太陽ニュートリノから示唆される非標準相互作用のT2HKKとDUNEによる検証可能性

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Ghosh & OY, arXiv:1709.08264

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1. Introduction

Framework of 3 flavor ν oscillation

Mixing matrix

$$\begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \end{pmatrix} = \begin{pmatrix} U_{e1} & U_{e2} & U_{e3} \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} \end{pmatrix} \begin{pmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \end{pmatrix}$$

Functions of mixing angles θ_{12}, θ_{23}, θ_{13}, and CP phase δ

All 3 mixing angles have been measured

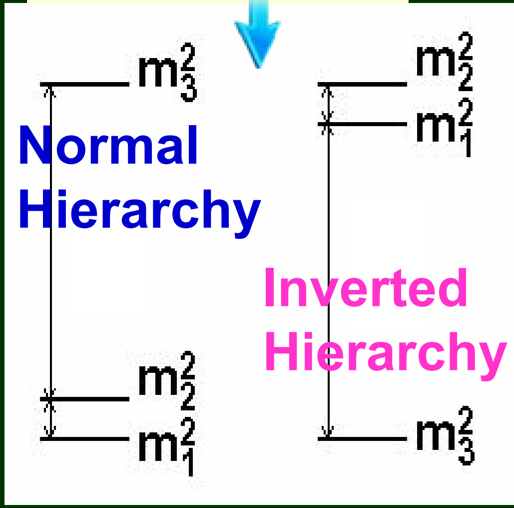
ν_{solar}+KamLAND (reactor) → $\theta_{12} \cong \frac{\pi}{6}, \Delta m^2_{21} \cong 8\times 10^{-5}\,\mathrm{eV}^2$

ν_{atm}, K2K,T2K,MINOS,Nova (accelerators) → $\theta_{23} \cong \frac{\pi}{4}, |\Delta m^2_{32}| \cong 2.5\times 10^{-3}\,\mathrm{eV}^2$

DCHOOZ+Daya Bay+Reno (reactors), T2K+MINOS+Nova → $\theta_{13} \cong \pi/20$

Next task is to measure sign(Δm^2_{31}), $\pi/4-\theta_{23}$ and δ

Both hierarchy patterns are allowed



Proposed experiments

- T2HK(JP, JPARC-->HK) L=295km, E~0.6GeV
- T2HHK(JP, JPARC-->Korea) L=1100km, E~1GeV
- DUNE (US, FNAL-->Homestake, SD) , L=1300km, E~2GeV

$$\overset{(-)}{\nu}_\mu \to \overset{(-)}{\nu}_\mu + \overset{(-)}{\nu}_\mu \to \overset{(-)}{\nu}_e$$

These experiments are expected to measure sign(Δm^2_{31}), $\pi/4-\theta_{23}$ and δ

Motivation for research on New Physics

High precision measurements of ν oscillation in future experiments can be used to probe physics beyond SM by looking at deviation from SM+m_ν (like at B factories).
→ Research on New Physics is important.

List of New Physics discussed in ν phenomenology

Scenario beyond SM+m_ν	Experimental indication ?	Phenomenological constraints on the magnitude of the effects
Light sterile ν	Maybe	O(10%)
NSI at production / detection	×	O(1%)
NSI in propagation	Maybe	e-τ: O(100%) Others: O(1%)
Unitarity violation due to heavy particles	×	O(0.1%)

NSI: discussed in this talk

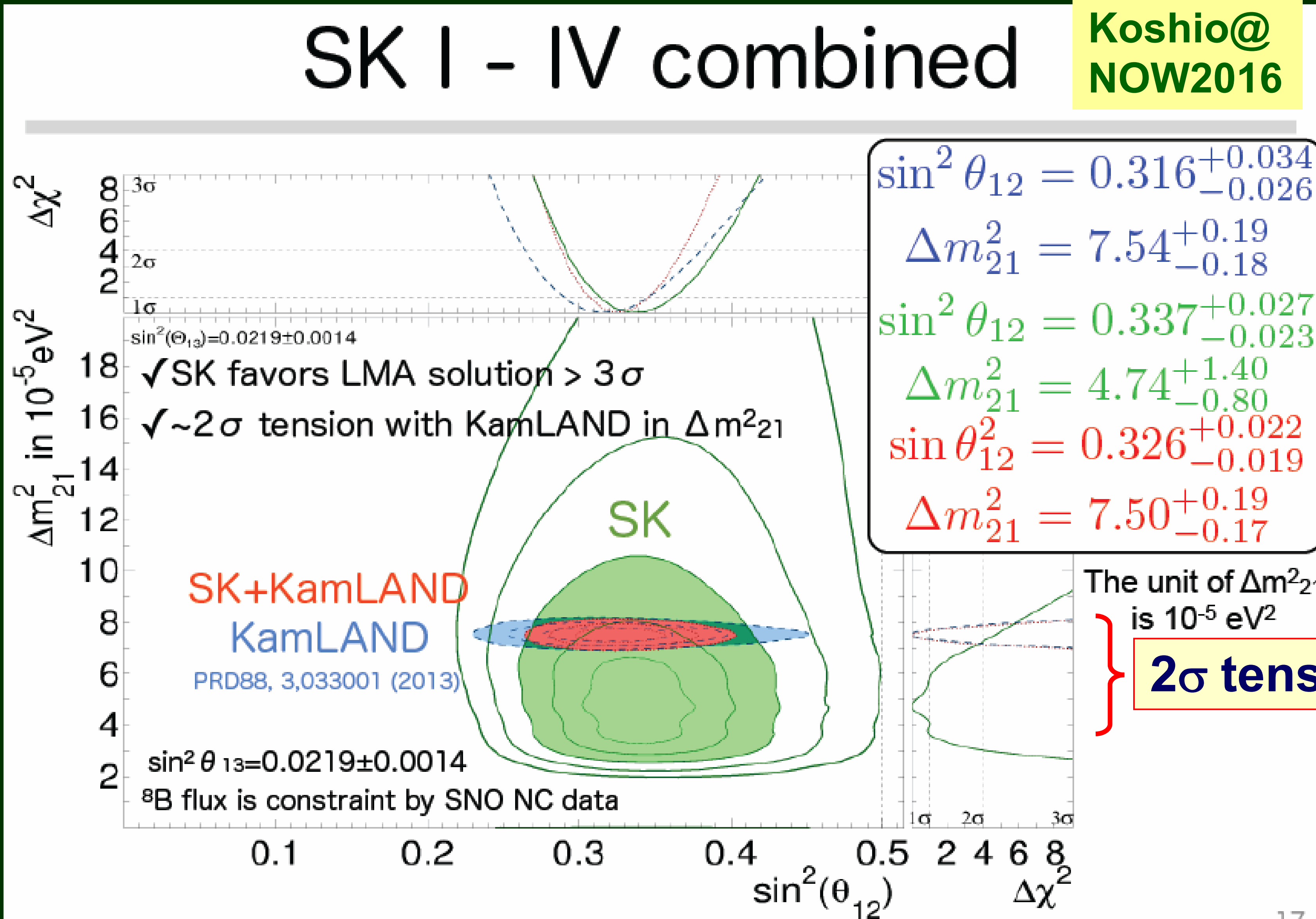
In the mean time we have had some possible tensions among the data within the standard oscillation scenario:

- ν_{solar} - KamLAND: Δm^2_{21} → **NSI** or sterile ν ($\Delta m^2 = O(10^{-5})\,eV^2$)
- MINOS - T2K: θ_{23} → Decoherence??
- LSND-MiniBooNE anomaly, Reactor anomaly, Gallium anomaly → sterile ν ($\Delta m^2 = O(1)\,eV^2$)

NSI: motivation to this talk

sterile ν : not directly related to this talk

Tension between Δm^2_{21}(solar) & Δm^2_{21}(KamLAND)

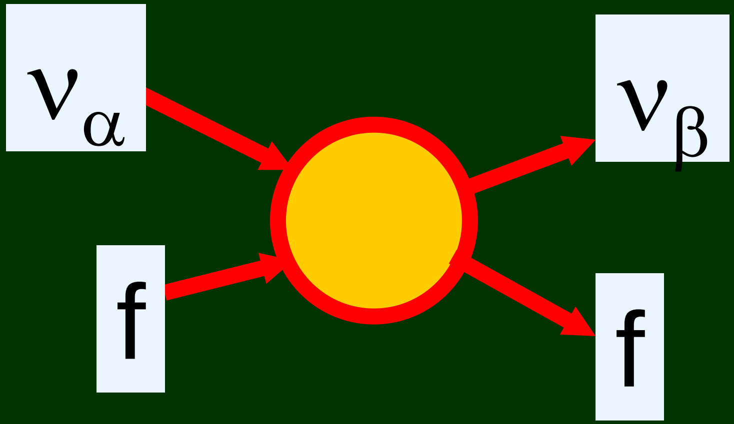


Koshio@ NOW2016

The unit of Δm^2_{21} is 10^{-5} eV^2

2σ tension

2. Nonstandard Interaction in propagation



Phenomenological New Physics considered in this talk: 4-fermi Non Standard Interactions:

$$\mathcal{L}_{eff} = G_{NP}^{\alpha\beta}\, \bar{\nu}_\alpha \gamma^\mu \nu_\beta\, \bar{f} \gamma_\mu f'$$

neutral current non-standard interaction

f = e, u or d

Modification of matter effect

$$i\frac{d}{dt}\begin{pmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{pmatrix} = \left[U\,\mathrm{diag}\left(E_1, E_2, E_3\right)U^{-1} + A\begin{pmatrix}1+\epsilon_{ee} & \epsilon_{e\mu} & \epsilon_{e\tau}\\ \epsilon_{\mu e} & \epsilon_{\mu\mu} & \epsilon_{\mu\tau}\\ \epsilon_{\tau e} & \epsilon_{\tau\mu} & \epsilon_{\tau\tau}\end{pmatrix}\right]\begin{pmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{pmatrix}$$

