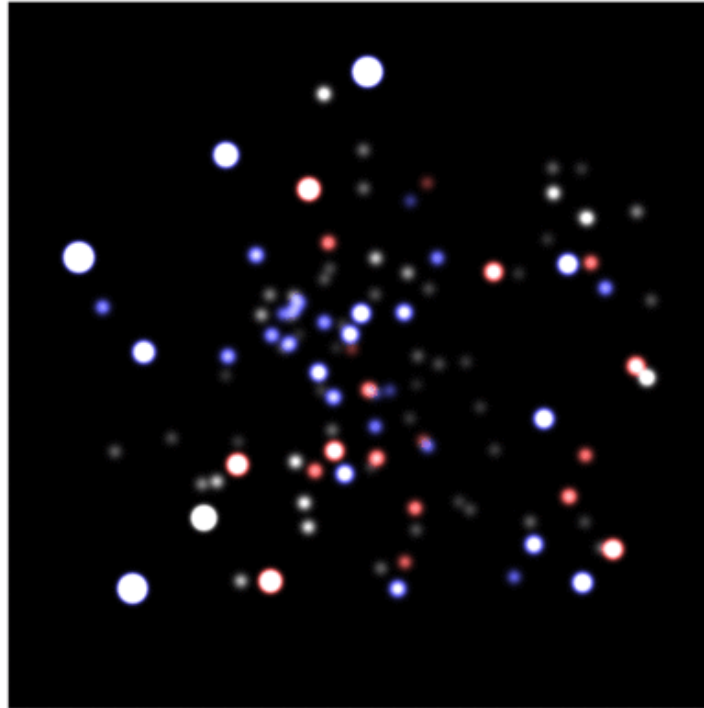


# Energy extraction from black holes

Kyoto 2012 | March 28, 2012 | Kyoto, Japan



Credit: ESO/MPE (2010)

**Vitor Cardoso** (CENTRA/IST & Olemiss)

**Cardoso et al, PRL107: 241101 (2011)**

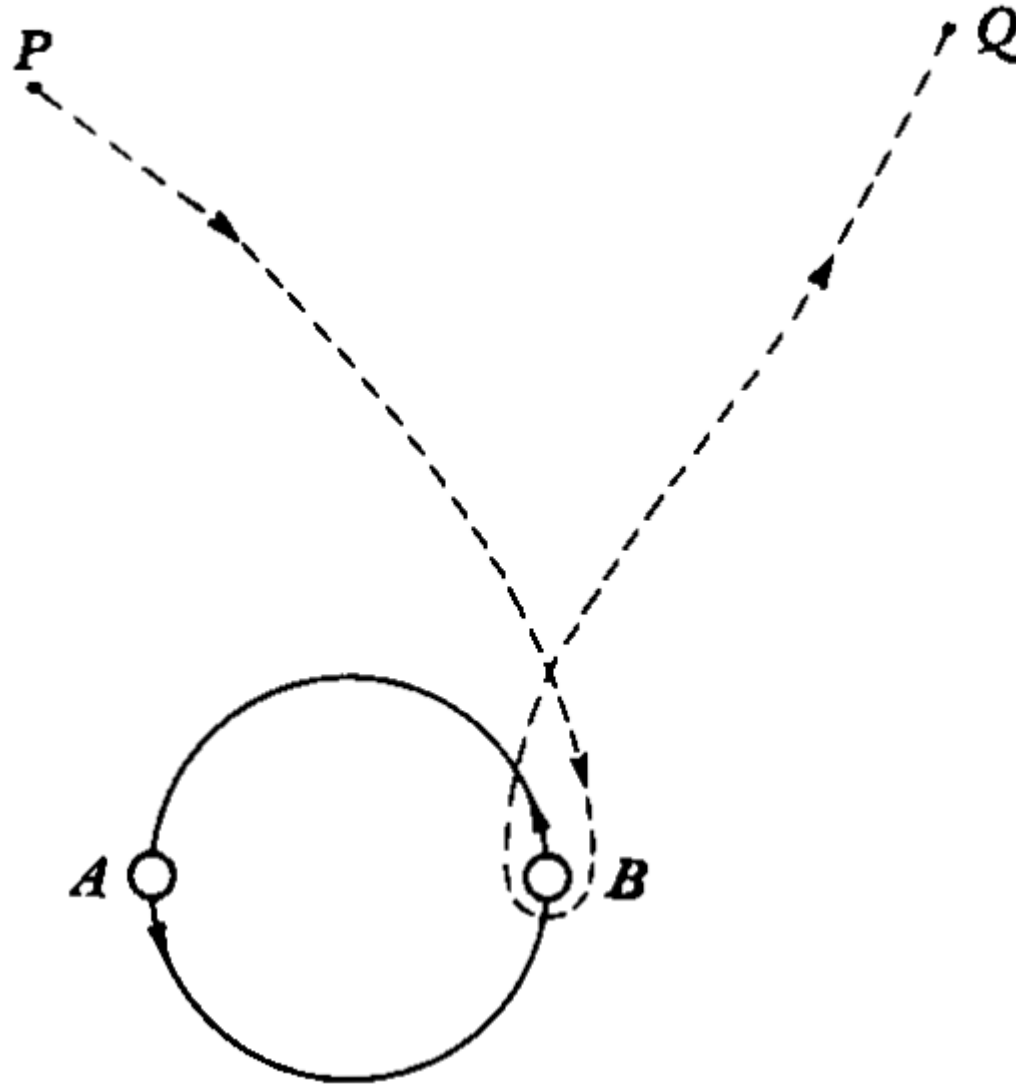
**Yunes, Pani, Cardoso, arXiv:1112.3351**

