left( \frac{2k_1}{k_3} \right)^{\frac{1+p}{2(1-p)}} \sin \left\{ S \log \left[ \left( \frac{2k_1}{k_3} \right)^{-1} \right] + \phi(p, \xi, k_3) \right\}$$

### Inflation

$$|p| > 1, \frac{1+p}{2(1-p)} < 0 \Rightarrow \text{decreasing with } k_1$$

### Alternative scenarios

$$|p| < 1, \frac{1+p}{2(1-p)} > 0 \Rightarrow \text{increasing with } k_1$$

## Conclusion

If we find the quantum clock signal from the power spectrum:

- $\sin(\sim k^{1/p})$  with  $|p| < 1$ : alternative scenarios;
- $\sin(\sim \log(\dots k))$ :
  - correlation function increases with  $k$ : alternative scenarios
  - correlation function decreases with  $k$ : inflation

(other nice non-standard clocks caused by others can be studied...)

# Reference

- [1] A.H. Guth, The Inflationary Universe: A Possible Solution to the Horizon and Flatness Problems, Phys. Rev. D 23 (1981) 347.
- [2] A.D. Linde, A New Inflationary Universe Scenario: A Possible Solution of the Horizon, Flatness, Homogeneity, Isotropy and Primordial Monopole Problems, Adv. Ser. Astrophys. Cosmol. 3 (1987) 149.
- [3] A. Albrecht and P.J. Steinhardt, Cosmology for Grand Unified Theories with Radiatively Induced Symmetry Breaking, Phys. Rev. Lett. 48 (1982) 1220.
- [4] A.A. Starobinsky, A New Type of Isotropic Cosmological Models Without Singularity, Adv. Ser. Astrophys. Cosmol. 3 (1987) 130.
- [5] J. Khoury, B.A. Ovrut, P.J. Steinhardt and N. Turok, The Ekpyrotic universe: Colliding branes and the origin of the hot big bang, Phys. Rev. D 64 (2001) 123522 [hep-th/0103239].
- [6] D. Wands, Duality invariance of cosmological perturbation spectra, Phys. Rev. D 60 (1999) 023507 [gr-qc/9809062].
- [7] F. Finelli and R. Brandenberger, On the generation of a scale invariant spectrum of adiabatic fluctuations in cosmological models with a contracting phase, Phys. Rev. D 65 (2002) 103522 [hep-th/0112249].
- [8] Y.-S. Piao and E. Zhou, Nearly scale invariant spectrum of adiabatic fluctuations may be from a very slowly expanding phase of the universe, Phys. Rev. D 68 (2003) 083515 [hep-th/0308080].
- [9] R.H. Brandenberger and C. Vafa, Superstrings in the Early Universe, Nucl. Phys. B 316 (1989) 391.

Thank you