$A \equiv \sqrt{2} G_F N_e$ $\quad N_e \equiv$ electron density

NP

● Constraints on $\varepsilon_{\alpha\beta}$ from non-oscillation experiments

Davidson et al., JHEP 0303:011,2003; Berezhiani, Rossi, PLB535 ('02) 207; Barranco et al., PRD73 ('06) 113001; Barranco et al., arXiv:0711.0698

Biggio et al., JHEP 0908, 090 (2009)

Constraints are weak

$$\begin{pmatrix} |\epsilon_{ee}| \lesssim 4 \times 10^{0} & |\epsilon_{e\mu}| \lesssim 3 \times 10^{-1} & |\epsilon_{e\tau}| \lesssim 3 \times 10^{0} \\ & |\epsilon_{\mu\mu}| \lesssim 7 \times 10^{-2} & |\epsilon_{\mu\tau}| \lesssim 3 \times 10^{-1} \\ & & |\epsilon_{\tau\tau}| \lesssim 2 \times 10^{1} \end{pmatrix}$$

Constraints from other experiments:

$|\varepsilon_{\alpha\mu}| \ll 1$ ($\alpha = e, \mu, \tau$) -> (ε_{ee}, $\varepsilon_{e\tau}$, $\varepsilon_{\tau\tau}$)

● NSI for solar ν: $\varepsilon_{\alpha\beta}$ vs $(\varepsilon_D, \varepsilon_N)$

Gonzalez-Garcia, Maltoni, JHEP 1309 (2013) 152

In solar ν analysis, $\Delta m_{31}{}^2$ -> infinity, H -> H^{eff}

$$H^{\rm eff} = \frac{\Delta m^2_{21}}{4E}\begin{pmatrix} -\cos 2\theta_{12} & \sin 2\theta_{12} \\ \sin 2\theta_{12} & \cos 2\theta_{12}\end{pmatrix}$$

$$+\begin{pmatrix} c_{13}^2 A & 0 \\ 0 & 0\end{pmatrix} + A\sum_{f=e,u,d}\frac{N_f}{N_e}\begin{pmatrix} -\epsilon_D^f & \epsilon_N^f \\ \epsilon_N^{f*} & \epsilon_D^f\end{pmatrix}$$

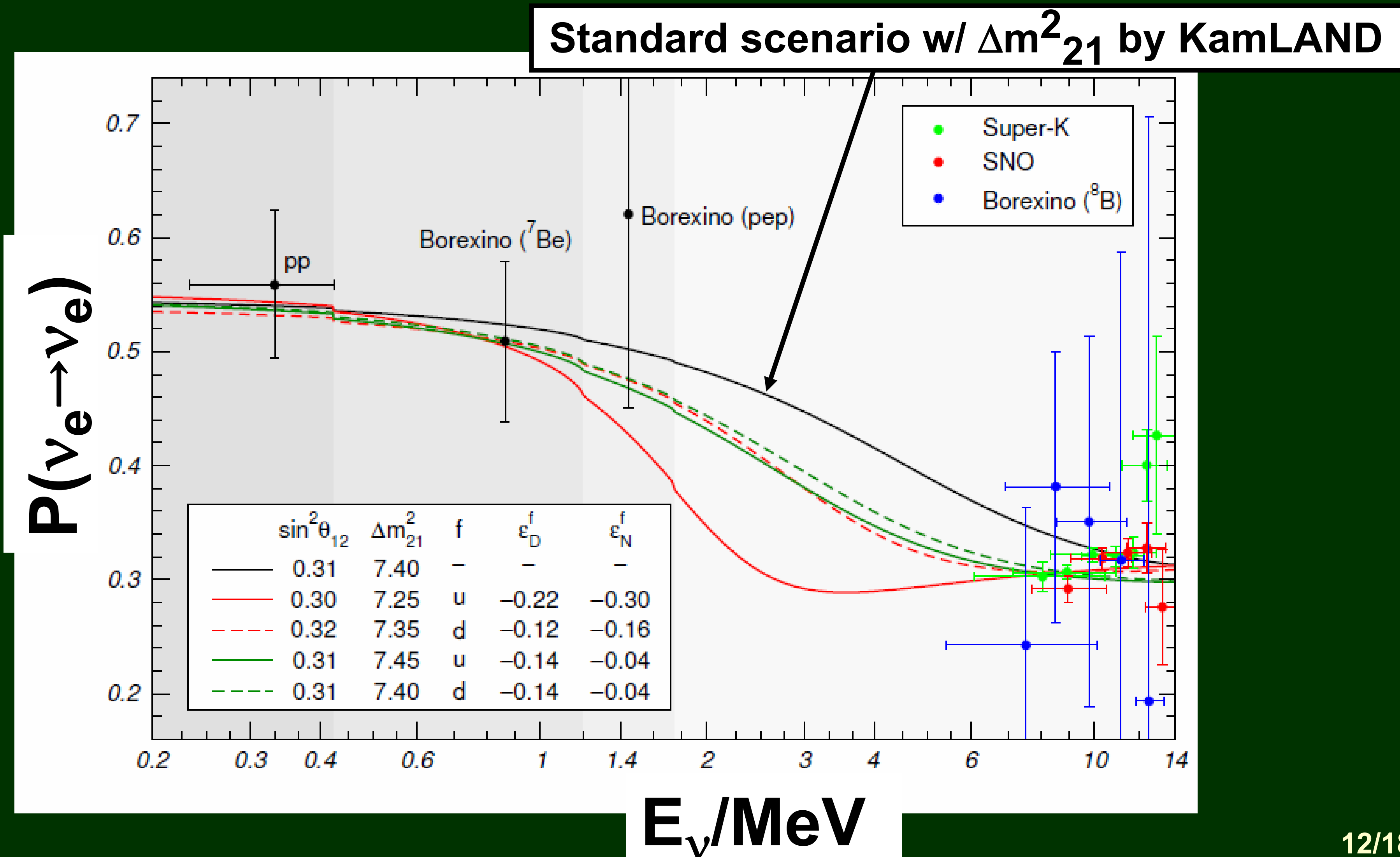
$$\begin{aligned}\epsilon_D^f &= c_{13}s_{13}{\rm Re}\left[e^{i\delta_{\rm CP}}\left(s_{23}\epsilon^f_{e\mu}+c_{23}\epsilon^f_{e\tau}\right)\right]-\left(1+s^2_{13}\right)c_{23}s_{23}{\rm Re}\left[\epsilon^f_{\mu\tau}\right] \\ &\quad -\frac{c^2_{13}}{2}\left(\epsilon^f_{ee}-\epsilon^f_{\mu\mu}\right)+\frac{s^2_{23}-s^2_{13}c^2_{23}}{2}\left(\epsilon^f_{\tau\tau}-\epsilon^f_{\mu\mu}\right) \\ \epsilon_N^f &= c_{13}\left(c_{23}\epsilon^f_{e\mu}-s_{23}\epsilon^f_{e\tau}\right)+s_{13}e^{-i\delta_{\rm CP}}\left[s^2_{23}\epsilon^f_{\mu\tau}-c^2_{23}\epsilon^{f*}_{\mu\tau}+c_{23}s_{23}\left(\epsilon^f_{\tau\tau}-\epsilon^f_{\mu\mu}\right)\right]\end{aligned}$$

f = e, u or d

ε^f_{ee}, $|\varepsilon^f_{e\tau}|$, $\varepsilon^f_{\tau\tau}$ have to be solved from $(\varepsilon^f_D, \varepsilon^f_N)$

Tension between solar ν & KamLAND data comes from little observation of upturn by SK & SNO