**Cardoso & Pani, unpublished; Brito et al (2012)**

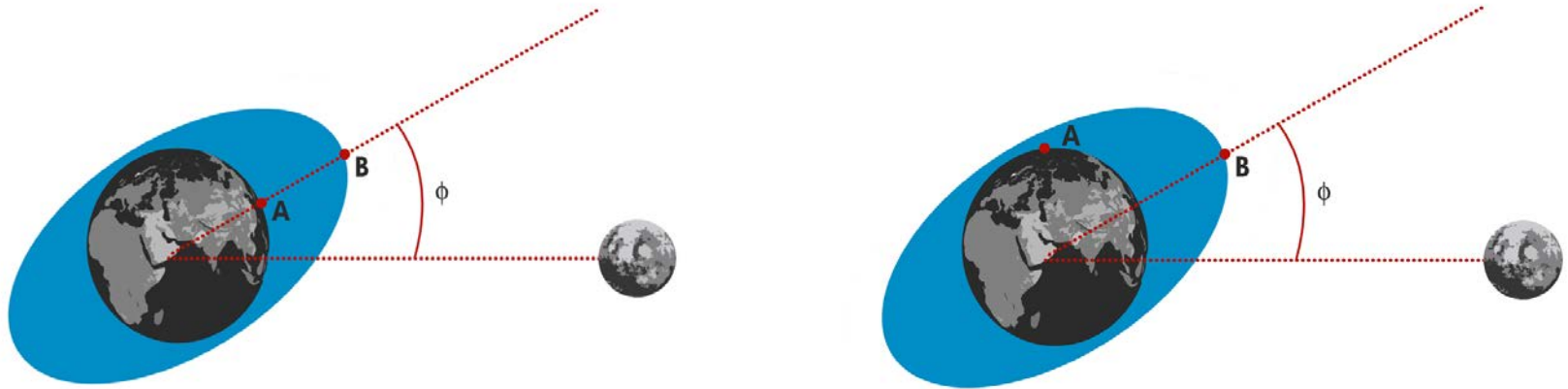


Credit: ESO/MPE/M.Schartmann (2011)

## Orbital energy: slingshot effect



# Rotational energy: tidal acceleration



Earth-moon: 0.002s/cent

4cm/yr

$$\mu = \frac{\kappa}{2} m_p \left( \frac{R}{r_0} \right)^3$$

$$\phi = (\Omega_H - \Omega)\tau$$

$$\dot{E}_{\text{orbital}} = 3G\kappa m_p^2 \frac{R^5}{r_0^6} \Omega (\Omega_H - \Omega) \tau$$