Gonzalez-Garcia, Maltoni, JHEP 1309 (2013) 152

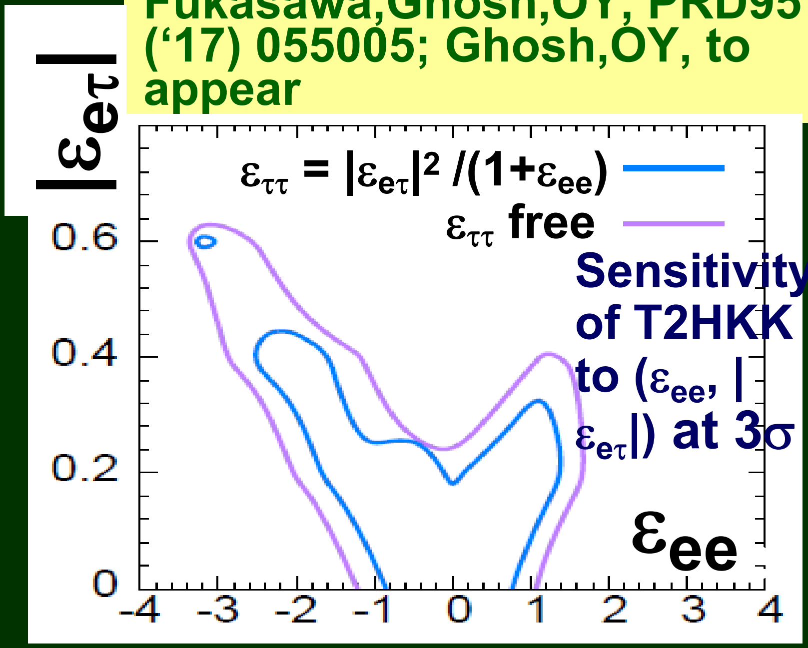


3. Sensitivity to NSI of propagation at T2HKK

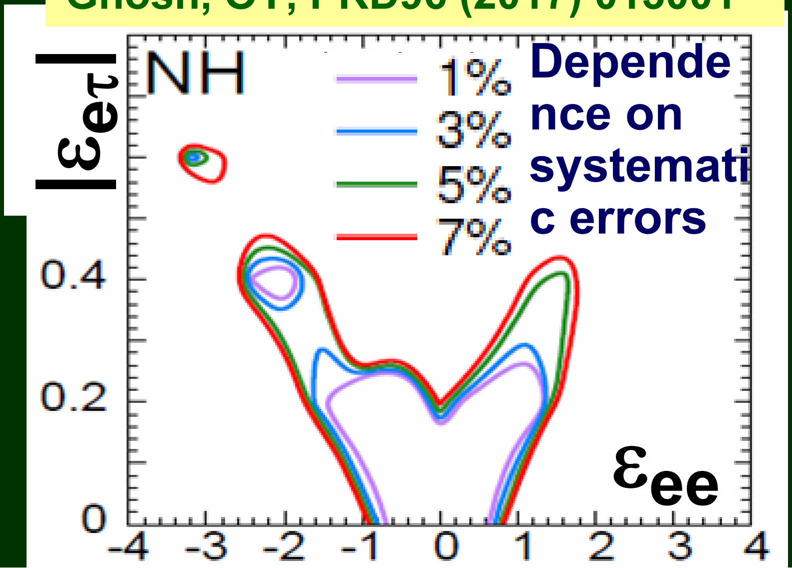
3.0 Motivation of our work

All the works on the sensitivity to NSI was expressed in terms of $\varepsilon_{\alpha\beta}$ typically in (ε_D, ε_N)-plane
-> Whether the LBL experiments have sensitivity to the region suggested by the solar tension is not clear.
-> Sensitivity given in (ε_D, ε_N)-plane is desired.

Fukasawa,Ghosh,OY, PRD95 ('17) 055005; Ghosh,OY, to appear



Ghosh, OY, PRD96 (2017) 013001



3.1 Outline of our Analysis

Strategy of our analysis:

● We assume $\varepsilon_{\alpha\beta}$(true) = 0 and minimize χ^2 (ε^f_D(test), ε^f_N(test)) by varying other $\varepsilon_{\alpha\beta}$(test).

We compare the sensitivities of

T2HKK, DUNE, HK(ν_{atm})

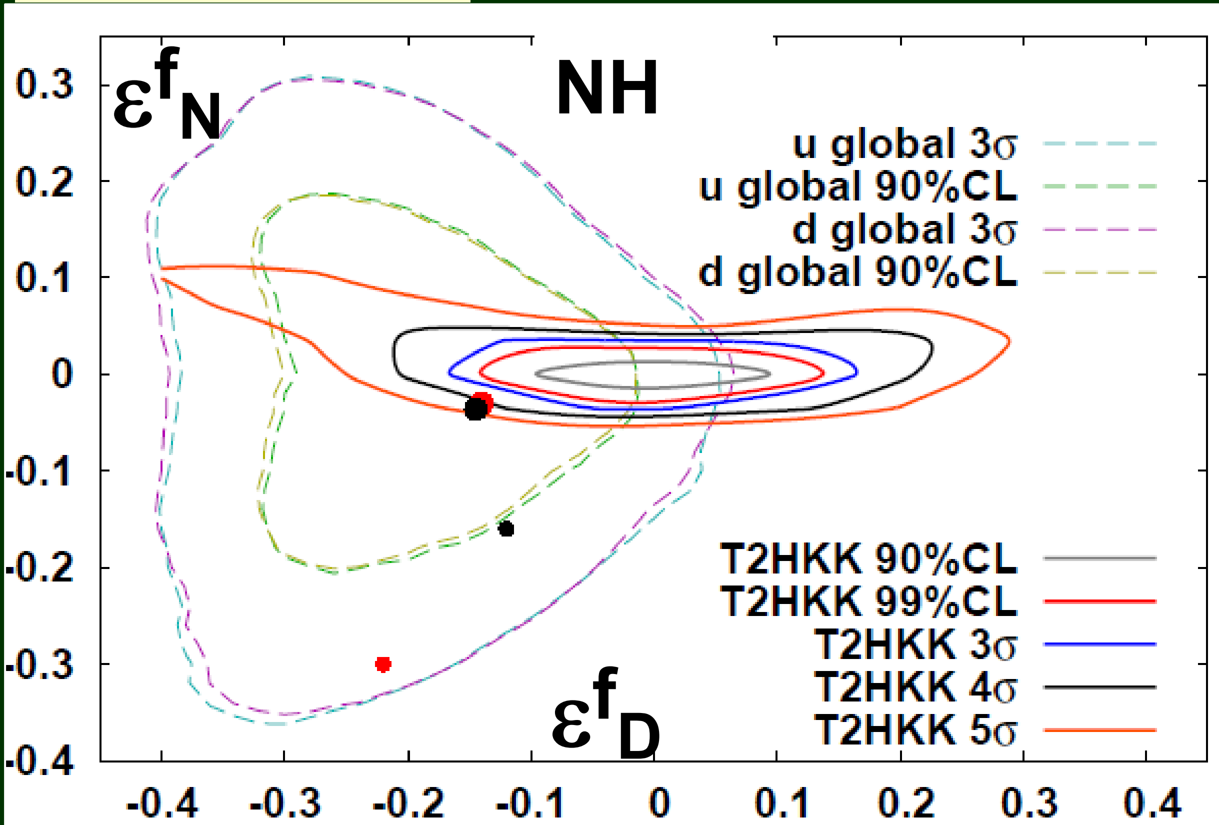
10km<L<13000km

L=1100km

L=1300km

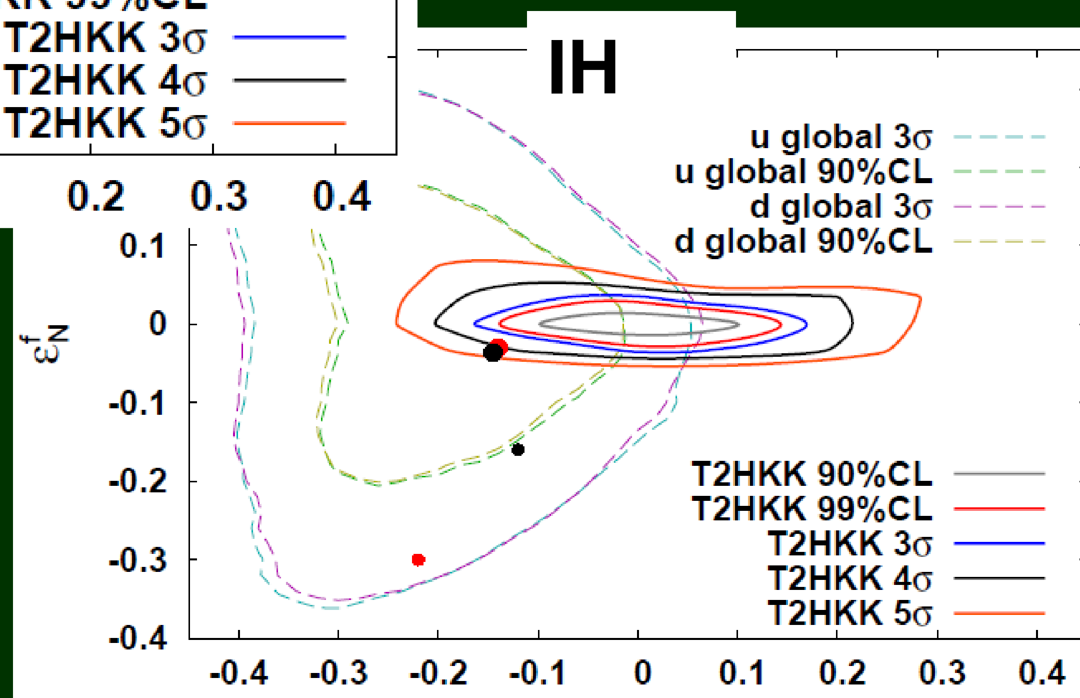
3.2 Results

Ghosh & OY, arXiv:1709.08264



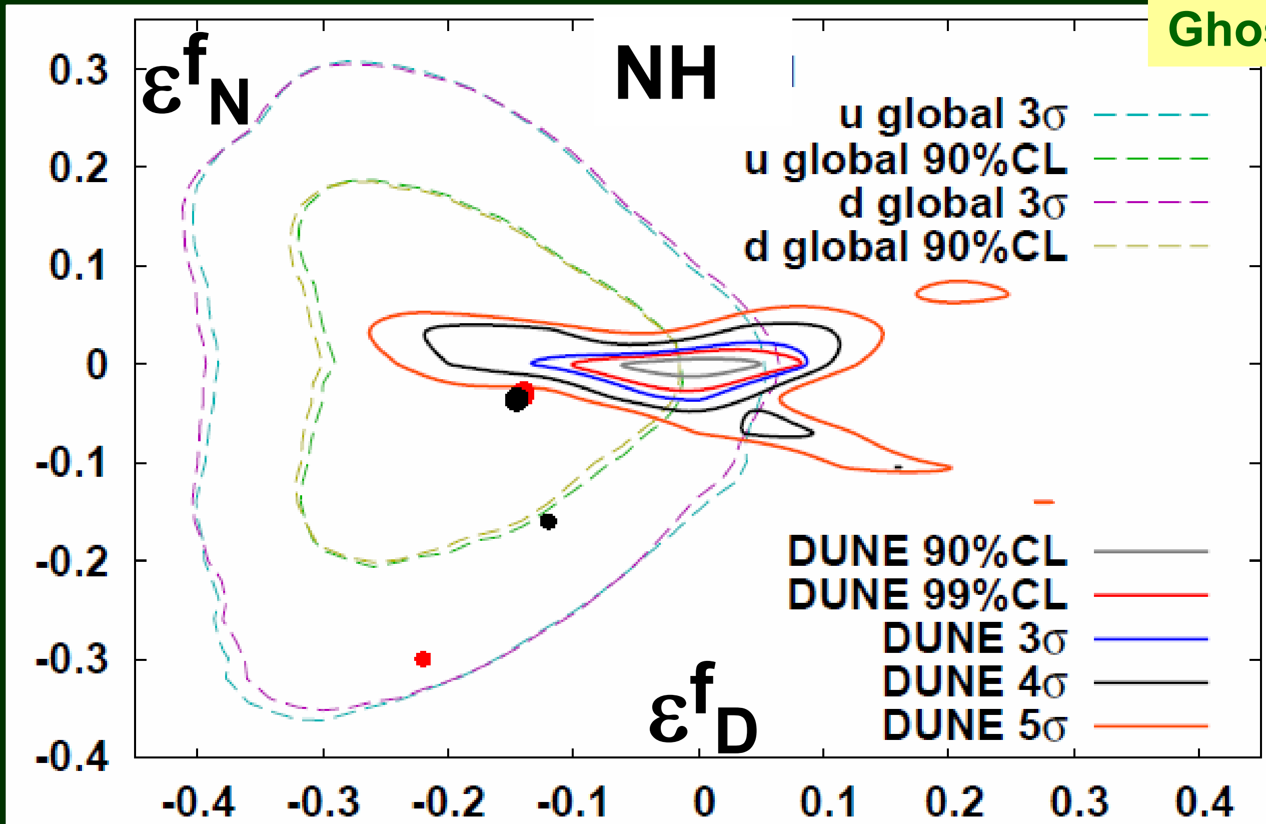
Excluded region by LBL is outside of the curve