## Tidal acceleration is in general impossible for BHs!

$$\dot{E}_H \sim \frac{G^7}{c^{13}} \frac{M^6 m_p^2}{r_0^6} \Omega (\Omega - \Omega_H)$$

$$\dot{E}_\infty \sim \frac{32}{5} \frac{G^4}{c^5} \frac{M^3 m_p^2}{r_0^5}$$

$$\frac{\dot{E}_H}{\dot{E}_\infty} = \left( \frac{GM}{c^2 r_0} \right)^3 \frac{r_0 \Omega}{c} \left( \frac{r_0 \Omega - r_0 \Omega_H}{c} \right) \sim (v/c)^8$$

*Press & Teukolsky, Nature (1973)*

## **Resonant excitation of BH modes? Resonant tides?**

Massless modes of BHs are localized around *the light ring...*

*...and the null geodesic lies inside the ISCO*

## ***(massive) scalar fields***

**Interesting as effective description**

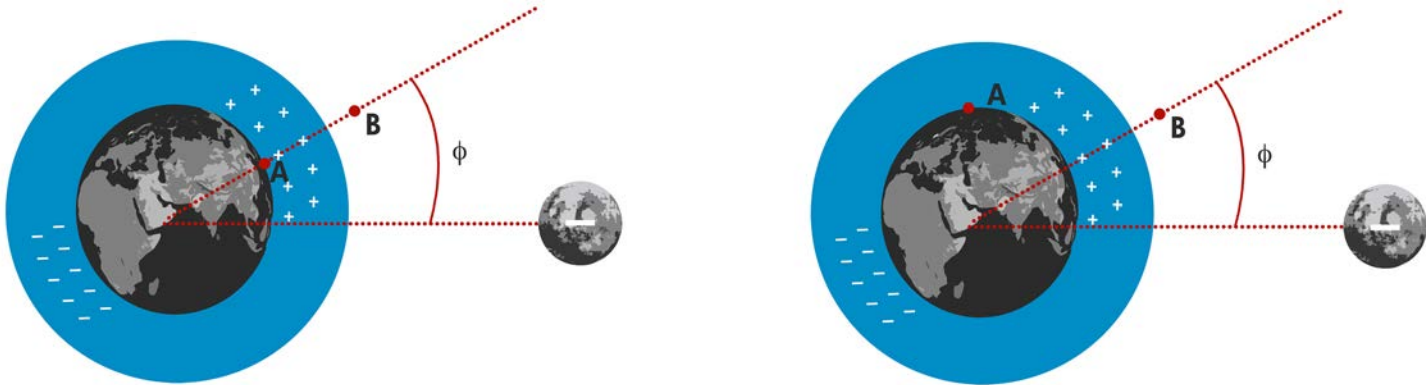
**Proxy for more complex interactions**

**Arise as interesting alternatives of GR**

Brans-Dicke or generic scalar-tensor theories; quadratic  $f(R)$

**Axiverse scenarios (moduli and coupling constants  
in QCD, Peccei-Quinn mechanism QCD, etc)**

**We are in the Yukawa Institute**



$$\sigma_{\text{pol}} = 3\epsilon_0 \left( \frac{\epsilon_r - 1}{2\epsilon_r + 1} \right) E_0 \cos \theta$$

$$p = 4\pi\epsilon_0 \left( \frac{\epsilon_r - 1}{2\epsilon_r + 1} \right) R^3 E_0$$