δ(true) = -90°



T2HKK

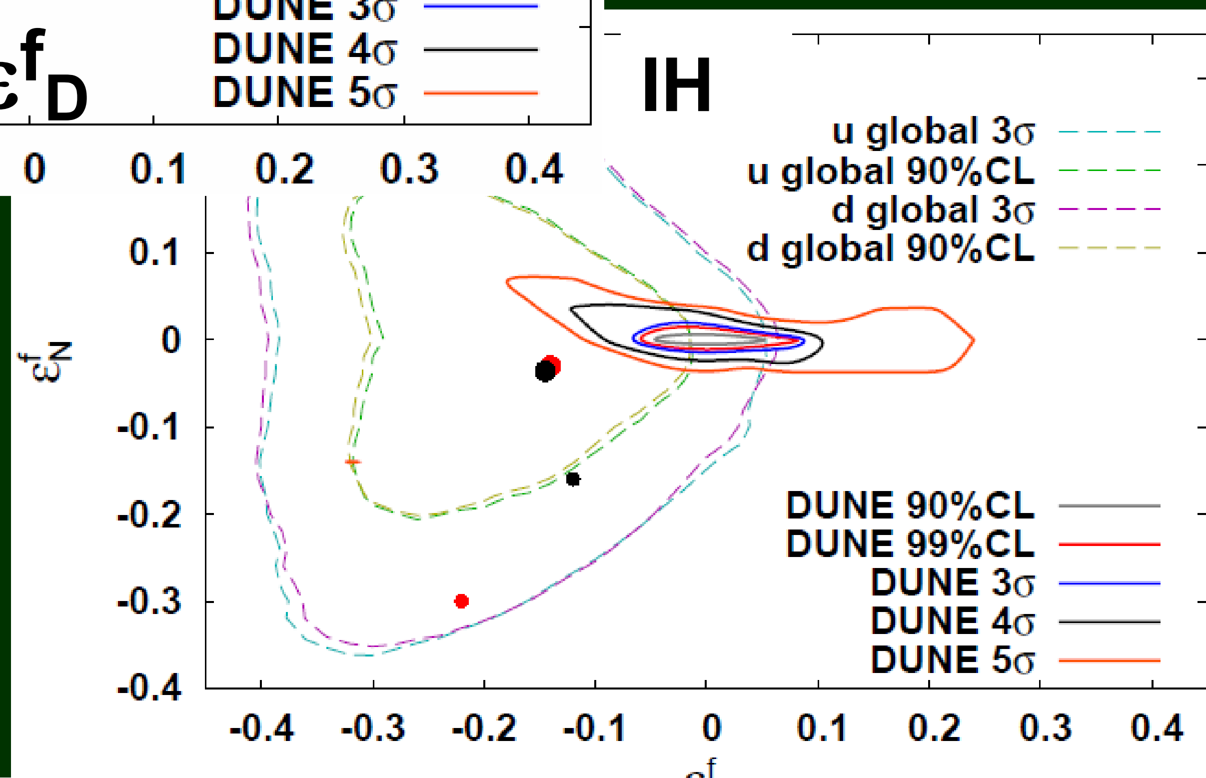
Ghosh & OY, arXiv:1709.08264



Sensitivity of DUNE is slightly better than T2HKK

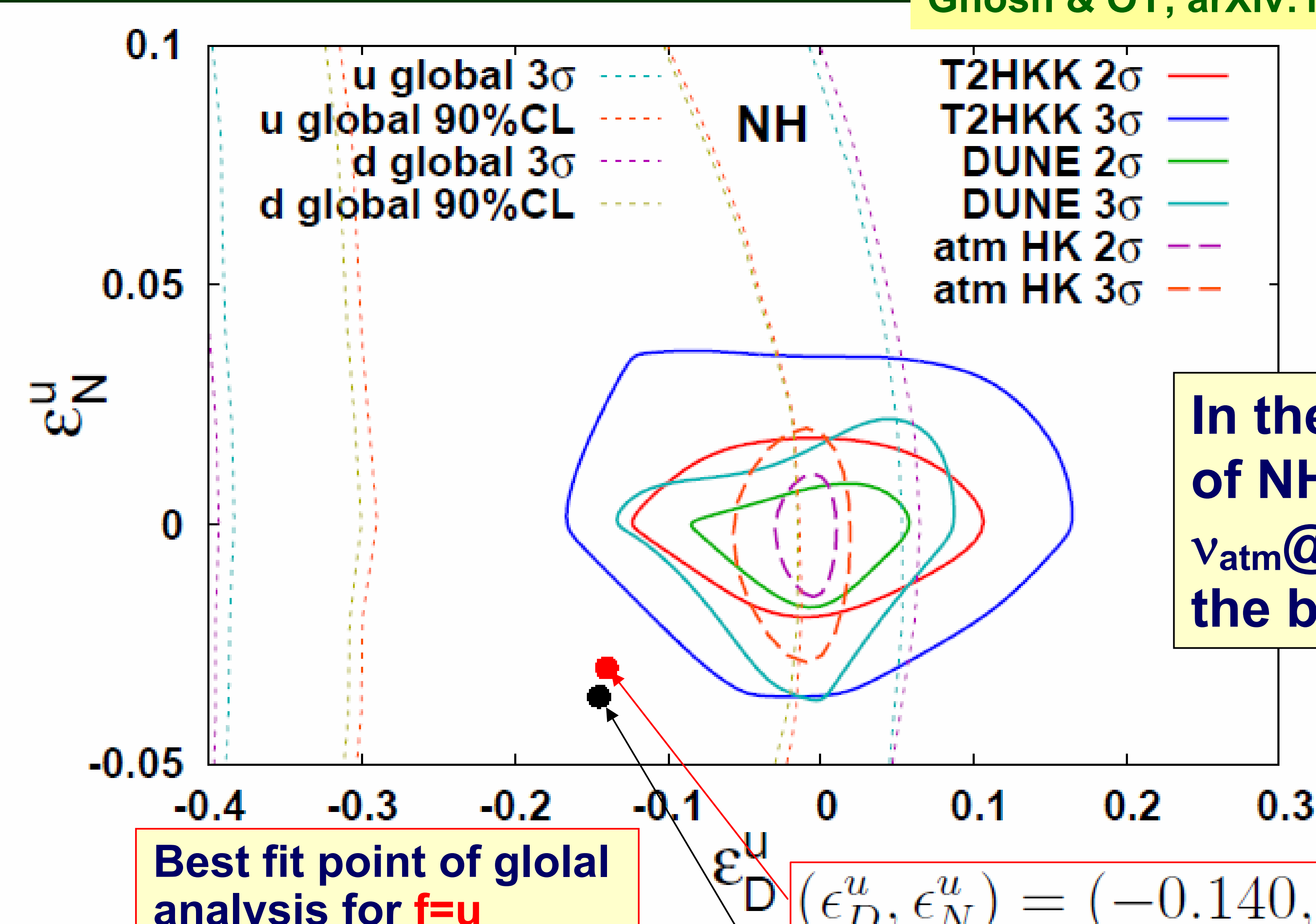
δ(true) = -90°

DUNE



●Comparison of sensitivity T2HKK, DUNE, ν_{atm}@HK

Ghosh & OY, arXiv:1709.08264



In the case of NH, ν_{atm}@HK is the best

Best fit point of glolal analysis for **f=u**

$$(\epsilon_D^u, \epsilon_N^u) = (-0.140, -0.030)$$

Best fit point of glolal analysis for f=d

$$(\epsilon_D^d, \epsilon_N^d) = (-0.145, -0.036)$$

4. Conclusions

- T2HKK and DUNE have sensitivity to NSI and they cover some of the allowed region in the (ϵ^f_D, ϵ^f_N)-plane suggested by the solar ν tension for δ(true) = -90°.
- Sensitivity of DUNE is slightly better than that of T2HKK because DUNE uses information of wide E_ν spectrum.

Backup slides

Observation of matter effect needs large L

ν oscillation in matter (in two flavor toy case)

$$P(\nu_\mu \to \nu_e) = \left(\frac{\Delta E}{\Delta \tilde{E}}\right)^2 \sin^2 2\theta \sin^2\left(\frac{\Delta \tilde{E} L}{2}\right)$$

$$\Delta E \equiv \Delta m^2 / 2E$$

$$\Delta \tilde{E} \equiv \left[(\Delta E \cos 2\theta - A)^2 + (\Delta E \sin 2\theta)^2\right]^{1/2}$$

$$A \equiv \sqrt{2} G_F n_e(x)$$

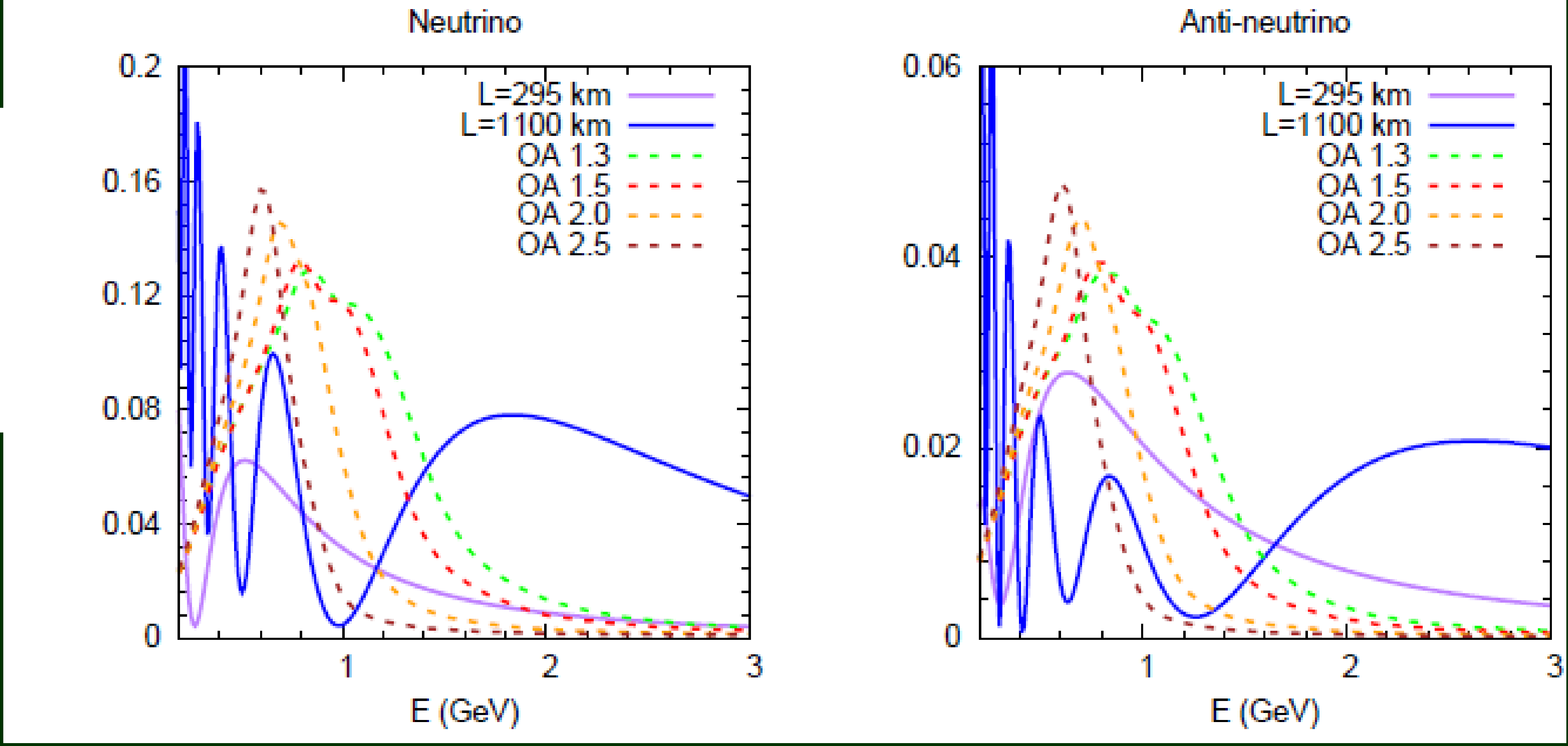
$$\tan 2\tilde{\theta} \equiv \frac{\Delta E \sin 2\theta}{\Delta E \cos 2\theta - A}$$

Matter effect becomes most conspicuous if $\Delta E \cos 2\theta = A$ is satisfied ($\tilde{\theta} = \pi/2$). In this case, the baseline length L has to be large:

$$\pi = \Delta \tilde{E} L = \Delta E \sin 2\theta L = A L \tan 2\theta$$

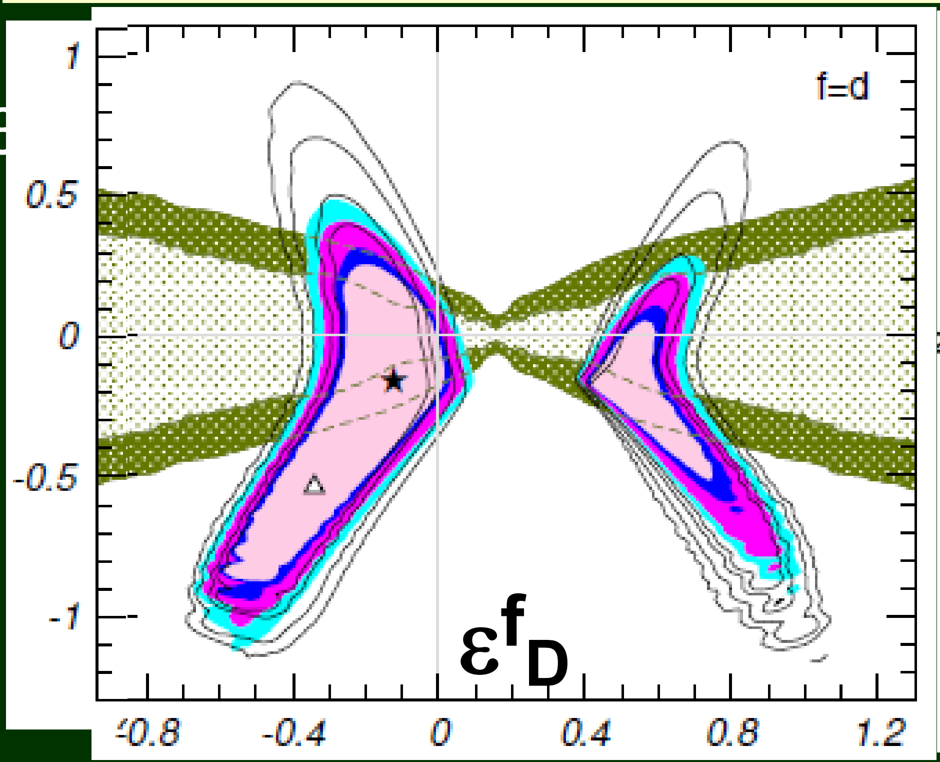
→L > π/A > O(1000km)

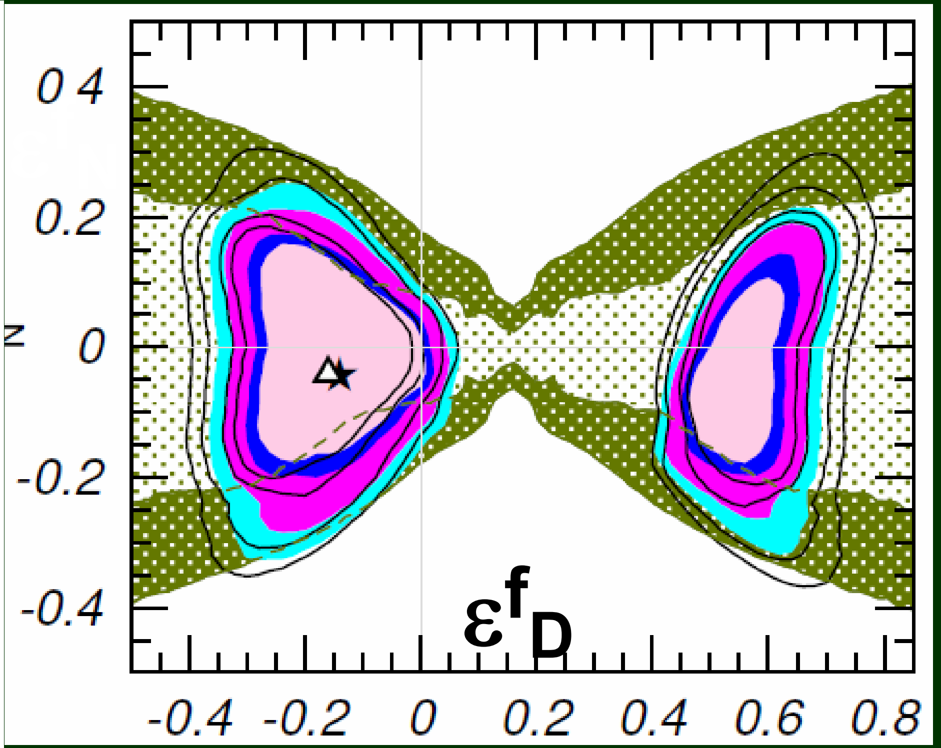
T2HKK:Appearance probability at L=1050km



Tension between solar ν & KamLAND can be solved by NSI

Gonzalez-Garcia, Maltoni, JHEP 1309 (2013) 152





Best fit value of solar-KL

$$(\epsilon^u_D, \epsilon^u_N) = (-0.22, -0.30)$$

$$(\epsilon^d_D, \epsilon^d_N) = (-0.12, -0.16)$$

Best fit value of global fit

$$(\epsilon^u_D, \epsilon^u_N) = (-0.140, -0.030)$$

$$(\epsilon^d_D, \epsilon^d_N) = (-0.145, -0.036)$$

Relation between $\varepsilon_{\alpha\beta}$ & $(\varepsilon_D, \varepsilon_N)$

We treat $\varepsilon^f_{\tau\tau}$, $|\varepsilon^f_{e\tau}|$, ε^f_{ee} as dependent variables:

$$|\epsilon^f_{e\tau}| = \frac{1}{c_{13}c_{23}\sin(\phi_{13}+\delta_{CP})}\left(-F\sin\delta_{CP} + G\cos\delta_{CP}\right)$$

$$\epsilon^f_{\tau\tau} = \frac{2}{s_{13}\sin 2\theta_{23}\ \sin(\phi_{13}+\delta_{CP})}\left(F\sin\phi_{13} + G\cos\phi_{13}\right)$$

$$\begin{aligned} F \equiv\ & \epsilon^f_N - c_{13}c_{23}\,|\epsilon^f_{e\mu}|\,\cos\phi_{12} \\ & - s_{13}|\epsilon^f_{\mu\tau}|\left\{s^2_{23}\cos(\phi_{23}-\delta_{CP}) - c^2_{23}\cos(\phi_{23}+\delta_{CP})\right\} \\ G \equiv\ & -c_{13}c_{23}\,|\epsilon^f_{e\mu}|\,\sin\phi_{12} \\ & - s_{13}|\epsilon^f_{\mu\tau}|\left\{s^2_{23}\sin(\phi_{23}-\delta_{CP}) + c^2_{23}\sin(\phi_{23}+\delta_{CP})\right\} \end{aligned}$$

↓

ϕ_{12} =arg($\varepsilon^f_{e\mu}$), ϕ_{13} =arg($\varepsilon^f_{e\tau}$), ϕ_{23} =arg($\varepsilon^f_{\mu\tau}$)

$$\epsilon^f_{ee} = \frac{2}{c^2_{13}} \left\{ \frac{s_{23}}{2} \sin 2\theta_{13} |\epsilon^f_{e\mu}| \cos(\delta_{\rm CP} + \phi_{12}) \right.$$

$$+ \frac{c_{23}}{2} \sin 2\theta_{13} |\epsilon^f_{e\tau}| \cos(\delta_{\rm CP} + \phi_{13})$$

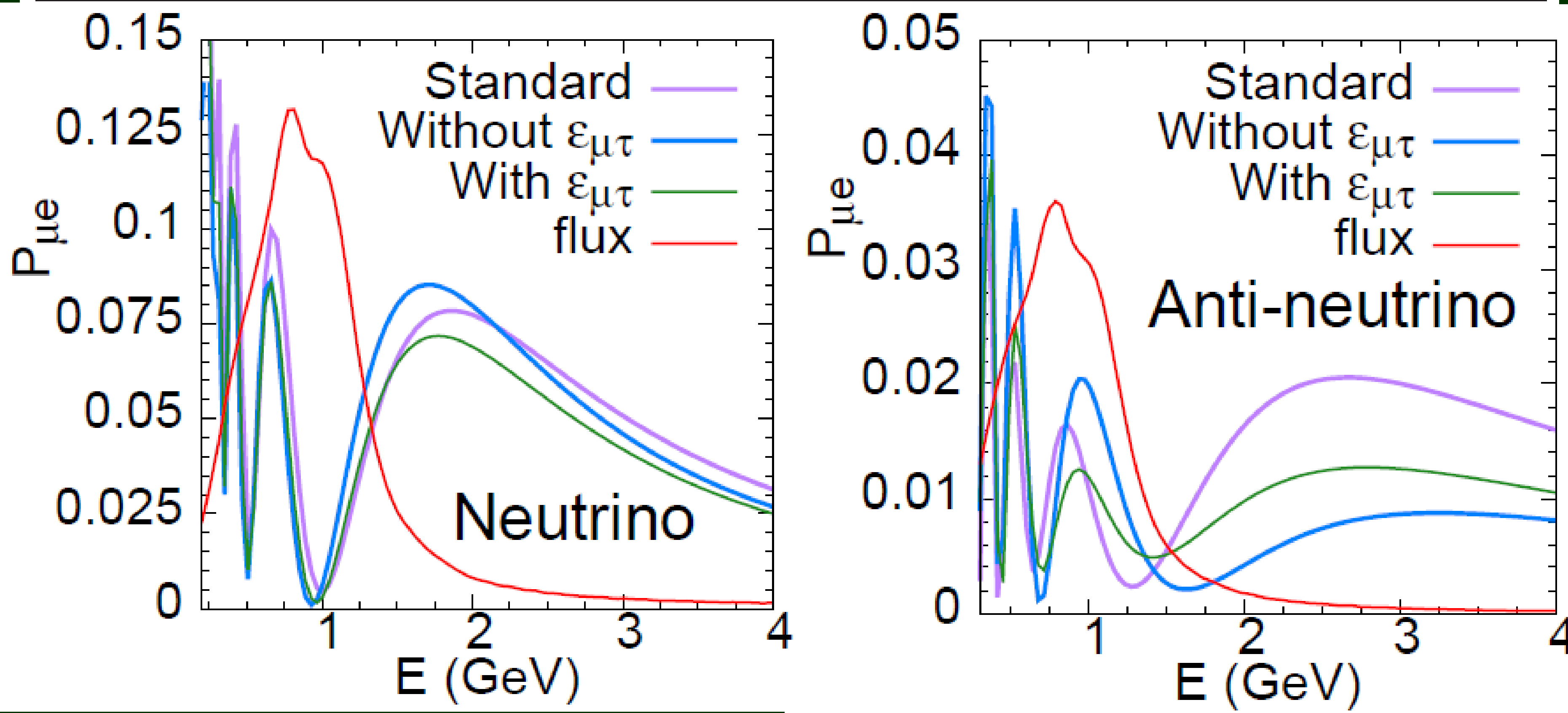
$$- (1 + s^2_{13})\, c_{23} s_{23} |\epsilon^f_{\mu\tau}| \cos(\phi_{23})$$

$$\left. - \epsilon^f_D + \frac{s^2_{23} - s^2_{13} c^2_{23}}{2} \epsilon^f_{\tau\tau} \right\}$$