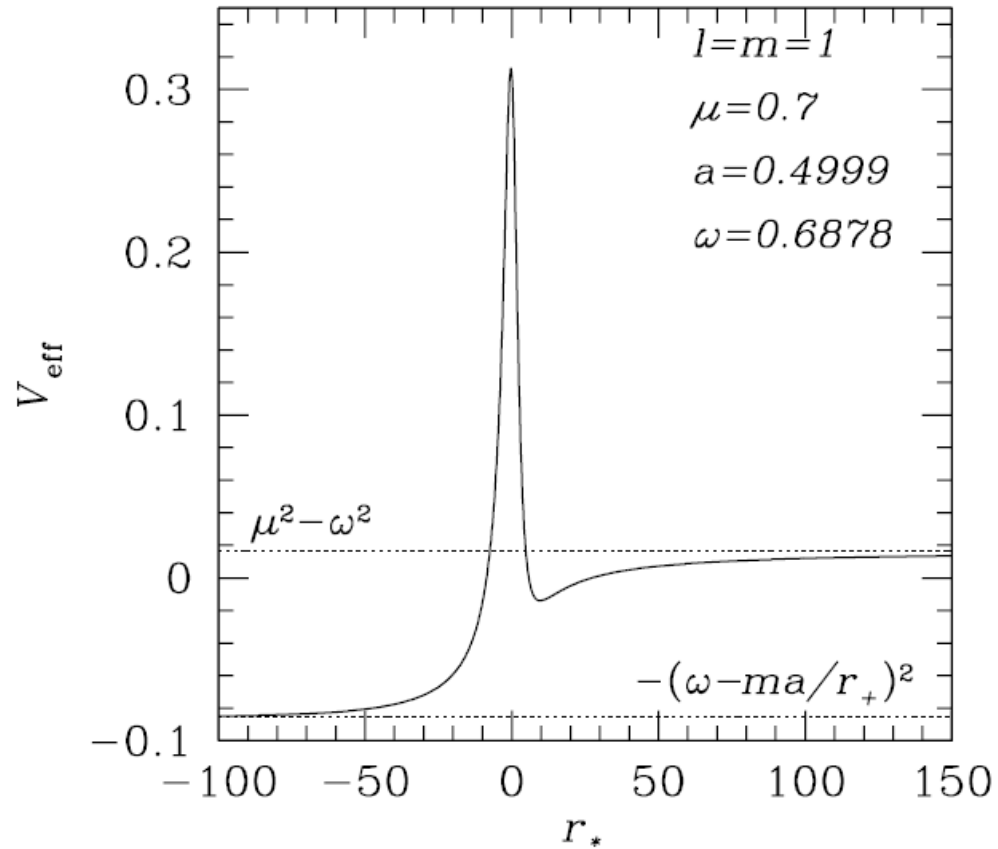
$$\dot{E}_{\text{orbital}} = \left( \frac{\epsilon_r - 1}{2\epsilon_r + 1} \right) \frac{q_p^2 R^3 \tau}{r_0^4} \Omega (\Omega_H - \Omega)$$



$$\omega_{\text{res}}^2 = \mu_s^2 - \mu_s^2 \left( \frac{\mu_s M}{l + 1 + n} \right)^2$$