$\phi_{12} = \arg(\varepsilon^f_{e\mu})$, $\phi_{13} = \arg(\varepsilon^f_{e\tau})$, $\phi_{23} = \arg(\varepsilon^f_{\mu\tau})$

In principle we could take into account $\varepsilon^f_{e\mu}$, but contribution from $\varepsilon^f_{e\mu}$ turns out to be small, so we put $\varepsilon^f_{e\mu}$ =0 for simplicity

ϵ_{ee}	$\|\epsilon_{e\tau}\|$	$\epsilon_{\tau\tau}$	δ_{CP}	θ_{23}	$\arg(\epsilon_{e\tau})$	$\|\epsilon_{\mu\tau}\|$	$\arg(\epsilon_{\mu\tau})$	$\|\epsilon_{e\mu}\|$	$\arg(\epsilon_{e\mu})$	χ^2
0.846	0.123	-0.021	-90	47	0	0	0	0	0	25.46
1.128	0.108	0.511	-90	45	30	0.15	90	0	0	17.54
0.917	0.146	0.114	-90	47	0	0	0	0.03	30	24.61



-> Independent variables to be marginalized over:
Δm^2_{32}, θ_{23}, δ, $|\varepsilon^f_{\mu\tau}|$, ϕ_{13}

$$\chi^2 = \min_{\xi_k,\ \text{osc. param}} \left(\chi^2_{\text{stat}} + \sum_k \xi_k^2 + \chi^2_{\text{prior}} \right)$$

$$\chi^2_{\text{stat}} = 2 \sum_i \left\{ \tilde{N}_i^{\text{test}} - N_i^{\text{true}} - N_i^{\text{true}} \log \left(\frac{\tilde{N}_i^{\text{test}}}{N_i^{\text{true}}} \right) \right\}$$

Pull variables for systematic errors

$$\tilde{N}_i^{\text{test}} \equiv \left(1 + \sum_k c_i^k \xi_k \right) N_i^{\text{test}}$$

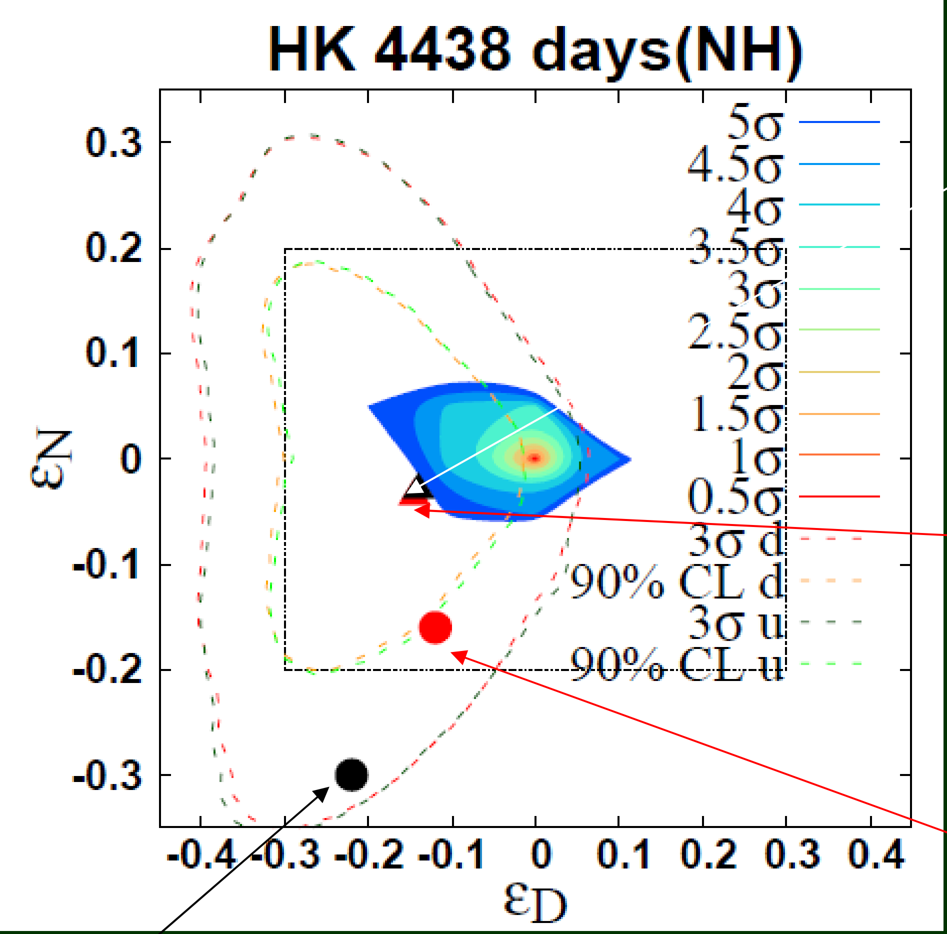
$$\chi^2_{\text{prior}} = 2.7 \left(\frac{|\epsilon_{e\mu}|}{0.15} \right)^2 + 2.7 \left(\frac{|\epsilon_{\mu\tau}|}{0.15} \right)^2$$

$$|\epsilon^f_{e\mu}| < 0.05, \quad |\epsilon^f_{\mu\tau}| < 0.05$$

$$\epsilon_{\alpha\beta} = 3\,\epsilon^f_{\alpha\beta}$$

Sensitivity of ν_{atm} at HK : Real ε_N

OY@nufact2016



$$(\epsilon^u_D, \epsilon^u_N) = (-0.140, -0.030)$$

Best fit point of glolal analysis for f=u: significance:5σ

$$(\epsilon^d_D, \epsilon^d_N) = (-0.145, -0.036)$$

Best fit point of glolal analysis for f=d: significance:5σ

$$(\epsilon^d_D, \epsilon^d_N) = (-0.12, -0.16)$$

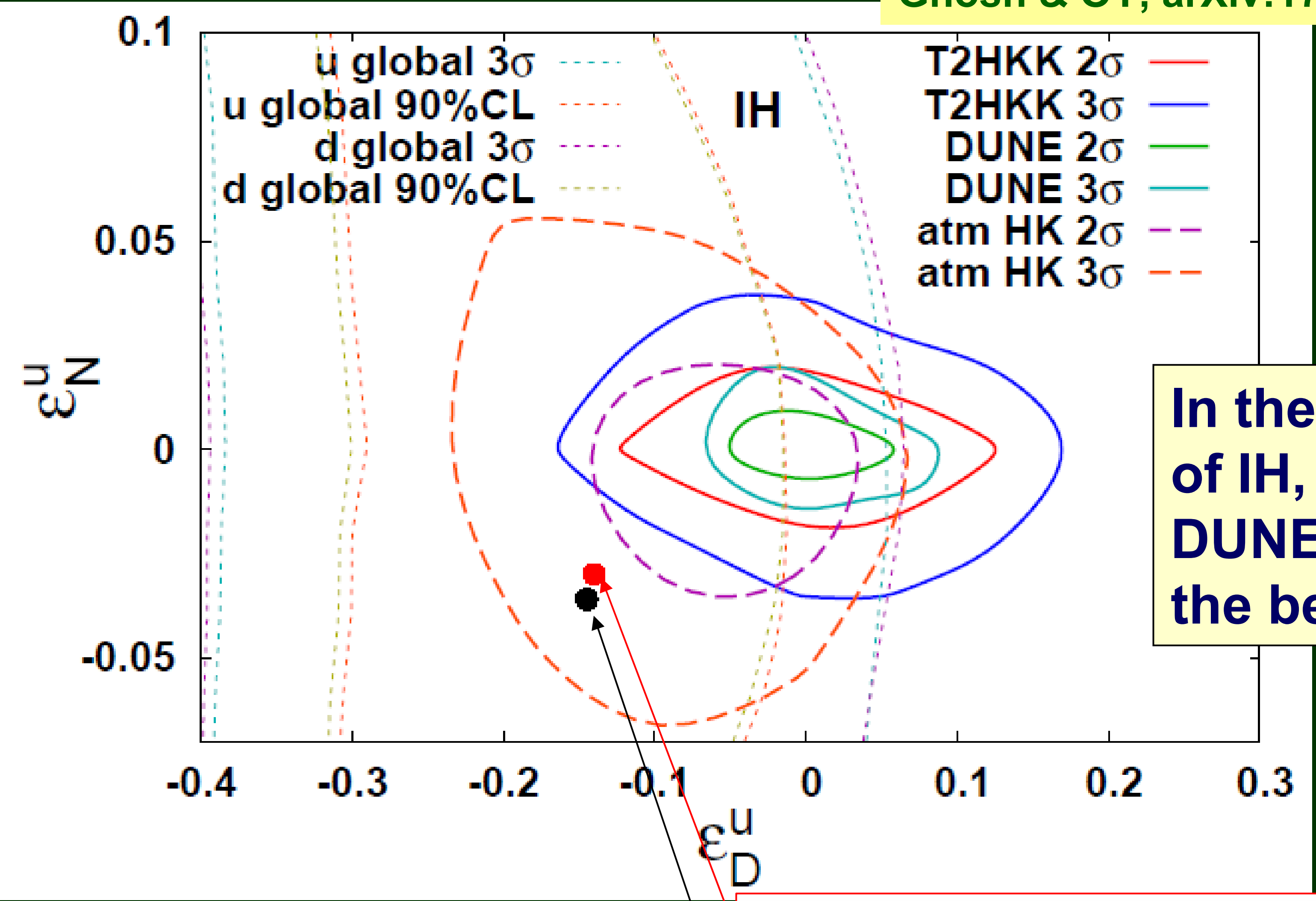
Best fit point of solar & KamLAND for f=d: significance:11σ

$$(\epsilon^u_D, \epsilon^u_N) = (-0.22, -0.30)$$

Best fit point of solar & KamLAND for f=u: significance:38σ

●Comparison of sensitivity T2HKK, DUNE, ν_{atm}@HK

Ghosh & OY, arXiv:1709.08264

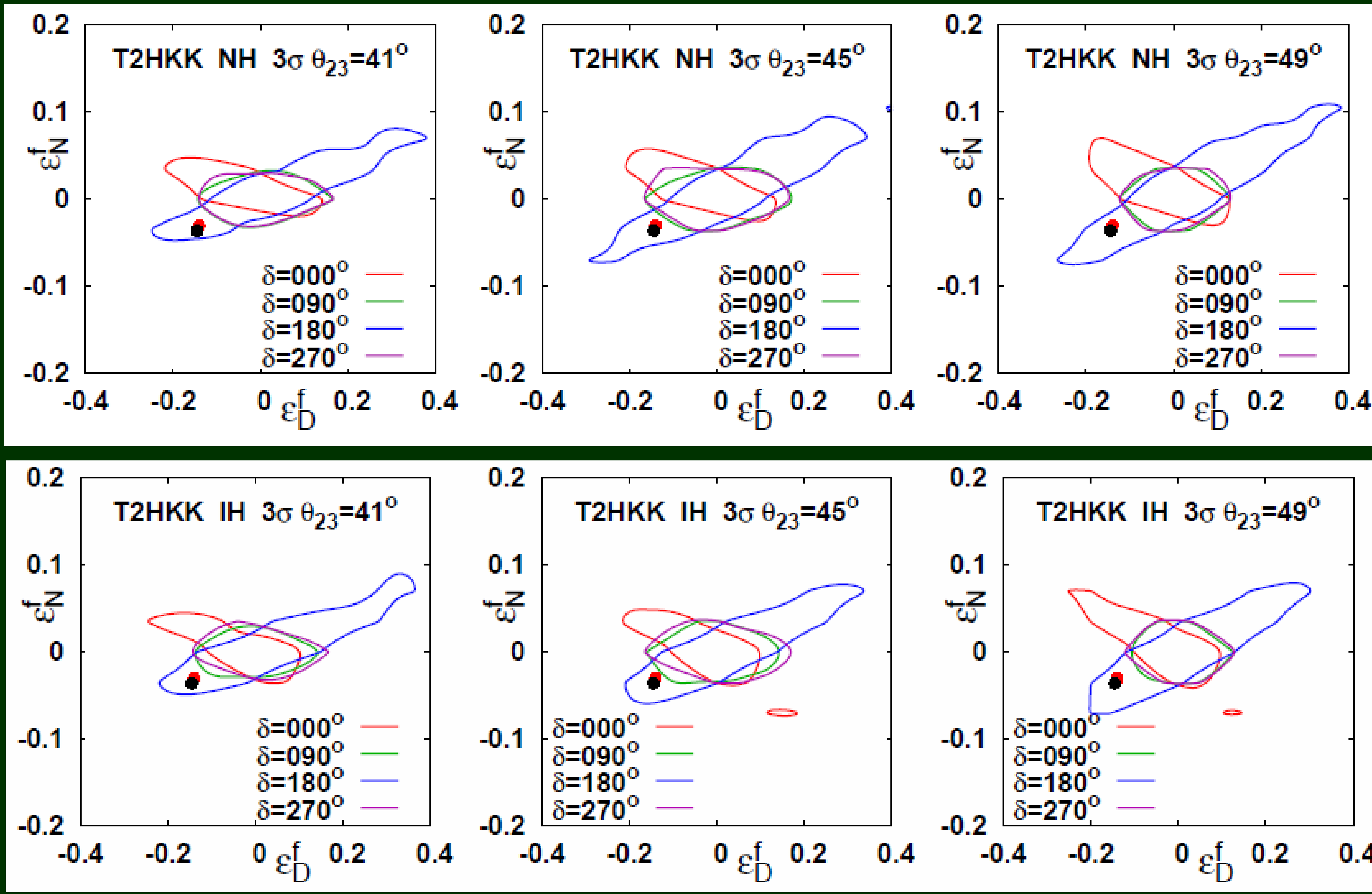


In the case of IH, DUNE is the best

$$(\epsilon^u_D, \epsilon^u_N) = (-0.140, -0.030)$$

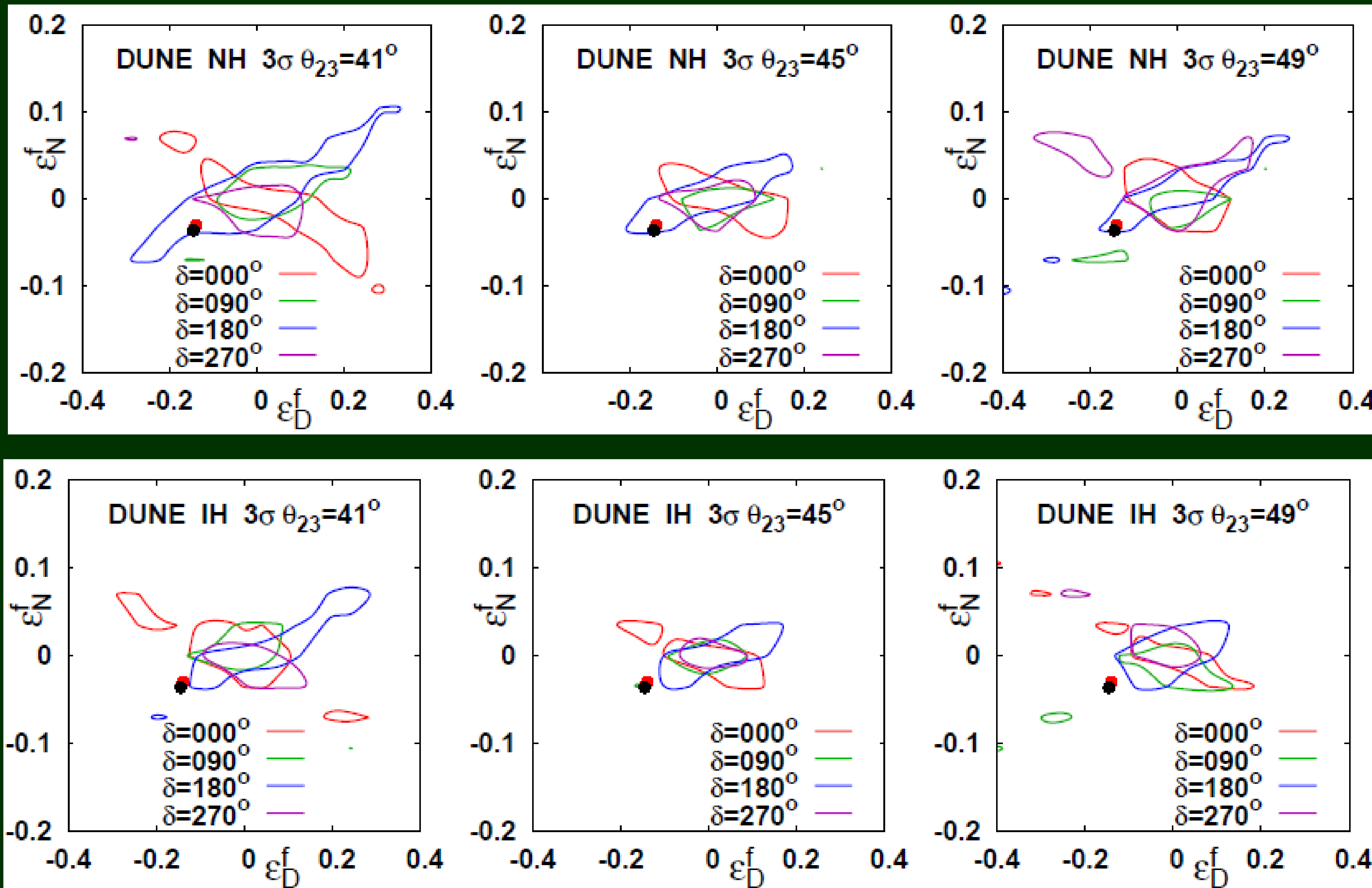
$$(\epsilon^d_D, \epsilon^d_N) = (-0.145, -0.036)$$

●Dependence of T2HKK on θ_{23}(true) & δ(true)



Ghosh & OY, arXiv:1709.08264

● Dependence of DUNE on θ_{23}(true) & δ(true)



Ghosh & OY, arXiv:1709.08264

- **Some model predicts large NSI** **(new gauge boson mass is of O(10MeV) and SU(2) invariance is broken):**

Farzan, PLB748 ('15) 311;
Farzan-Shoemaker, JHEP,1607 ('16)033;
Farzan-Heeck, PRD94 ('16) 053010.