$$\omega_I = \mu_s \frac{(\mu_s M)^8}{24} (a/M - 2\mu_s r_+)$$

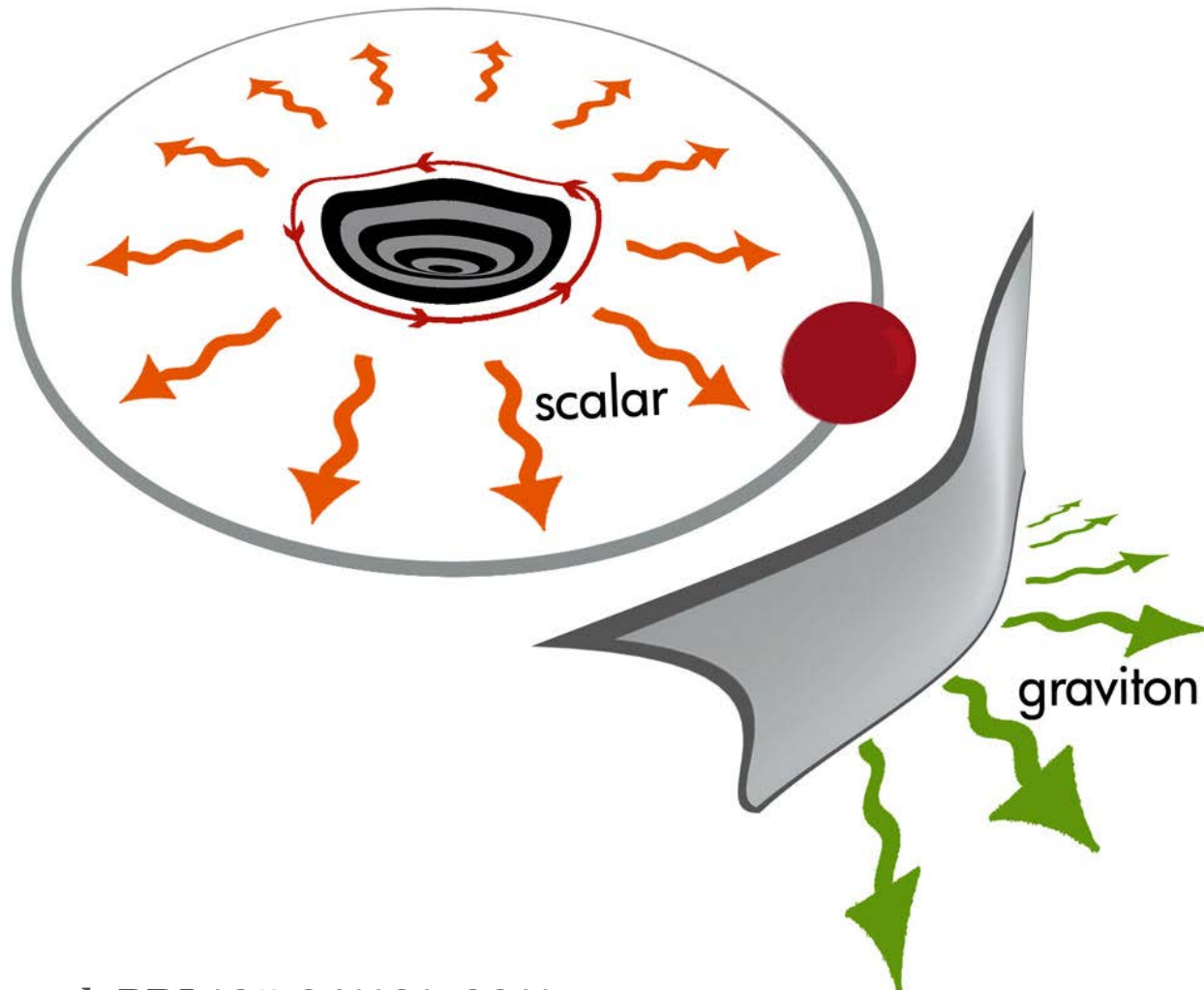
# Massive fields



**Massive fields around Kerr are unstable**

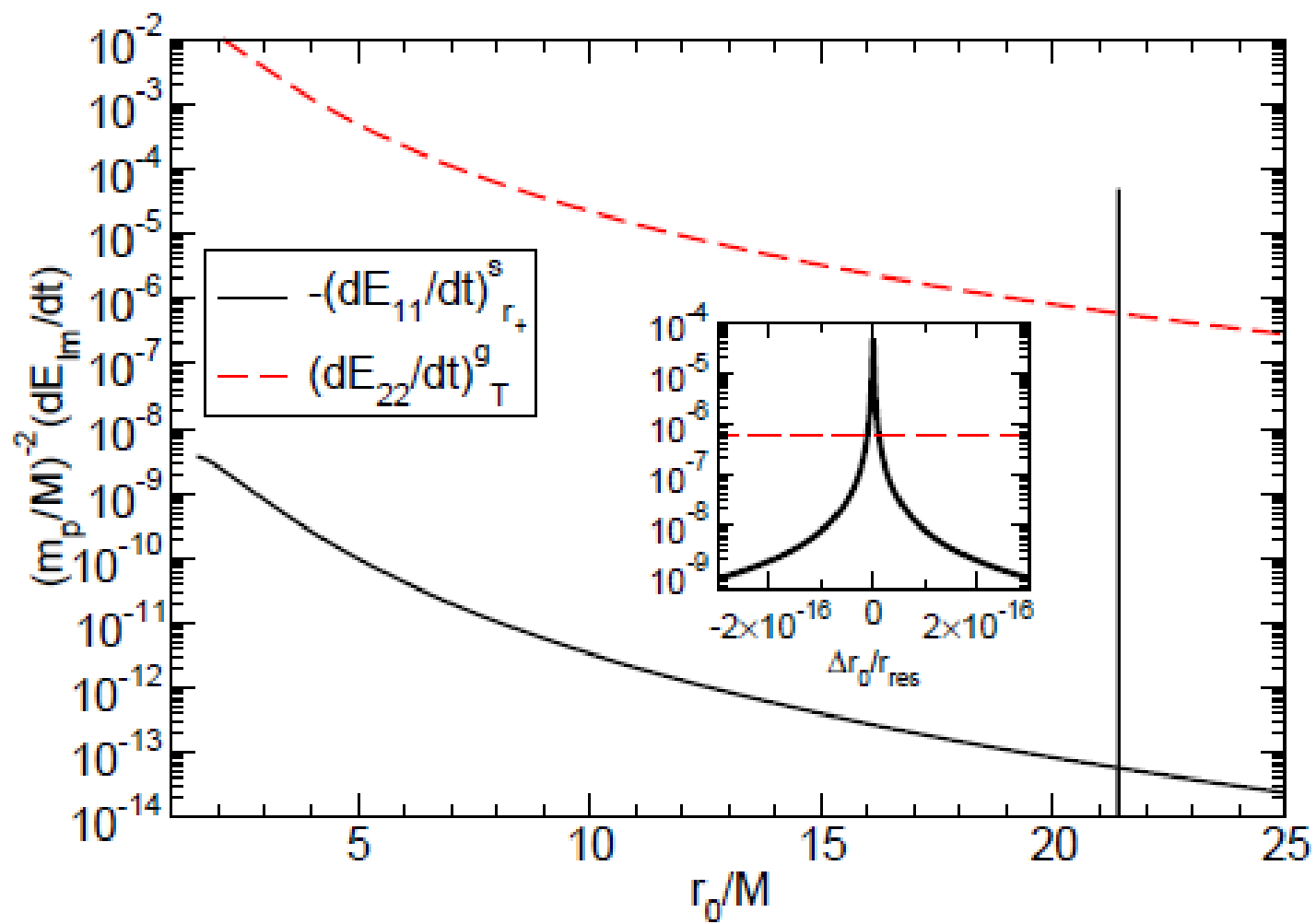
*(Damour et al '76; Detweiler, '80; Cardoso & Yoshida '05; Dolan '07)*

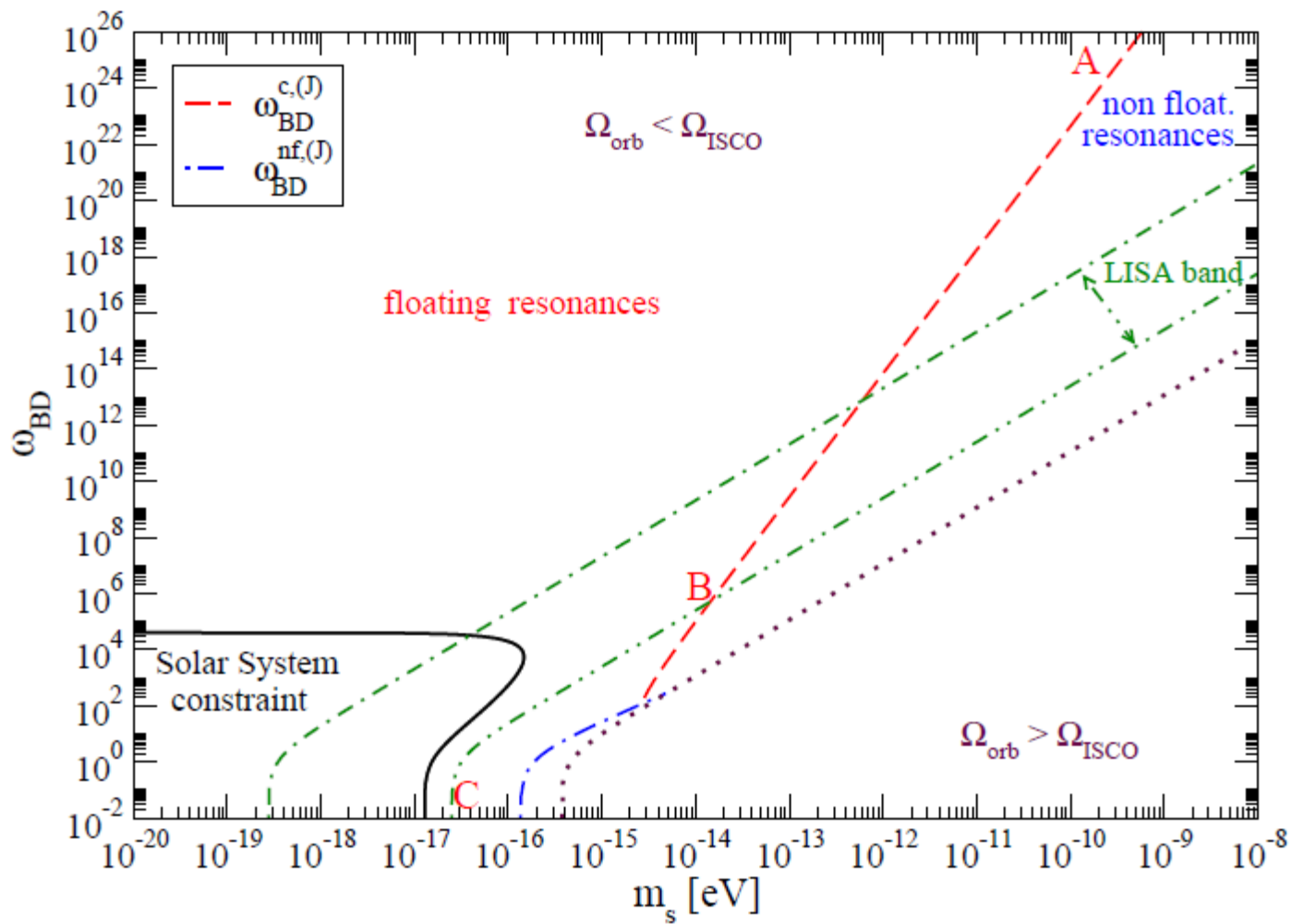
## Floating orbits



$$[\square - \mu_s^2] \varphi = \alpha \mathcal{T}$$

$$\dot{E}_{r_+}^{s,\text{peak}} \sim - \frac{3\alpha^2 \sqrt{\frac{r_0}{M}} m_p^2 M}{16\pi r_+ (M^2 - a^2) \left( \frac{a}{2r_+} - \left( \frac{M}{r_0} \right)^{3/2} \right) \mathcal{F}}$$





## Tidal acceleration in higher dimensions

$$\frac{\dot{E}_H}{\dot{E}_\infty} \sim (v/c)^{\frac{-(D-5)(D+1)}{D-3}} \quad s = 2$$

$$\frac{\dot{E}_H}{\dot{E}_\infty} \sim (v/c)^{\frac{-(D-5)(D-1)}{D-3}} \quad s = 0$$

$D > 5$  particles do not merge, tidal effects are too large?

*(scalar field case: yes - Brito et al, in progress)*

Circular orbits are unstable, on much smaller timescale...what happens?!

Tidal acceleration is equivalent to superradiance in BH physics

In absence of other dissipative effects tidal dissipation leads to floating



Floating orbits can be instrumental to constrain or prove existence of massive scalars coupled to matter... still a lot to do:

- Equal-mass case, what happens to floating?
- Eccentricity OK, what about other sources of noise? Higher multipoles?
- Spinning companion, is floating enhanced? Does it still require massive fields?
- Tidal acceleration requires dissipation (EH). Can it occur for spinning objects without horizon? In principle no, but Blandford-Znajek seems to, or does it? (see Ruiz et al, arXiv:1203.4125)
- Higher dimensional spacetime: tidal dissipation is dominant mechanism. Consequence for mergers?





**Thank you**